Neutrino Interaction Physics

Lecture 1: Introduction of neutrino interactions
1. Overview
2. Neutrino lepton scattering
3. Neutrino quark scattering (DIS)
4. Neutrino nucleus reactions

Lecture 2: Neutrino interactions for long baseline oscillation experiments
1. Overview
2. CCQE interaction
3. Baryonic resonances
4. Shallow inelastic scattering (SIS)

Teppei Katori
King's College London
Nov. 1, 2022

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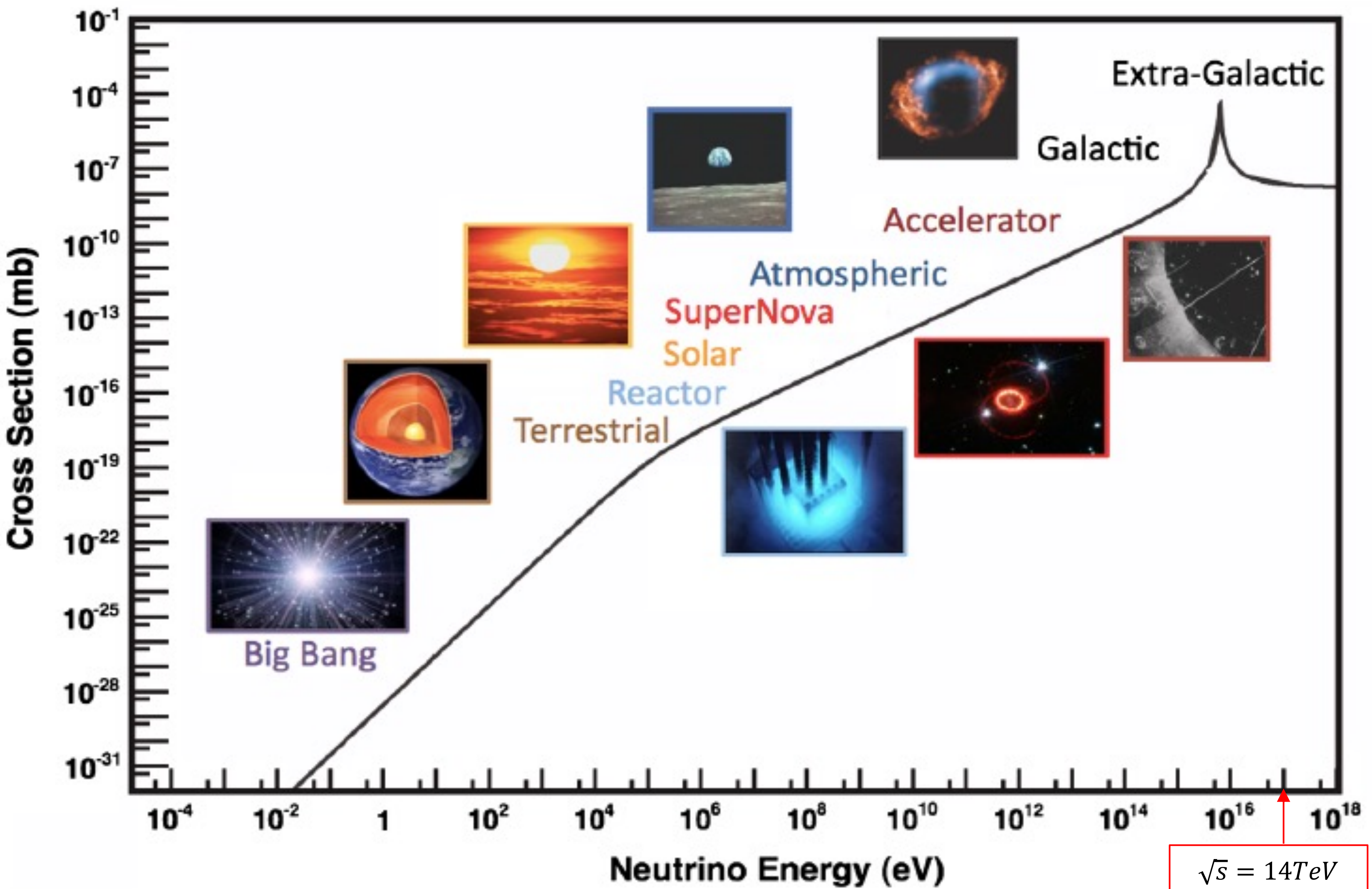
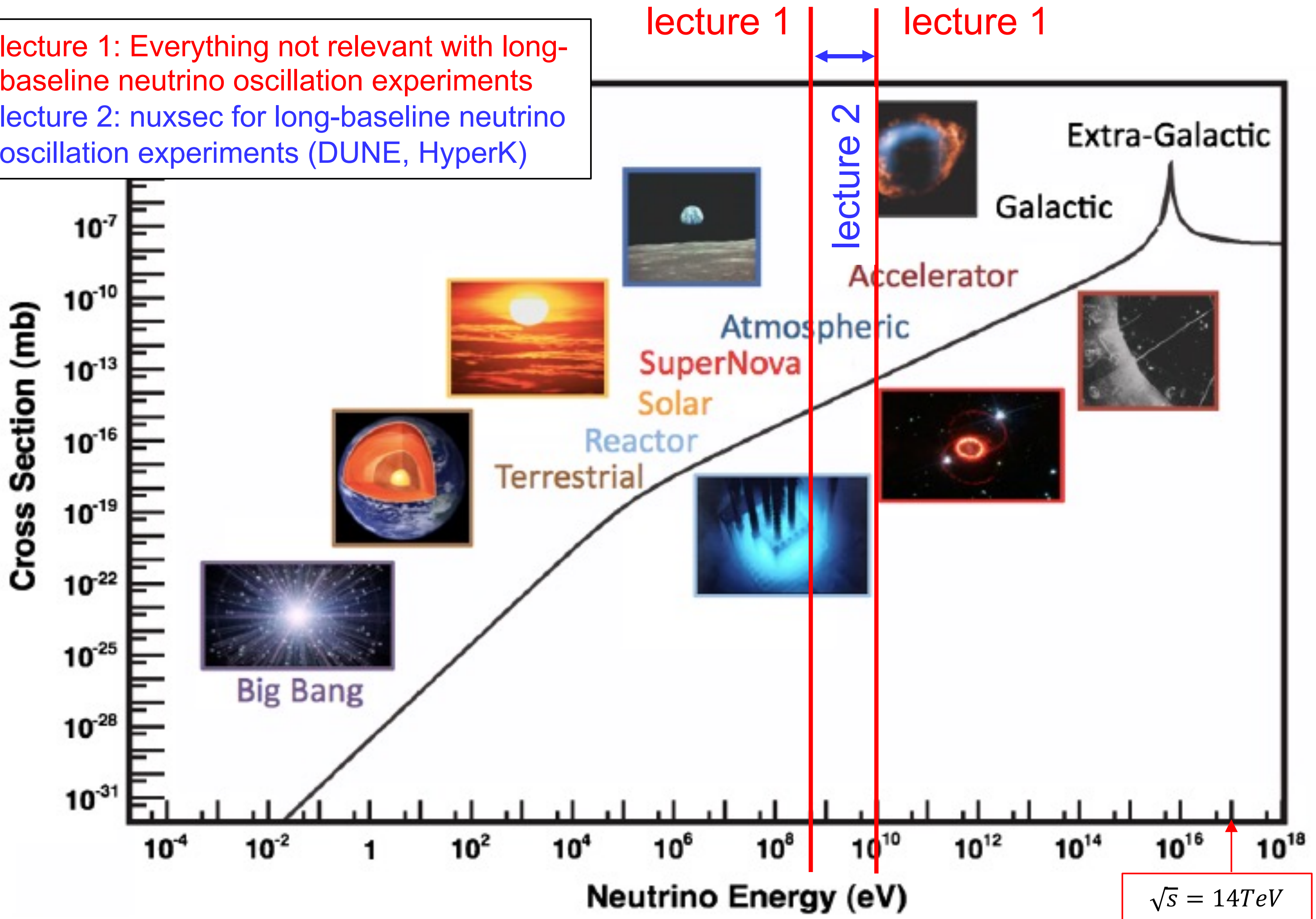
1. From eV to EeV: Neutrino cross sections across energy scales



Formaggio and Zeller, Rev.Mod.Phys.,84 (2012) 1307

1. From eV to EeV: Neutrino cross sections across energy scales

lecture 1: Everything not relevant with long-baseline neutrino oscillation experiments
lecture 2: nuxsec for long-baseline neutrino oscillation experiments (DUNE, HyperK)

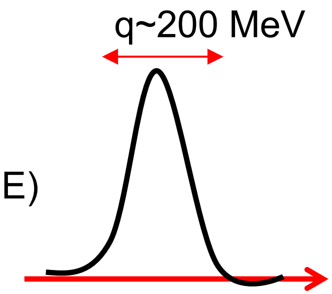


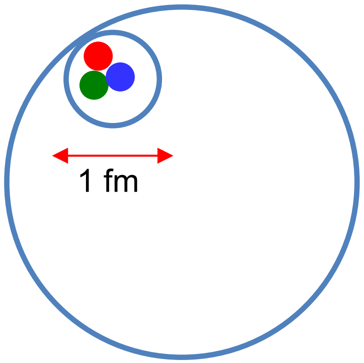
1. Scattering measurements

Size of wave packet ~ momentum transfer (~energy)

$\hbar c = 197\ MeV \cdot fm \rightarrow$ 200 MeV ~ 1 fm (size of nucleon)

~1 GeV neutrino beam
(T2K, NOvA, HyperK, DUNE)





1. Scattering measurements

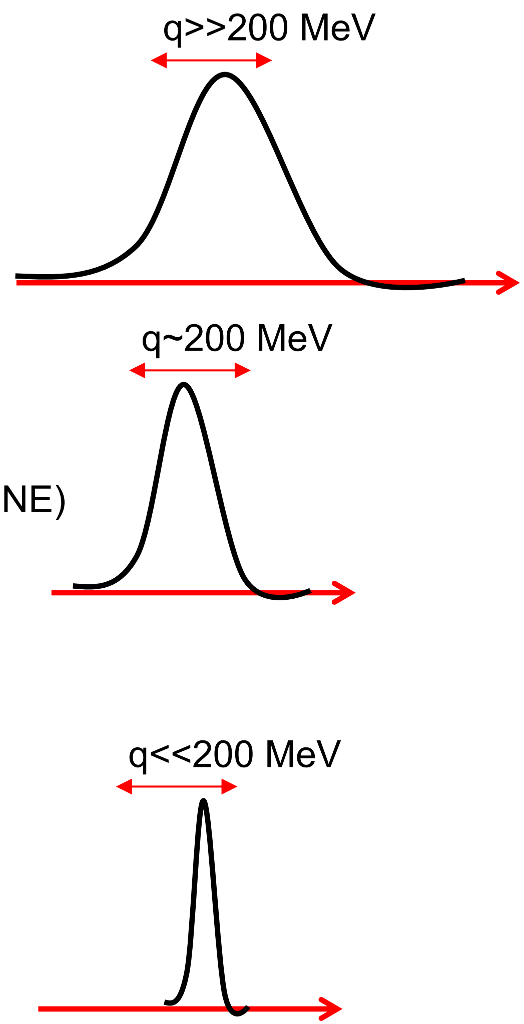
Size of wave packet ~ momentum transfer (~energy)

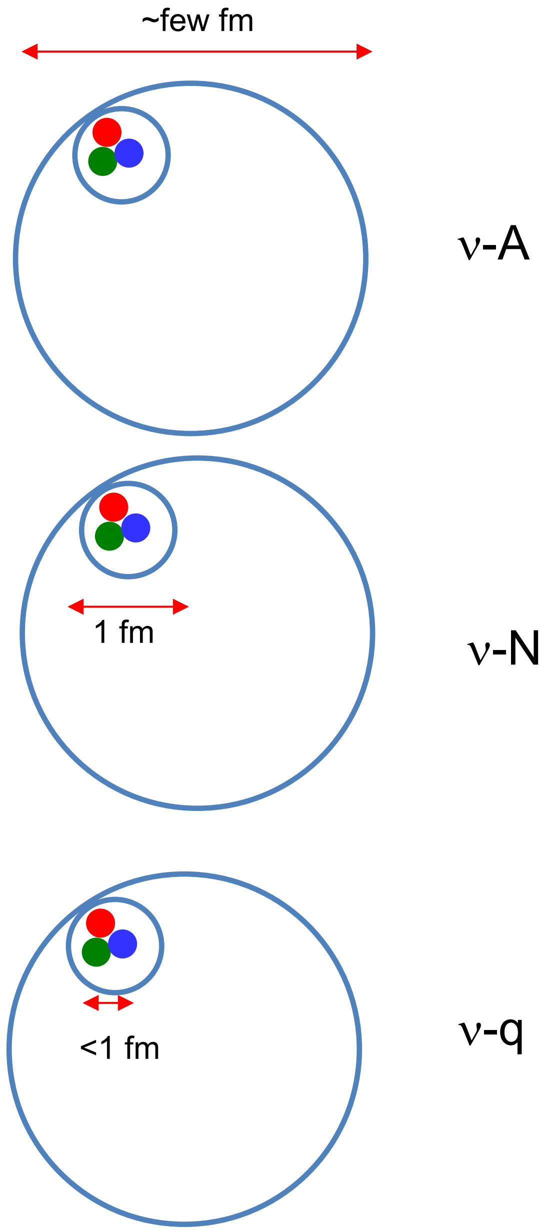
$\hbar c = 197\ MeV \cdot fm \rightarrow$ 200 MeV ~ 1 fm (size of nucleon)

<<1 GeV neutrino beam
(solar neutrinos, etc)

~1 GeV neutrino beam
(T2K, NOvA, HyperK, DUNE)

>>1 GeV neutrino beam
(IceCube, FASERnu)





1. Scattering measurements

Size of wave packet ~ momentum transfer (~energy)

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Lecture 1: Introduction of neutrino interactions
1. Overview
2. Neutrino lepton scattering (Standard Model)
3. Neutrino quark scattering (ν-q scattering)
4. Neutrino nucleus reactions (ν-A scattering)

Lecture 2: Neutrino interactions for long baseline oscillation experiments (ν-N scattering)
1. Overview
2. CCQE interaction
3. Baryonic resonances
4. Shallow inelastic scattering

2. Neutrino-electron scattering

Neutrino – electron differential cross section

$$\frac{d\sigma}{dT} = \frac{2G_F^2 m_e}{\pi}\left[c_L^2 + c_R^2\left(\frac{E-T}{E}\right)^2 - C_L C_R \frac{m_e T}{E^2}\right]$$

T=recoil electron kinetic energy

E=neutrino energy

	C_L	C_R
ν_e - e^-	$½+\sin^2\theta_W$	$\sin^2\theta_W$
$\overline{\nu_e}$ - e^-	$\sin^2\theta_W$	$½+\sin^2\theta_W$
ν_μ - e^-	$-½+\sin^2\theta_W$	$\sin^2\theta_W$
$\overline{\nu_\mu}$ - e^-	$\sin^2\theta_W$	$-½+\sin^2\theta_W$

2. MINERvA neutrino flux tuning

Neutrino – electron differential cross section

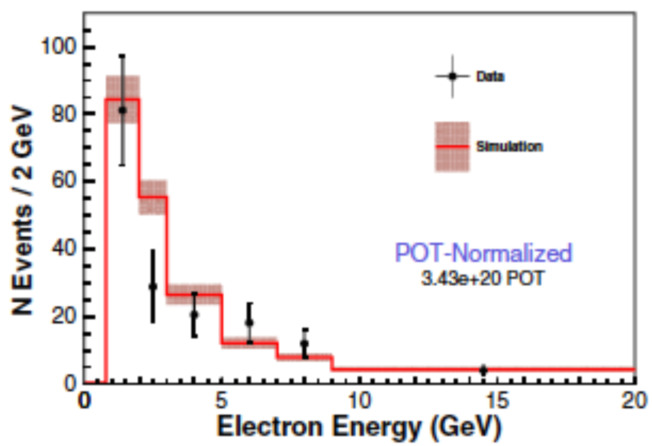
$$\frac{d\sigma}{dT} = \frac{2G_F^2 m_e}{\pi}\left[c_L^2 + c_R^2\left(\frac{E-T}{E}\right)^2 - C_L C_R \frac{m_e T}{E^2}\right]$$

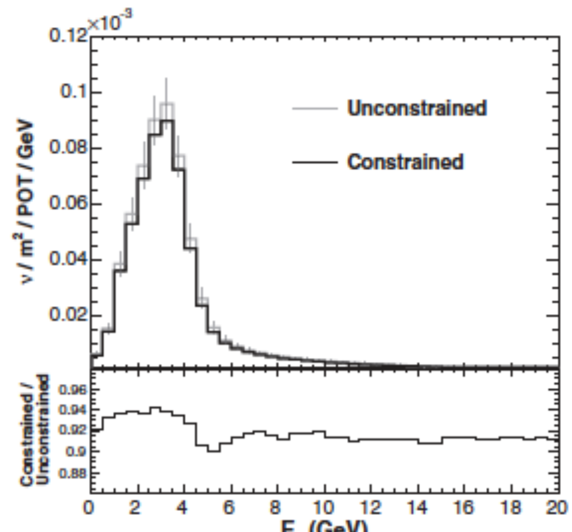
T=recoil electron kinetic energy

E=neutrino energy

$$\#events = \left(\int flux \otimes cross\ section \otimes efficiency \otimes target\ number\right) \times exposure$$

By assuming detector efficiency and cross-section are known, you can measure neutrino flux





TEXONO, PRD81(2010)072001

2. Neutrino magnetic moment

Neutrino – electron differential cross section

$$\frac{d\sigma}{dT} = \frac{2G_F^2 m_e}{\pi}\left[c_L^2 + c_R^2\left(\frac{E-T}{E}\right)^2 - C_L C_R \frac{m_e T}{E^2}\right]$$

T=recoil electron kinetic energy

E=neutrino energy

Neutrino – electron differential cross section with neutrino magnetic moment

$$\frac{d\sigma}{dT} = \underbrace{\frac{2G_F^2 m_e}{\pi}\left[c_L^2 + c_R^2\left(\frac{E-T}{E}\right)^2 - C_L C_R \frac{m_e T}{E^2}\right]}_{\text{SM}} + \underbrace{\frac{\pi\alpha\mu_\nu^2}{m_e^2}\left(\frac{1}{T} - \frac{1}{E}\right)}_{\text{large neutrino magnetic moment (BSM)}}$$

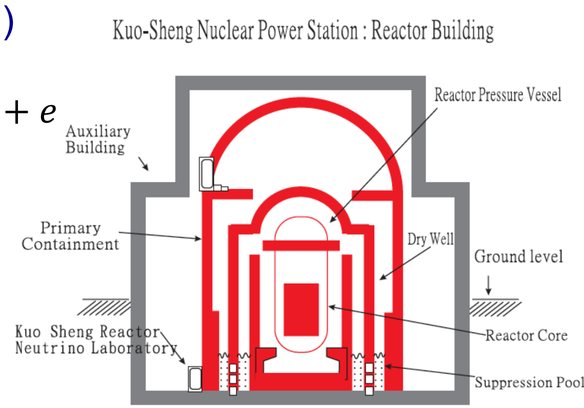
SM

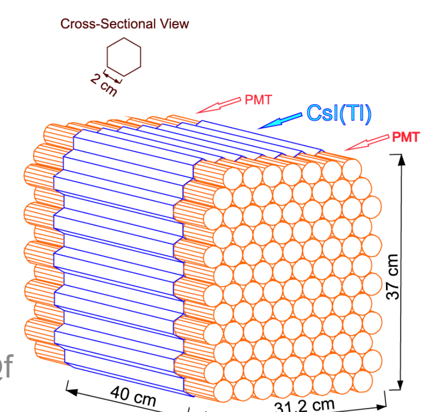
large neutrino magnetic moment (BSM)

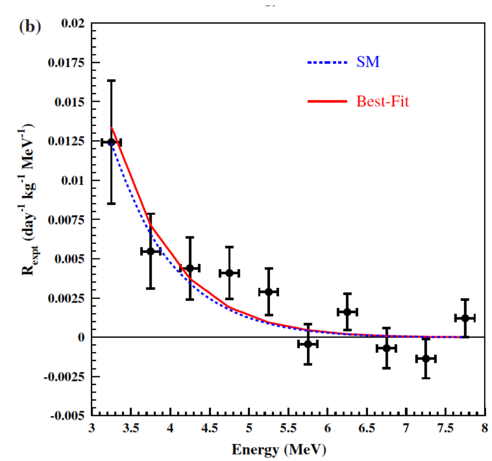
Lepton-only process (pure Standard Model) is often used to test new physics

TEXONO (Taiwan)

- reactor neutrino
- ES: $\bar{\nu}_e + e \rightarrow \bar{\nu}_e + e$
- CsI (Tl) crystal array (187kg)







3. Neutrino-DIS cross section

Neutrino – single d-quark cross section

$$\frac{d\sigma}{dy}(\nu d \to \mu u) = \frac{G_F^2 xs}{\pi}$$

Neutrino – d-quark cross section

$$\frac{d\sigma}{dy}(\nu d \to \mu u) = \int_0^1 \frac{G_F^2 xs}{\pi} d(x) dx$$

Neutrino-nucleon DIS cross section

$$\frac{d\sigma}{dy}(\nu N \to \mu X) = \int_0^1 \frac{G_F^2 xs}{\pi} [(d(x) + s(x) \ldots) + [\bar{u}(x) + \bar{c}(x) \ldots](1-y)^2] dx$$

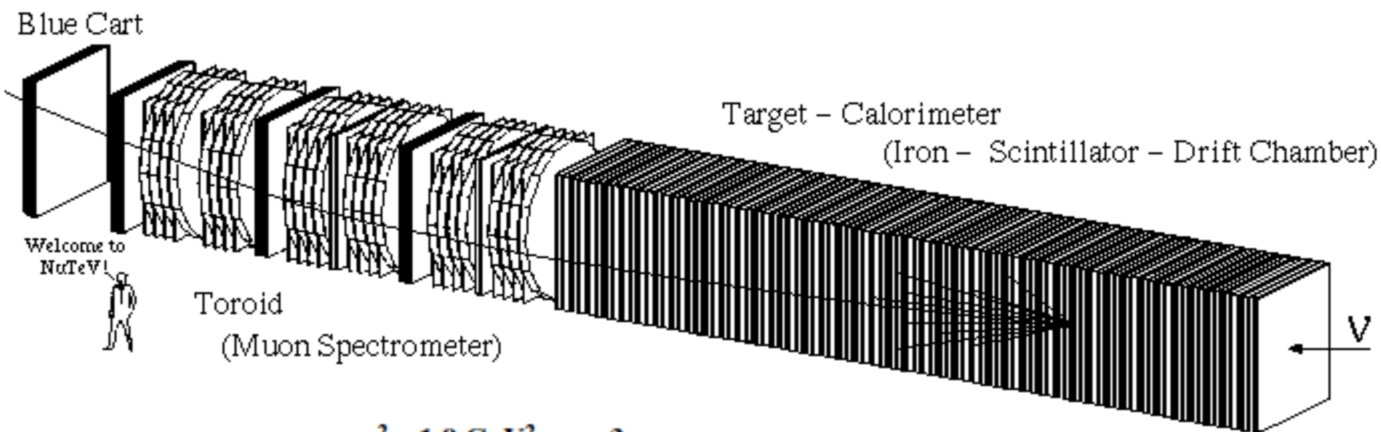
Neutrino-nucleus DIS cross section with isoscalar assumption

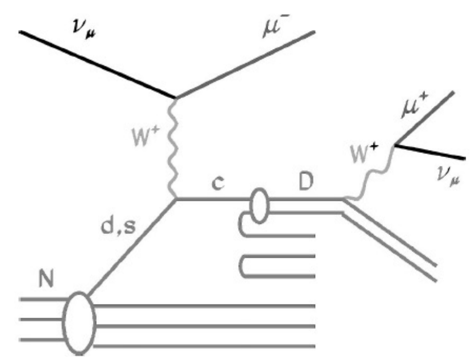
$$\frac{d\sigma}{dy}(\nu A \to \mu X) = A \int_0^1 \frac{G_F^2 xs}{\pi} [Q(x) + \bar{Q}(x)(1-y)^2] dx$$

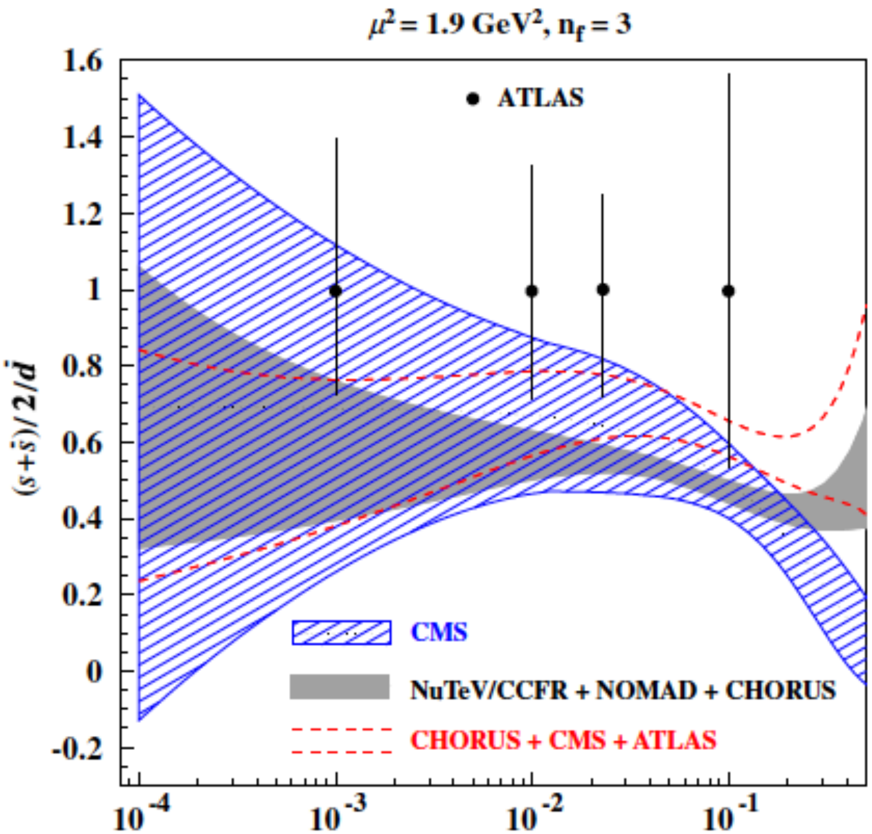
$$u^p(x) + u^n(x) = d^n(x) + d^p(x) = u(x) + d(x) \equiv Q(x)$$
$$\bar{u}^p(x) + \bar{u}^n(x) = \bar{u}^n(x) + \bar{u}^p(x) = \bar{u}(x) + \bar{d}(x) \equiv \bar{Q}(x)$$

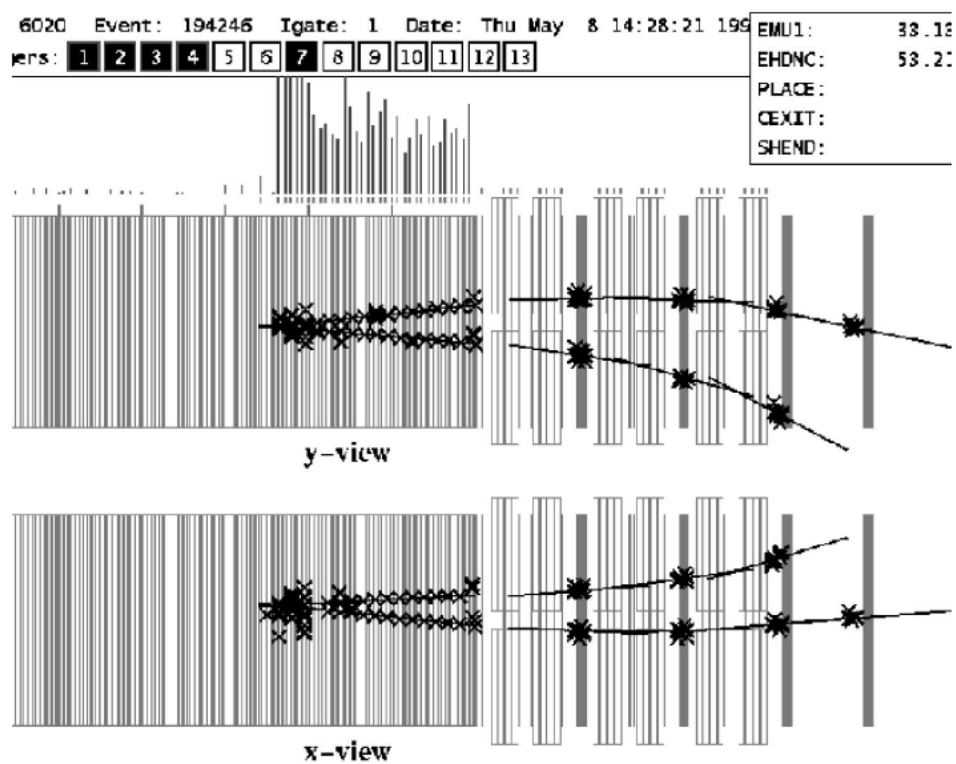
NuTeV, PRD64(2001)112006
Alekhin et al.,PRD91(2015)094002

3. Di-muon production

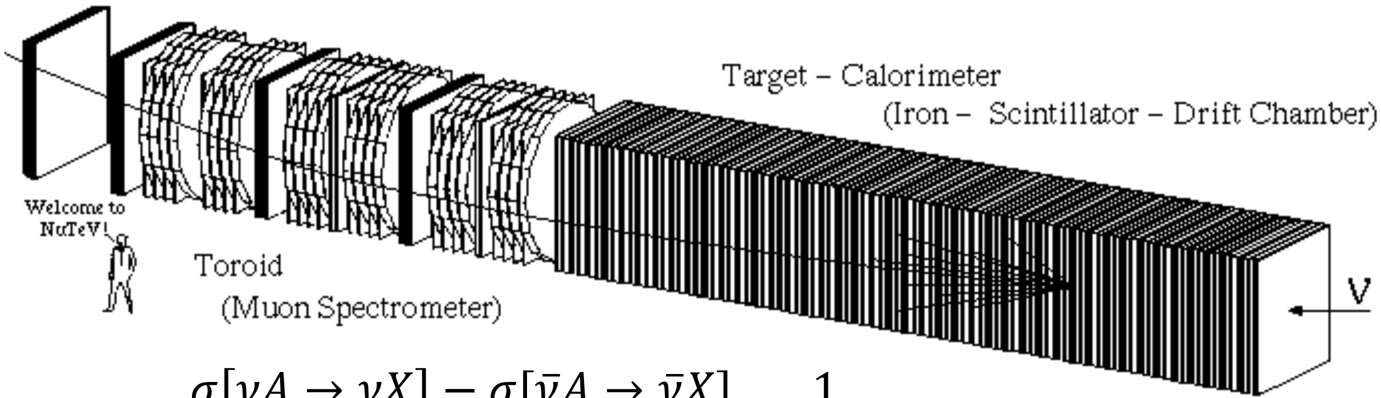








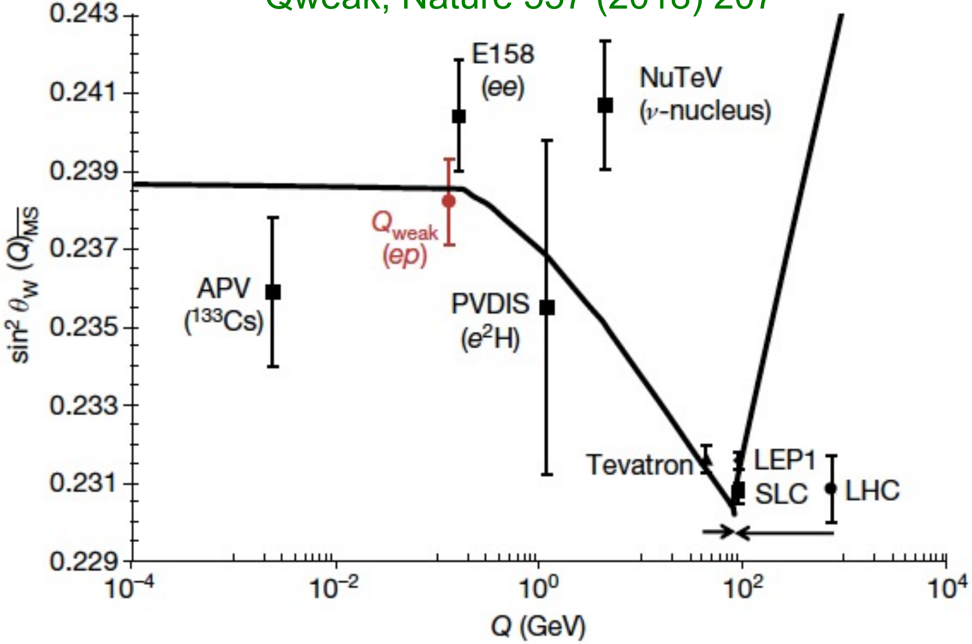
3. Paschos-Wolfenstein ratio and NuTeV anomaly



Manny Paschos (Dortmund)

$$R_{PW}=\frac{\sigma[\nu A\rightarrow\nu X]-\sigma[\bar{\nu}A\rightarrow\bar{\nu}X]}{\sigma[\nu A\rightarrow\mu X]-\sigma[\bar{\nu}A\rightarrow\mu^{+}X]}=\frac{1}{2}-sin^{2}\theta_{W}$$

Qweak, Nature 557 (2018) 207

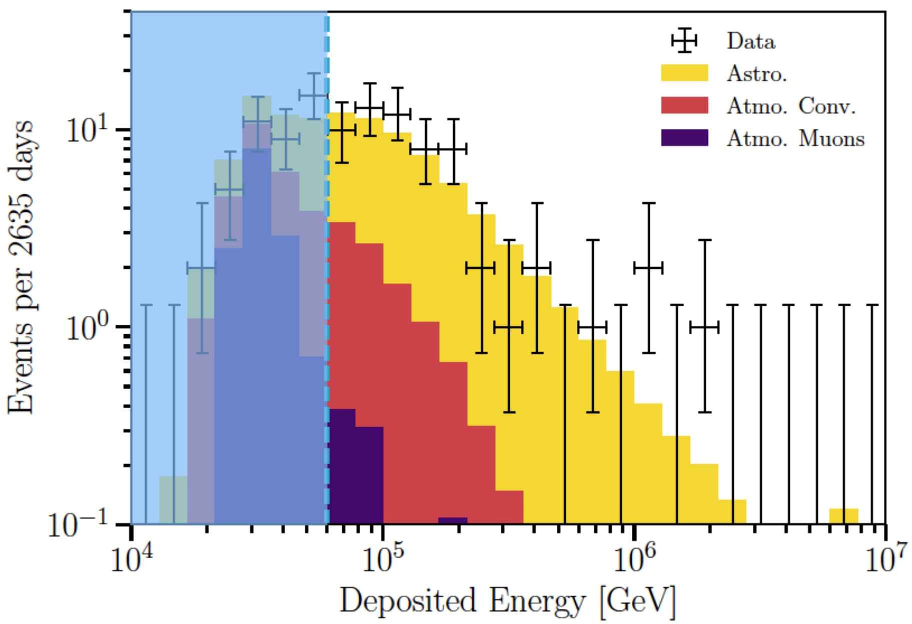


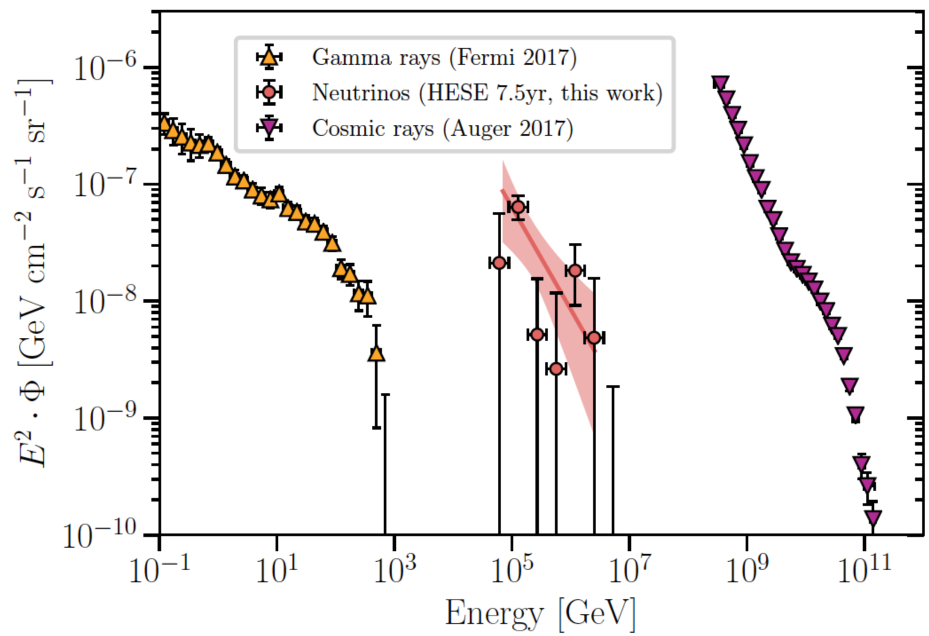
3. Astrophysical high-energy neutrino measurement

Data and MC agree up to ~PeV.
→ We more or less understand neutrino interactions up to ~PeV.

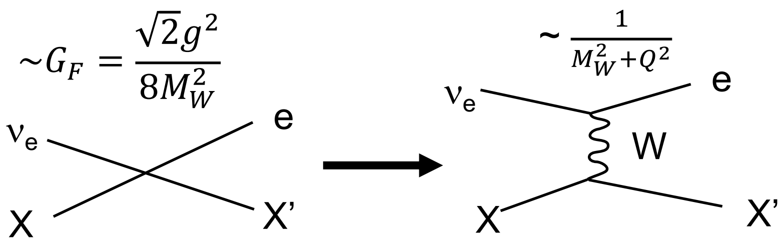
cross section, efficiency, flux, target

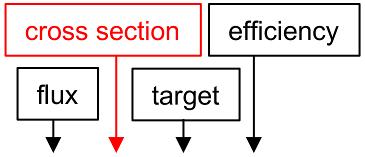
Event rate $N = \Phi \times \sigma \times T \times \varepsilon$



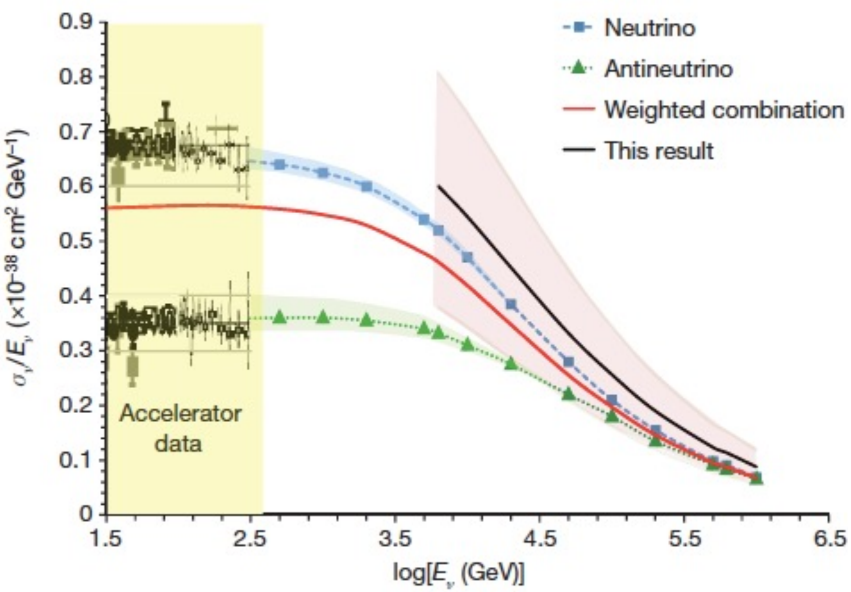


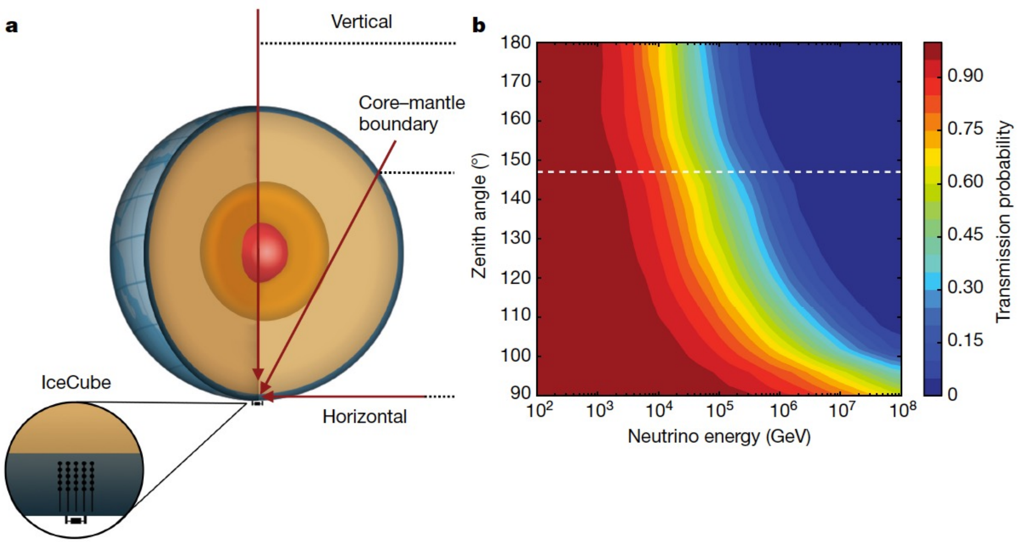
3. Neutrino DIS saturation





Event rate $N = \Phi \times \sigma \times T \times \varepsilon$



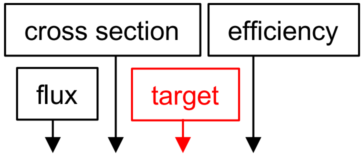


Dziewonski, Anderson (PREM), Phys. Earth Planet.Inter.25,(1981)297
Donini, Palomares-Ruiz, Salvado, Nature Physics 15(2019)37

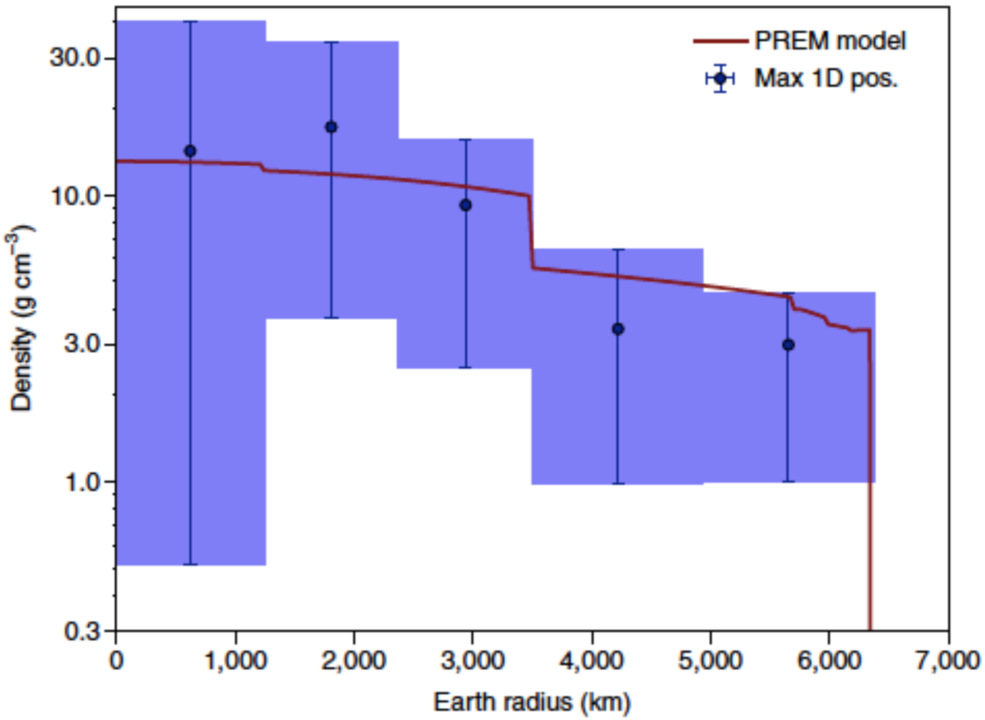
3. Earth tomography

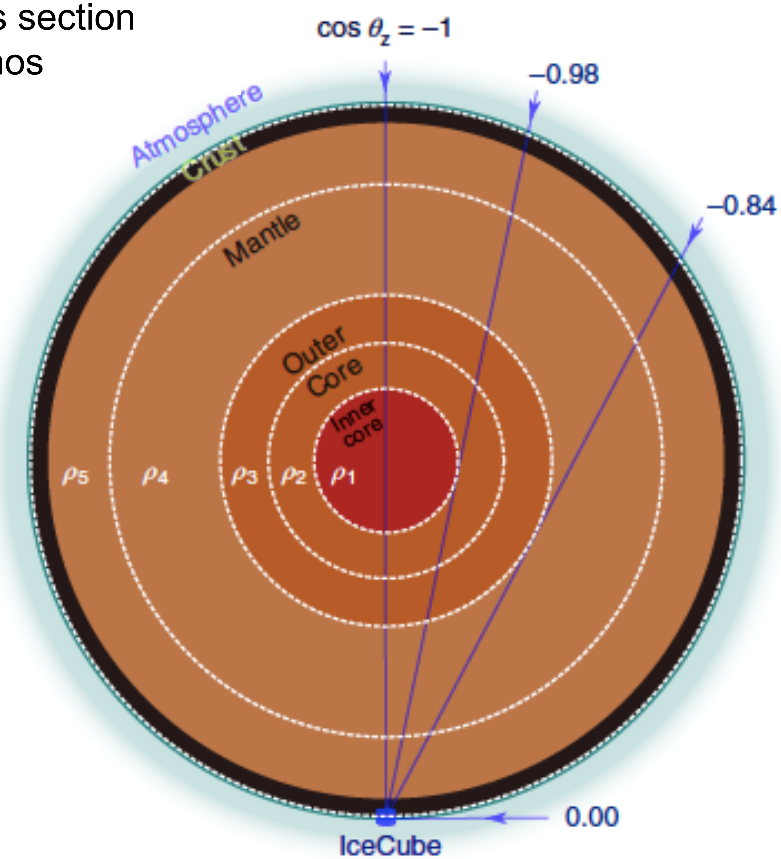
Earth absorption for Earth density measurement

- PREM (Preliminary reference Earth model)
- Standard earth density model used by T2K, NOvA, etc
- Earth density profile is extracted by assuming flux and cross section
- Measure Earth moment of inertia and Earth mass by neutrinos

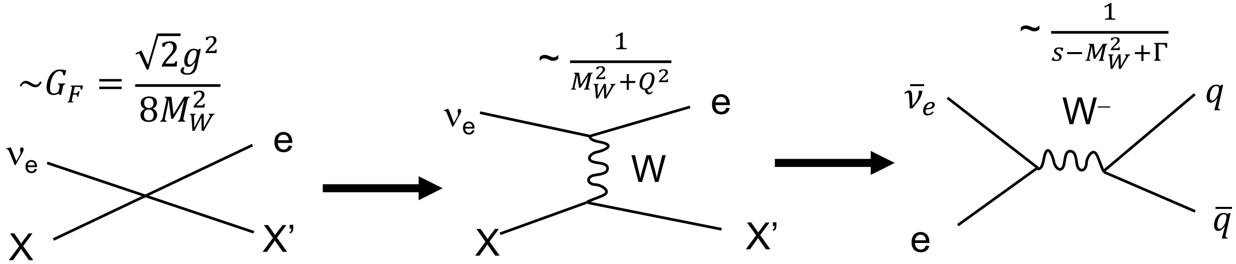


Event rate $\mathrm{N} = \Phi\times\sigma\times T\times\varepsilon$

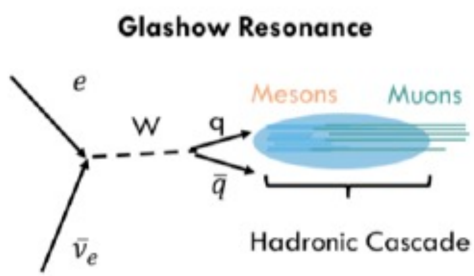




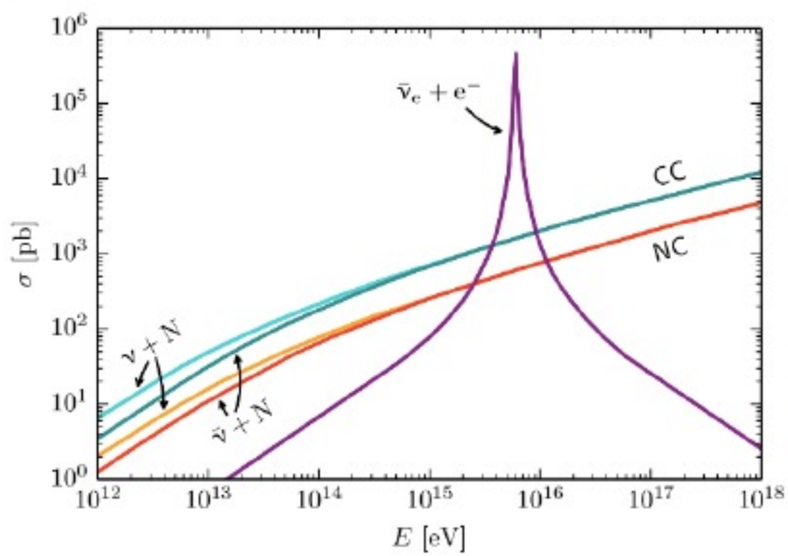
3. Glashow resonance

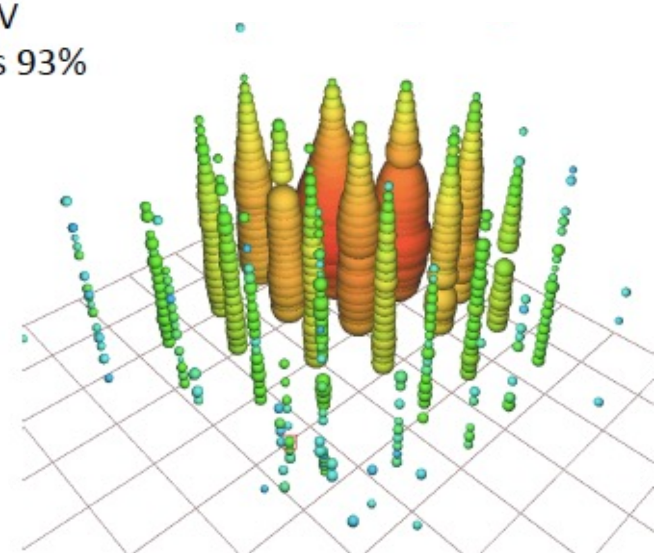


A 5.9 PeV event in IceCube



Resonance: E_ν = 6.3 PeV
Typical visible energy is 93%





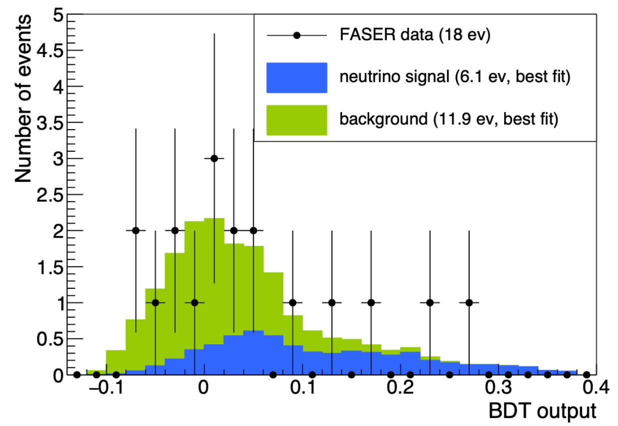
Event identified in a partially-contained PeV search (PEPE)

Deposited energy: 5.9±0.18 PeV (stat only)

ICRC 2017 arXiv:1710.01191

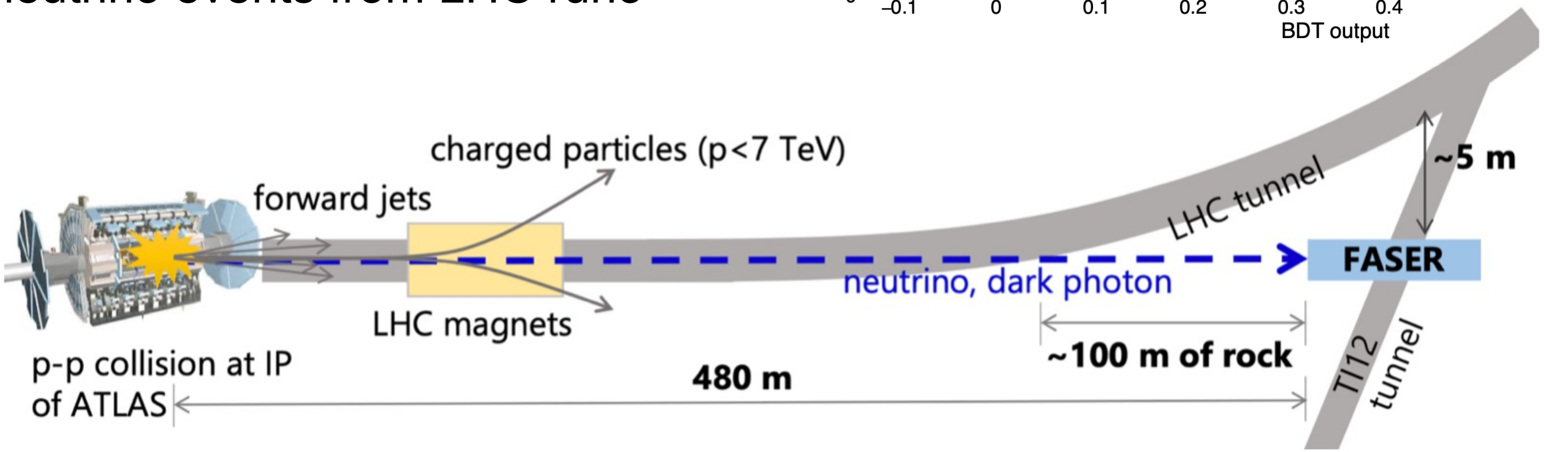
KING'S College LONDON

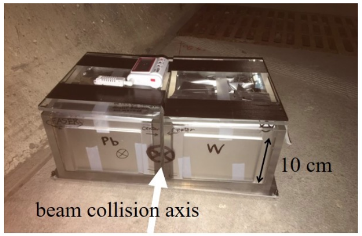
3. Collider neutrino

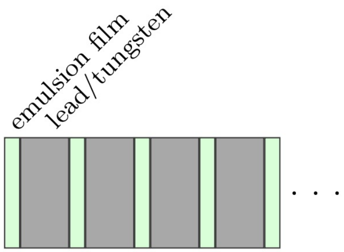


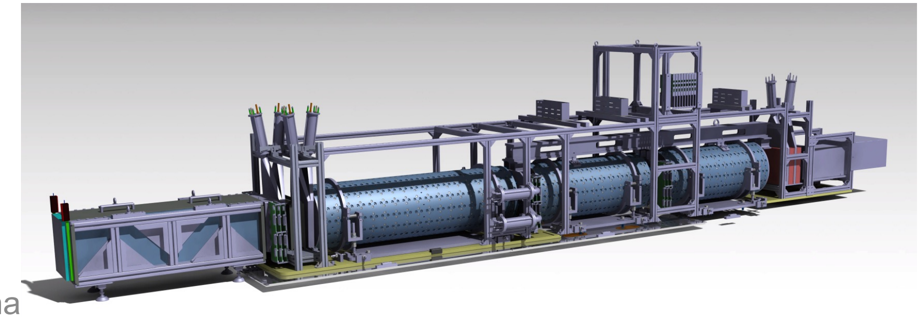
FASERnu

- Emulsion detector (high-resolution)
- neutrinos from ATLAS collision point
- neutrino excess from pilot run
- ~10,000 neutrino events from LHC run3





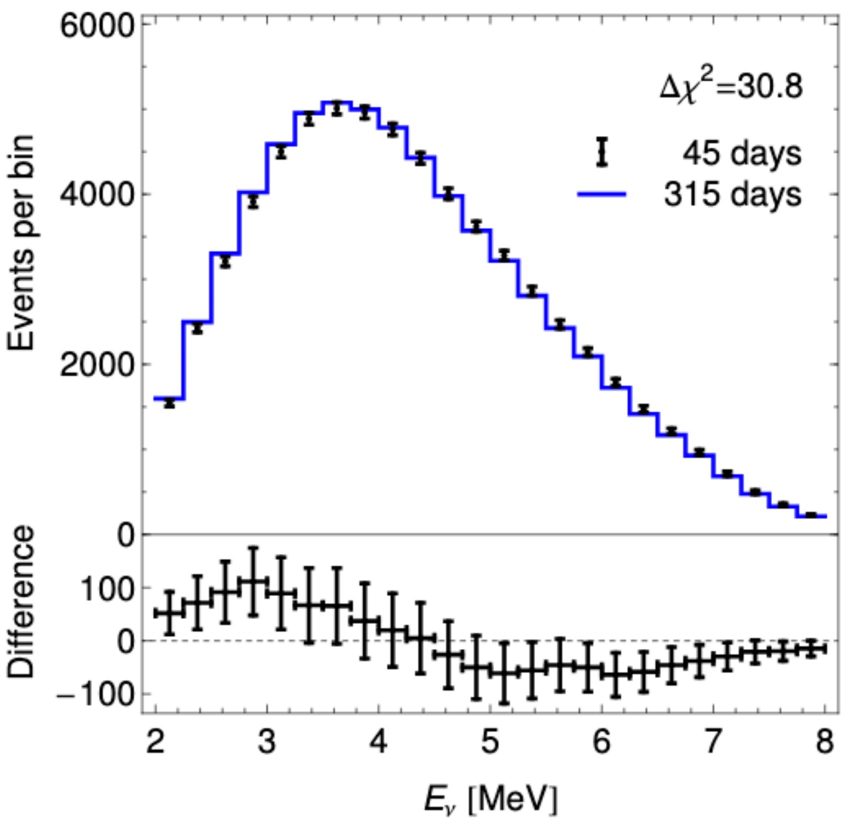


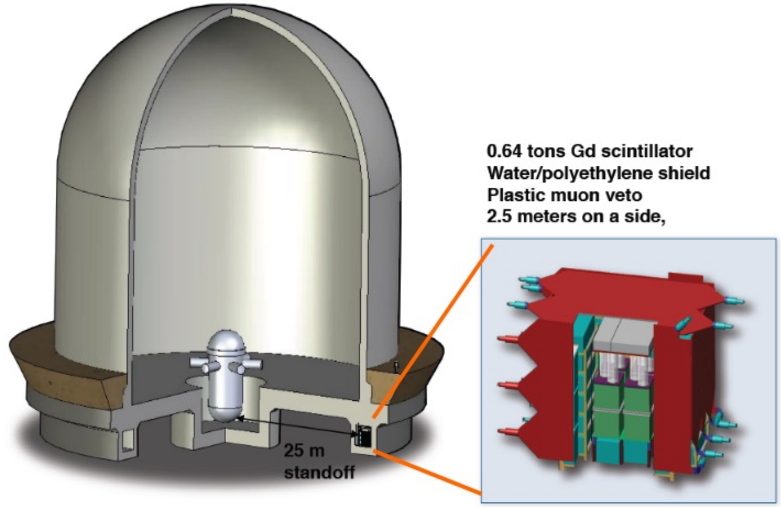


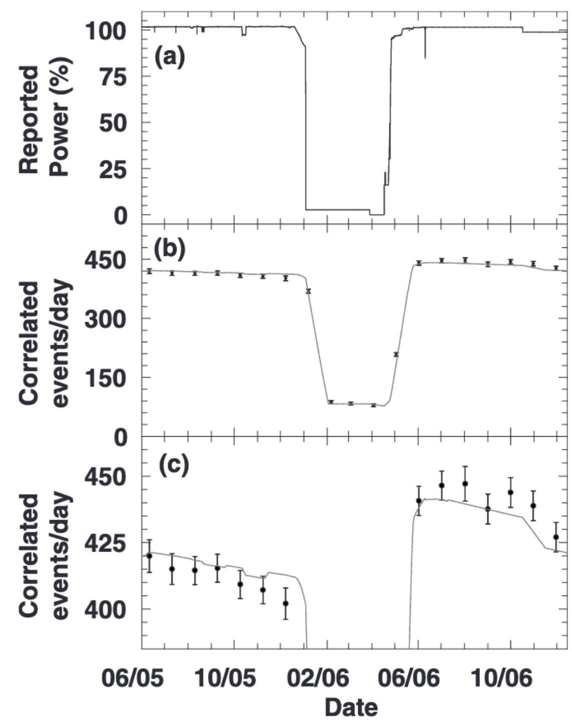
Akindele (Neutrieno2020), https://zenodo.org/record/3959532
SONGS, Journal of Applied Physics, 105(6), 064902 (2009)

4. Reactor neutrino

Low energy electron anti-neutrinos
- High-precision spectrum prediction
- Monitoring fission reactor







katori@fnal.gov

Cleveland et al.,Astrophys.J.496(1998)505
Pontecorvo,Phys.Lett.28B(1969)493

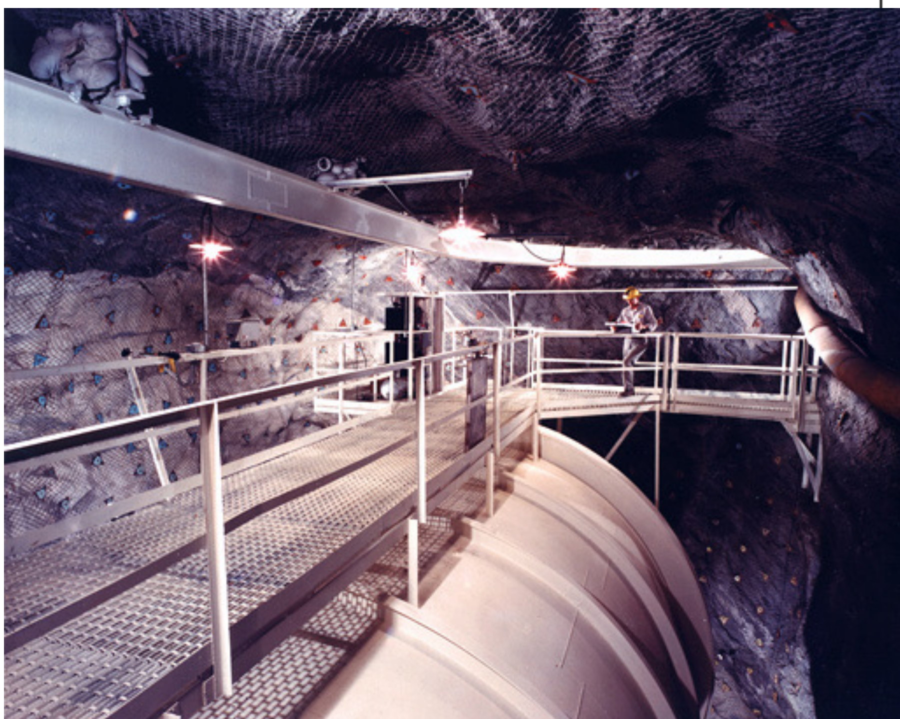
4. Solar neutrino

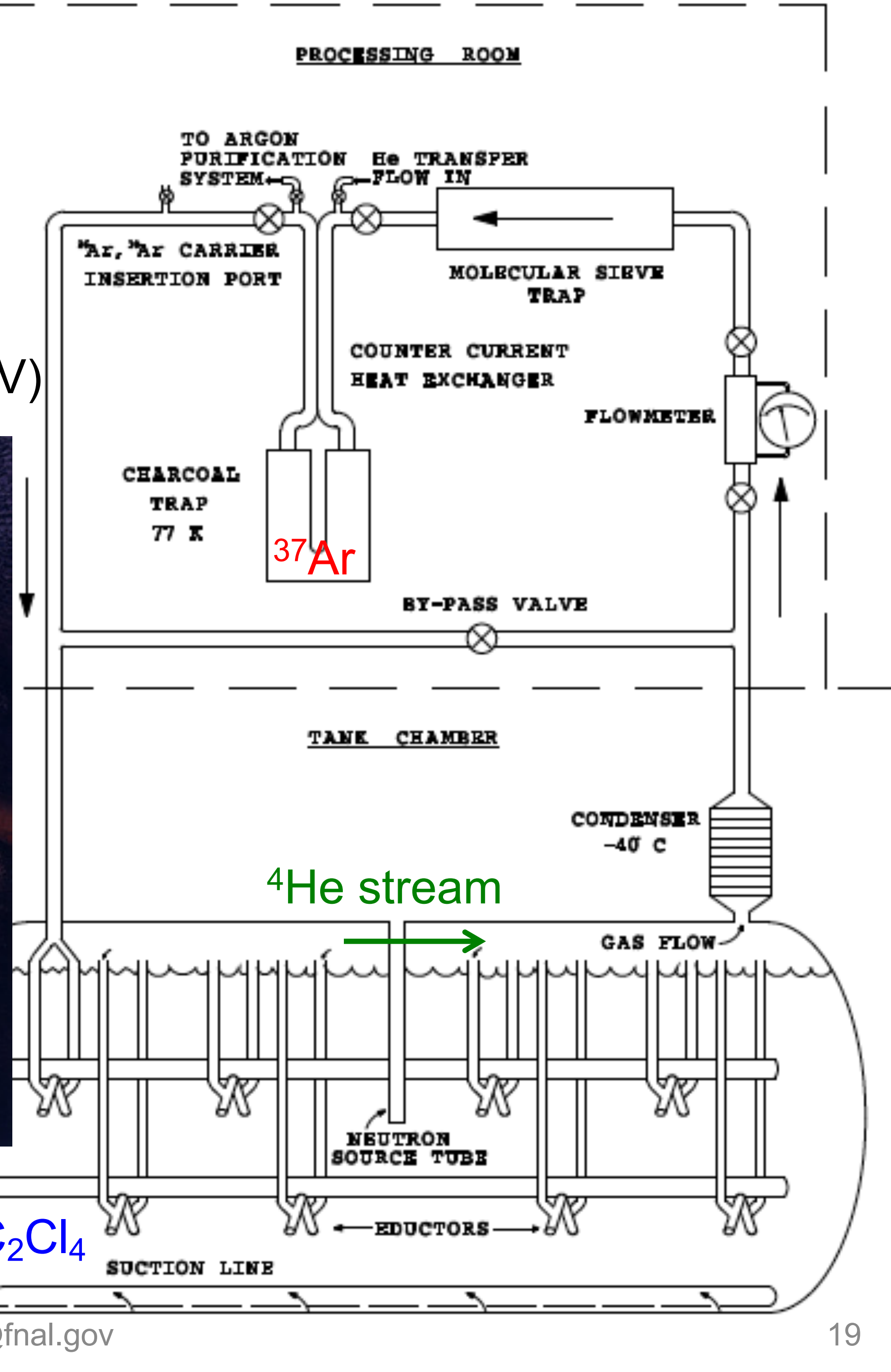
Homestake experiment

$$\nu_e + {}^{37}Cl \rightarrow e^- + {}^{37}Ar$$

(proposed by Pontecorvo)

- mainly sensitive to ^{8}B neutrino (~10 MeV)



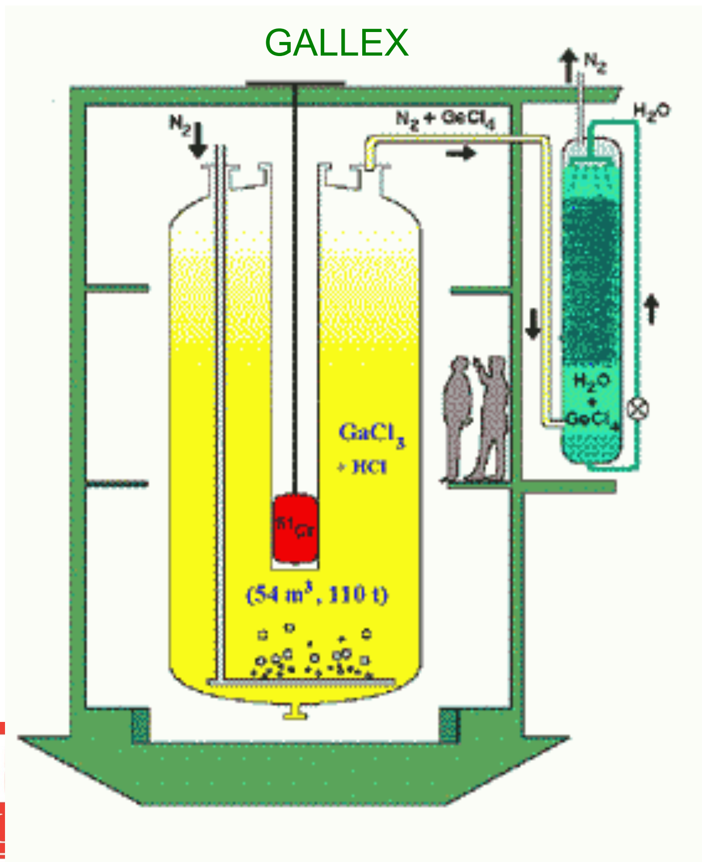


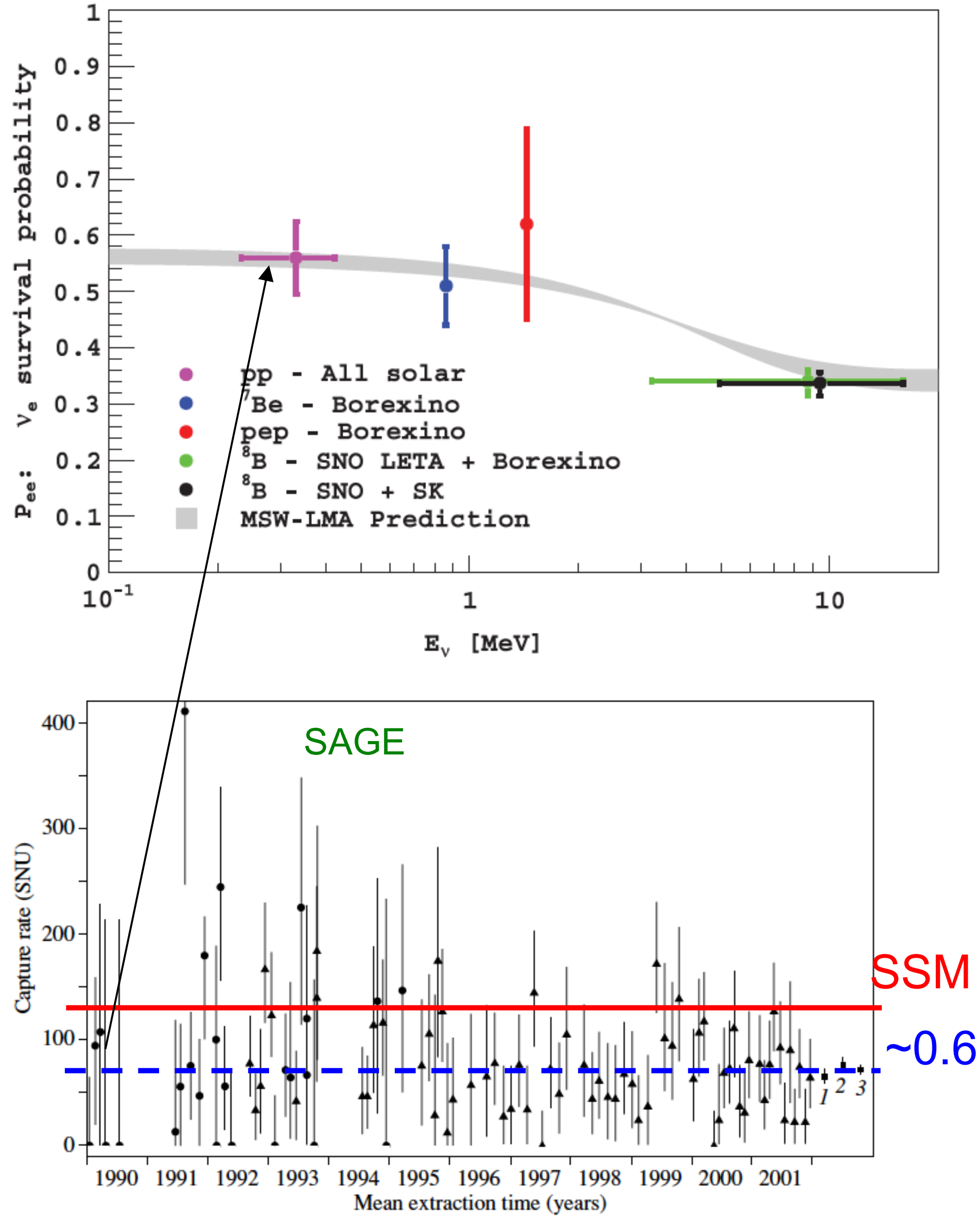
4. Solar neutrino

Gallium experiment

$$\nu_e + {}^{71}Ga \rightarrow e^- + {}^{71}Ge$$

- Sensitive to pp-neutrino (0.42 MeV), 90% of total solar neutrino flux.
- Both experiments observed deficit, but higher than Homestake result



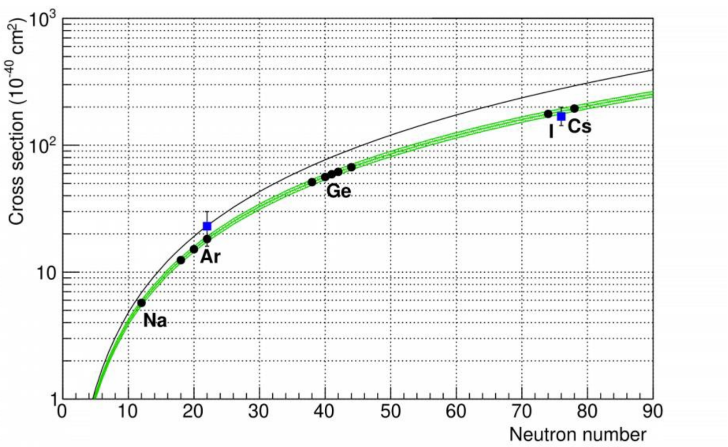


COHERENT, Science10.1126/science.aao0990 (2017),PRL126(2021)012002;129(2022)081801

4. Neutrino-Nucleus coherent scattering

Low energy neutrinos from neutron sources at SNS (spallation neutron source), ORNL (Oak Ridge National Lab)

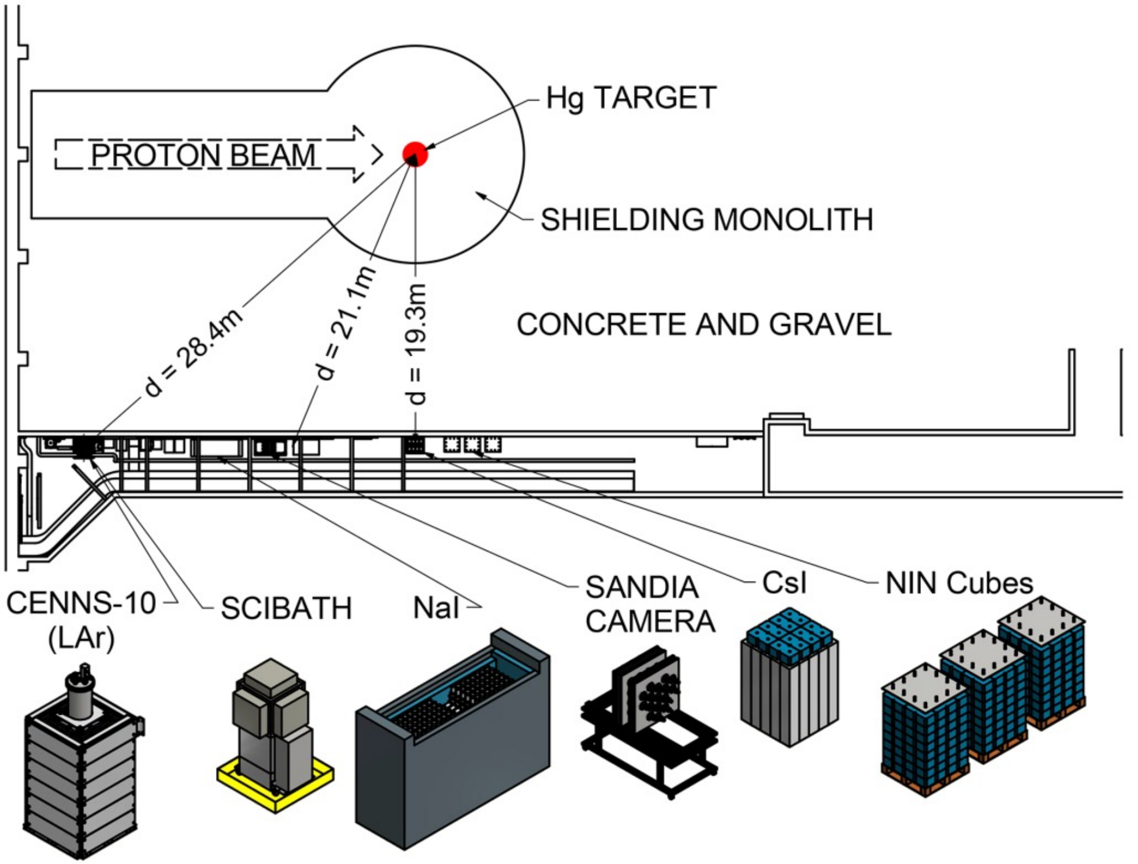
$$\nu + A \rightarrow \nu + A$$




Science REPORTS

Cite as: D. Akimov *et al.*, *Science* 10.1126/science.aao0990 (2017).

Observation of coherent elastic neutrino-nucleus scattering



Conclusion

ν-l scattering : well-known, test of weak theory

- Neutrino-electron scattering for neutrino flux measurement
- Anti-electron neutrino scattering for neutrino magnetic moment search (BSM)

ν-q scattering : test of weak theory, test of quark model

- DIS cross sections
- Di-muon production
- Paschos-Wolfenstein ratio
- Astrophysical neutrinos
- collider neutrinos

ν-A scattering :

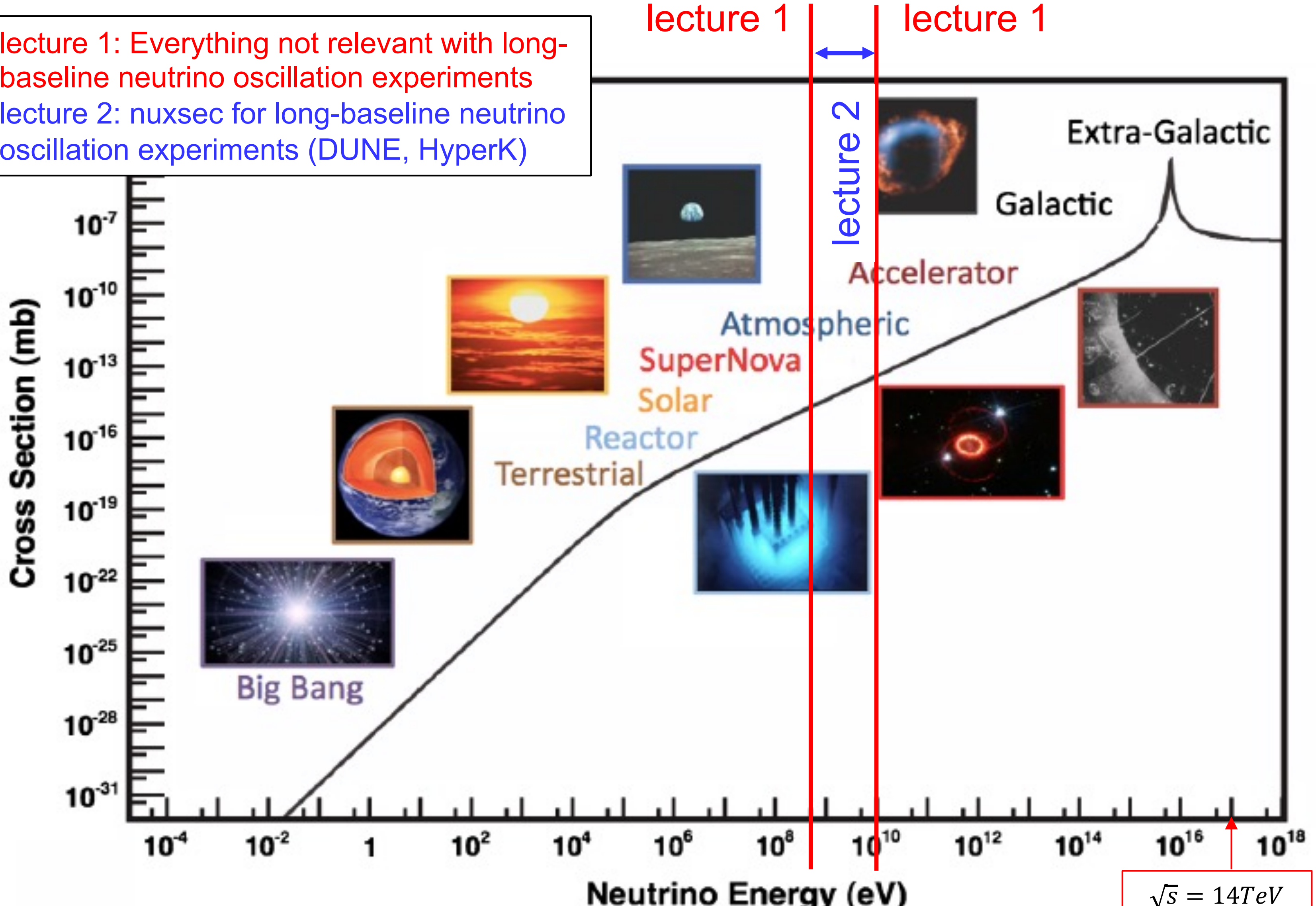
- Reactor neutrino experiments
- Neutrino nuclear capture by Cl and Ga, important for solar neutrinos
- Neutrino coherent scattering, important for supernova (2017)

ν-N scattering : important reactions for long baseline neutrino oscillation experiment
(T2K, NOvA, DUNE, Hyper-Kamiokande)

Formaggio and Zeller, Rev.Mod.Phys.84 (2012) 1307

1. From eV to EeV: Neutrino cross sections across energy scales

lecture 1: Everything not relevant with long-baseline neutrino oscillation experiments
lecture 2: nuxsec for long-baseline neutrino oscillation experiments (DUNE, HyperK)



$\sqrt{s} = 14TeV$

1. NuInt22 in Seoul (Oct. 24-29, 2022)

Neutrino interaction physics community
https://nuint22.org/

NuINT 2022

The 13th International Workshop on Neutrino-Nucleus Interactions
in the Few GeV Regions

October 24 to 29, 2022 (OFFLINE)

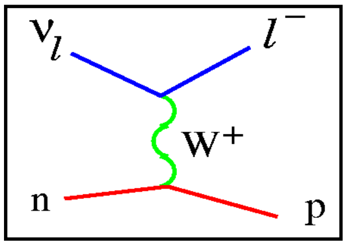
Hoam Faculty House
Seoul National University
Seoul, Korea

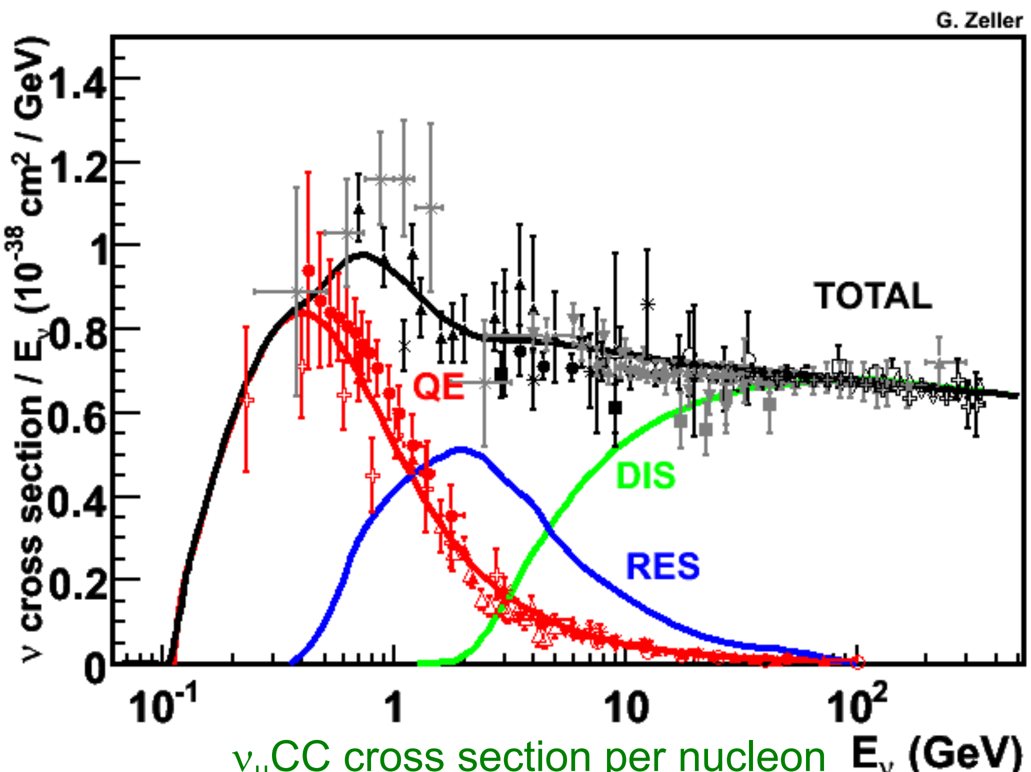
1. Next generation neutrino oscillation experiments

Accelerator-based neutrino oscillation experiments

- Present to Future: T2K, NOvA, Hyper-Kamiokande, DUNE

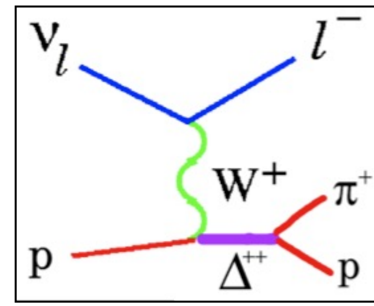
Quasi Elastic

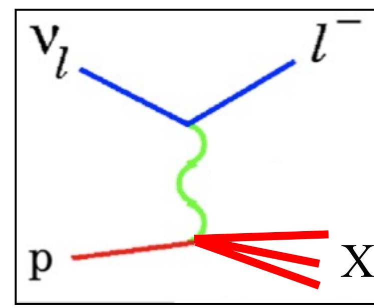




ν_μCC cross section per nucleon

baryonic RESonance





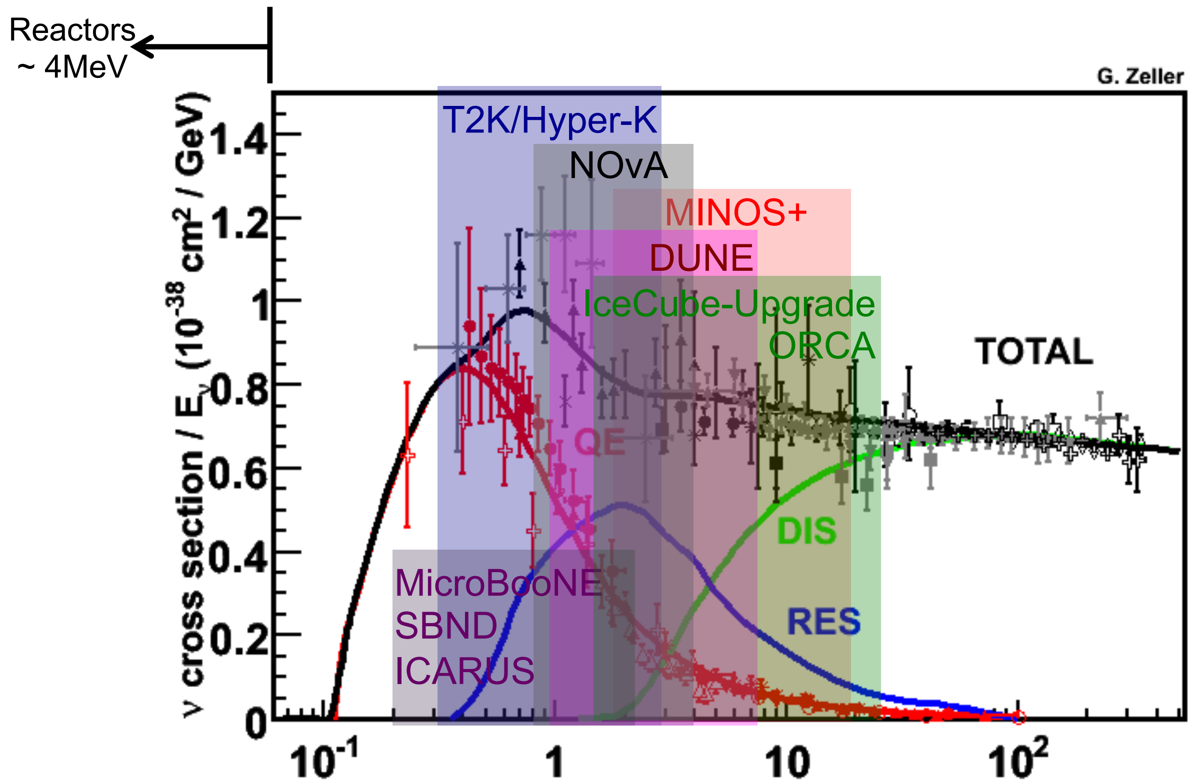
Deep Inelastic Scattering

Formaggio and Zeller, Rev.Mod.Phys.84(2012)1307

1. Next generation neutrino oscillation experiments

Accelerator-based neutrino oscillation experiments

- Present to Future: T2K, NOvA, Hyper-Kamiokande, DUNE…



ν_μCC cross section per nucleon

$$P_{\mu\to e}(L/E) = \sin^2 2\theta \sin^2\left(1.27\Delta m^2(eV^2)\frac{L(km)}{E(GeV)}\right)$$

1. Next generation neutrino oscillation experiments

Accelerator-based neutrino oscillation experiments

- Present to Future: T2K, NOvA, Hyper-Kamiokande, DUNE…

Most of data are from muon neutrino beam

- create by π-DIF, K-DIF (pion and kaon decay-in-flight)
- $\Phi(\nu_\mu) > \Phi(\bar{\nu}_\mu)$: more π^+ and K^+ than π^- and K^- (for low energy accelerators)
- μ-decay can make electro-neutrinos but they are background
- δ_{CP} study need electro-neutrino and antineutrino cross-sections (ν_e appearance)

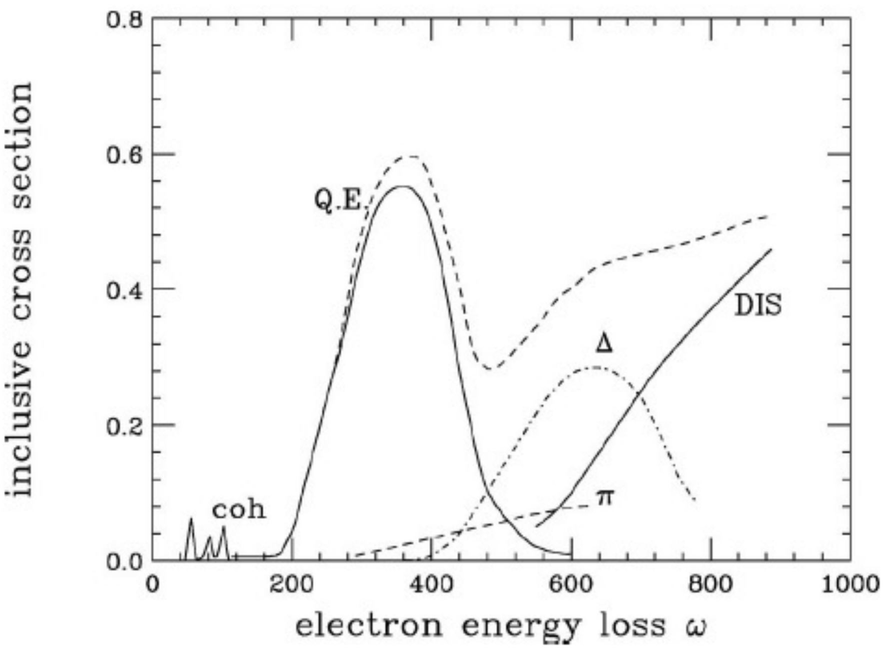
Nuclear physics sucks

- Simple extrapolation may be broken due to nuclear physics
- We are not good at nuclear physics because we are not nuclear physicists
- Nuclear physics = non-perturbative QCD (many models, no theory)
- Particle physics is developed by avoiding nuclear physics…

$$P_{\mu \to e}(L/E) = \sin^2 2\theta \sin^2\left(1.27\Delta m^2(eV^2)\frac{L(km)}{E(GeV)}\right)$$

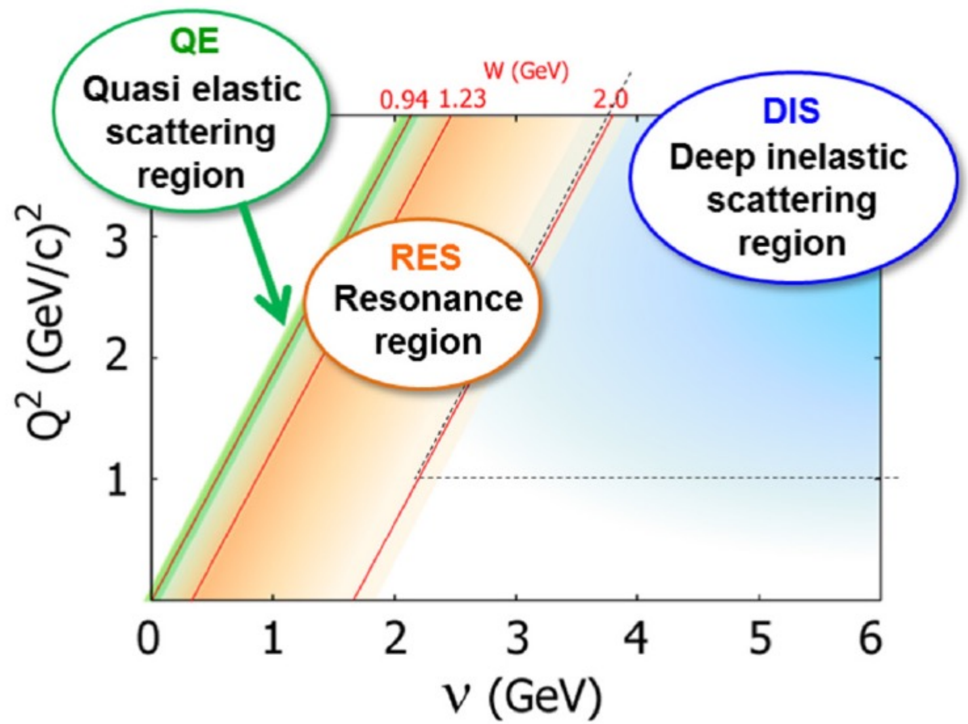
Benhar,Day,Sick,Rev.Mod.Phys.80(2008)189
Nakamura et a;,Rep.Prog.Phys.80(2017)056301

1. Particle Physics vs. Nuclear Physics



Nuclear physics

Interactions are classified in q (3-momnetum transfer) and ω (energy transfer)



Particle physics (neutrino physics)

Interactions are classified in Q^2 (4-momentum transfer) and ν (energy transfer) or W^2 (invariant mass)

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katc

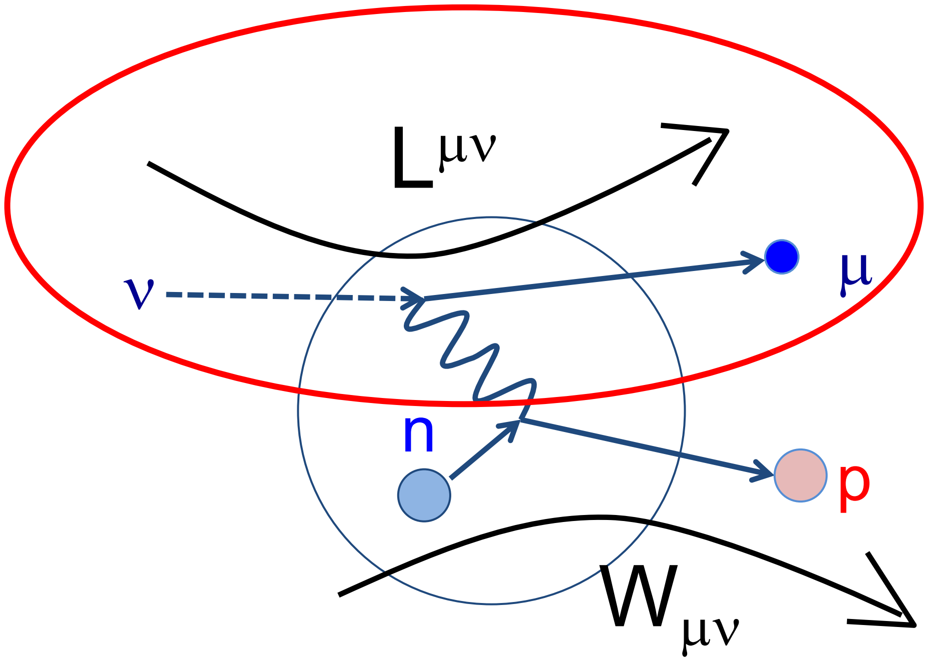
1. Neutrino cross-section formula

Cross-section
 - product of Leptonic and Hadronic tensor

$$d\sigma \sim L^{\mu\nu}W_{\mu\nu}$$

Leptonic tensor
 → the Standard Model (easy)

Hadronic tensor
 → nuclear physics (hard)



1. Neutrino cross-section formula

Cross-section

- product of Leptonic and Hadronic tensor

$$d\sigma \sim L^{\mu\nu}W_{\mu\nu}$$

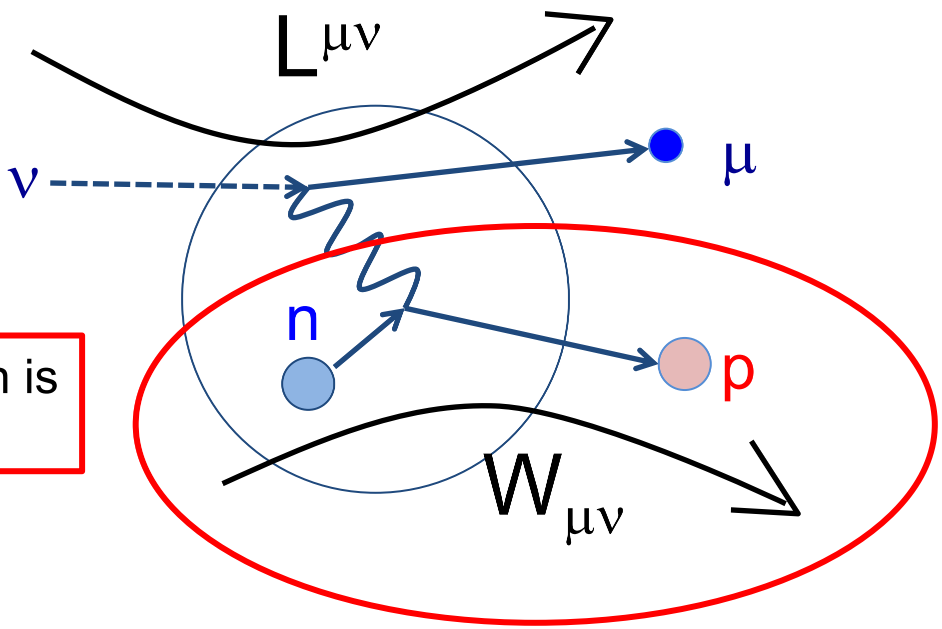
Leptonic tensor

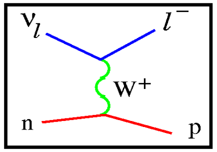
→ the Standard Model (easy)

Hadronic tensor

→ nuclear physics (hard)

All complication of neutrino cross-section is how to model the hadronic tensor part

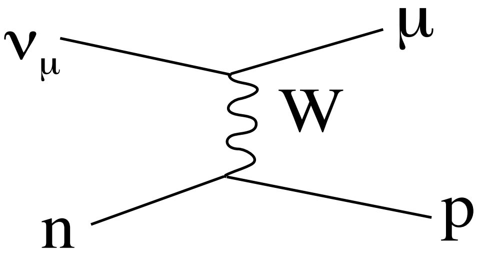




2. Charged Current Quasi-Elastic scattering (CCQE)

The simplest and the most abundant interaction around ~1 GeV.

$$\nu_\mu + n \rightarrow p + \mu^- \quad (\nu_\mu + X \rightarrow X' + \mu^-)$$



$d\sigma \sim L_{\mu\nu} T^{\mu\nu}$

$L_{\mu\nu} \sim J_\mu J_\nu$: Lepton tensor

$W_{\mu\nu} = \int f(\vec{k}, \vec{q}, \omega) T_{\mu\nu} dE$: hadronic tensor

$f(\vec{k}, \vec{q}, \omega)$: nucleon phase space

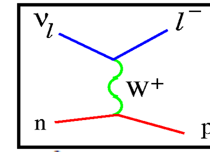
$T_{\mu\nu} = T_{\mu\nu}(F_1, F_2, F_A, F_P)$: form factors

Form factors can be parameterized with dipole form

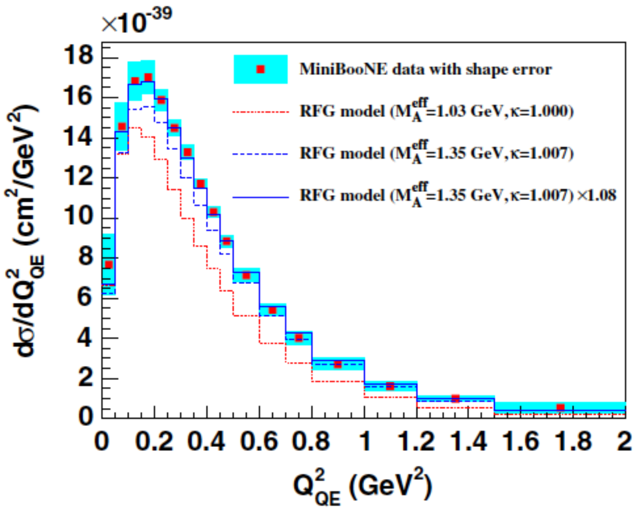
$$F(Q^2) = \frac{g}{\left(1 + \frac{Q^2}{M^2}\right)^2}$$

2. Form factors

Quasi Elastic



MiniBooNE CCQE cross section
PRD81(2010)092005



Form factors can be parameterized with dipole form

$$F(Q^2) = \frac{g}{\left(1 + \frac{Q^2}{M^2}\right)^2}$$

ρ(r)	$\|F(q^2)\|$	Example
pointlike	constant	Electron
exponential	dipole	Proton
gauss	gauss	^{6}Li
homogeneous sphere	oscillating	–
sphere with a diffuse surface	oscillating	^{40}Ca

r ⟶ $\|q\|$ ⟶

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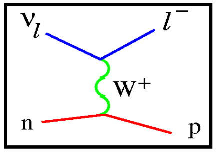
Quasi Elastic



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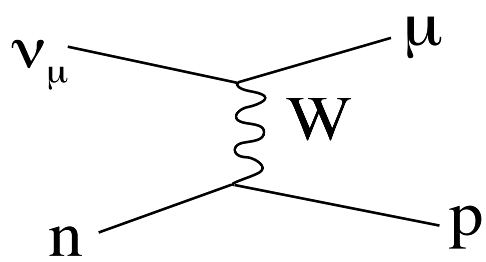
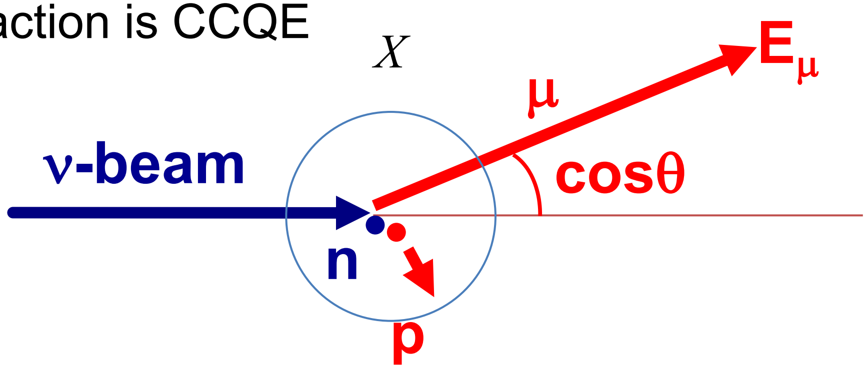


Neutrino energy is reconstructed from the observed lepton kinematics
"QE assumption"

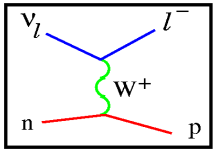
1. assuming neutron at rest
2. assuming interaction is CCQE



$$E_\nu^{QE} = \frac{ME_\nu - 0.5m_\mu^2}{M - E_\mu + p_\mu cos\theta}$$

Neutrinos hit nucleons inside of nucleus, and the energy reconstruction is possible only with QE assumption

Quasi Elastic



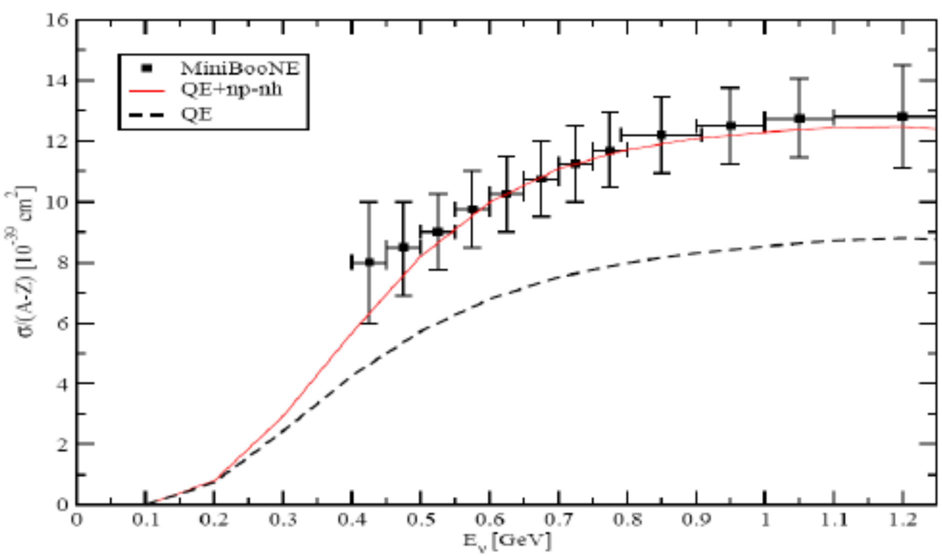
2. Nucleon correlations

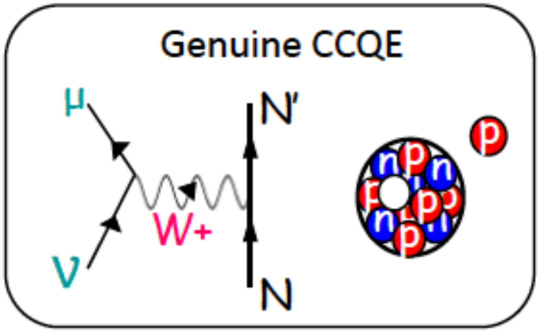
2-particle 2-hole (2p2h) effect

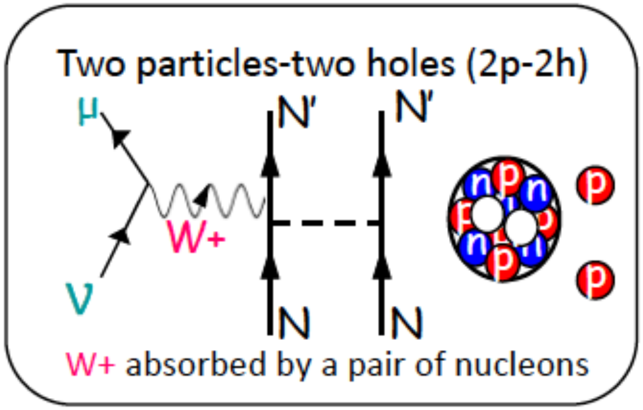
- Mimic CCQE interaction, significant change cross section (both shape and normalization)
- The biggest topic in nuxsec community (T2K, NOvA, MINERvA, MicroBooNE)

An explanation of this puzzle

Inclusion of the multinucleon emission channel (np-nh)

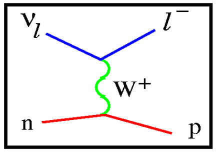






Martini et al, PRC80(2009)065501
NOvA, EPJC80,1119(2020)

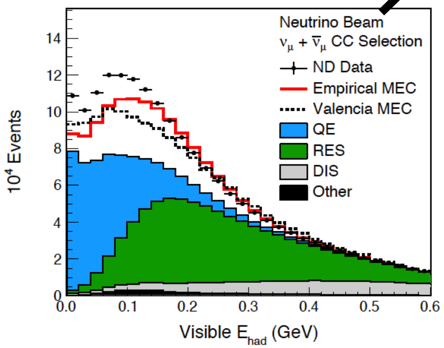
2. Nucleon correlations

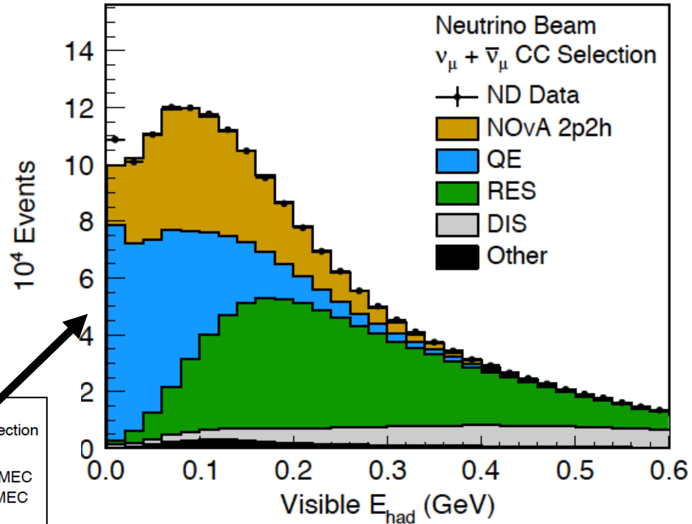


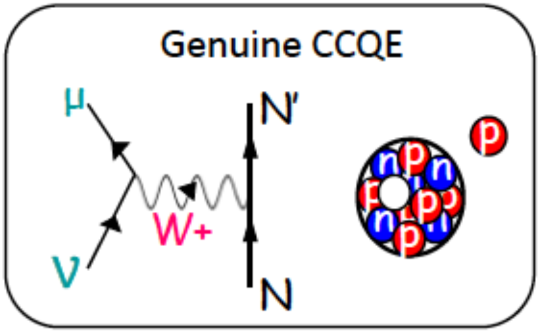
2-particle 2-hole (2p2h) effect

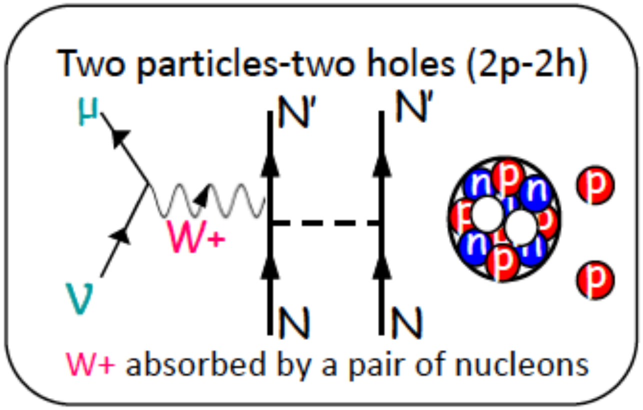
- Mimic CCQE interaction, significant change cross section (both shape and normalization)
- The biggest topic in nuxsec community (T2K, NOvA, MINERvA, MicroBooNE)
- 2p2h models in generators don't describe data well?
- High resolution detector (LArTPC, emulsion, etc) can find what is going on?

NOvA near detector data-MC comparison after fit





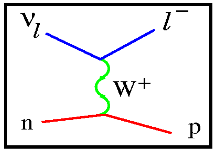




Smith and Moniz, NPB43(1972)605
Khachatryan et al., Nature 599(2021)565

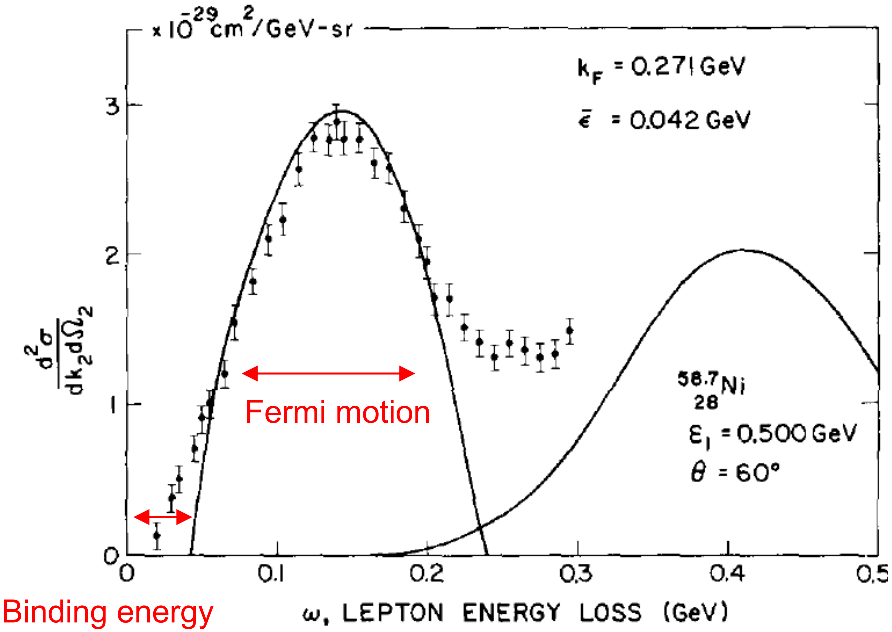
2. Fermi motion

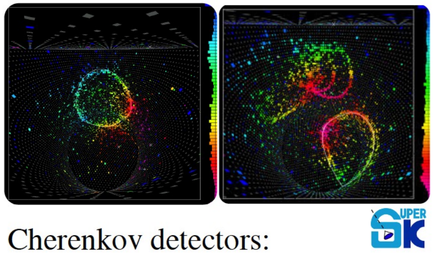
Quasi Elastic



Fermi motion

- Measured energy is smeared from the true energy if you assume nucleon at rest
- High resolution detector can measure all outgoing hadrons
  → initial nucleon momentum can be reconstructed (no Fermi motion smearing)



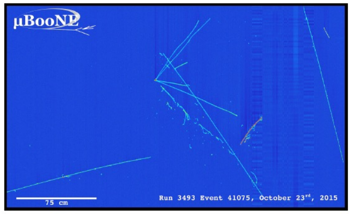


Cherenkov detectors:

Assuming QE interaction

Using lepton only

$$E_{QE} = \frac{2M\epsilon + 2ME_l - m_l^2}{2(M - E_l + |k_l|\cos\theta_l)}$$



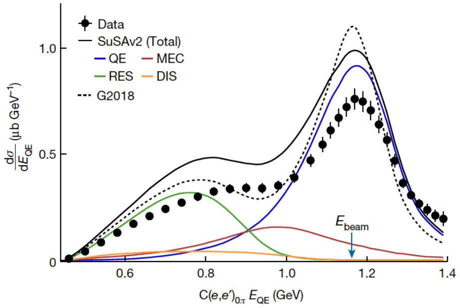
Tracking detectors:

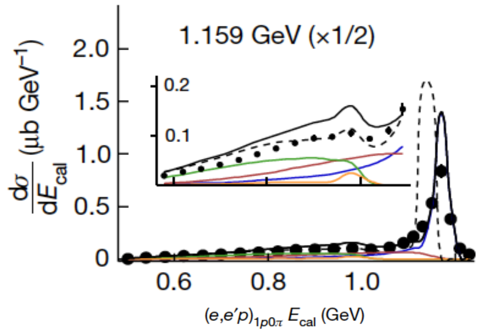
Calorimetric sum

Using All detected particles

$$E_{\mathrm{cal}} = E_l + E_p^{\mathrm{kin}} + \epsilon$$

[1p0π]



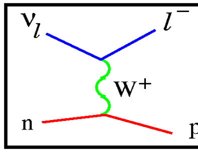


QE formula (HyperK) Calorimetric (DUNE)

Smith and Moniz, NPB43(1972)605
MiniBooNE, PRD81(2010)092005

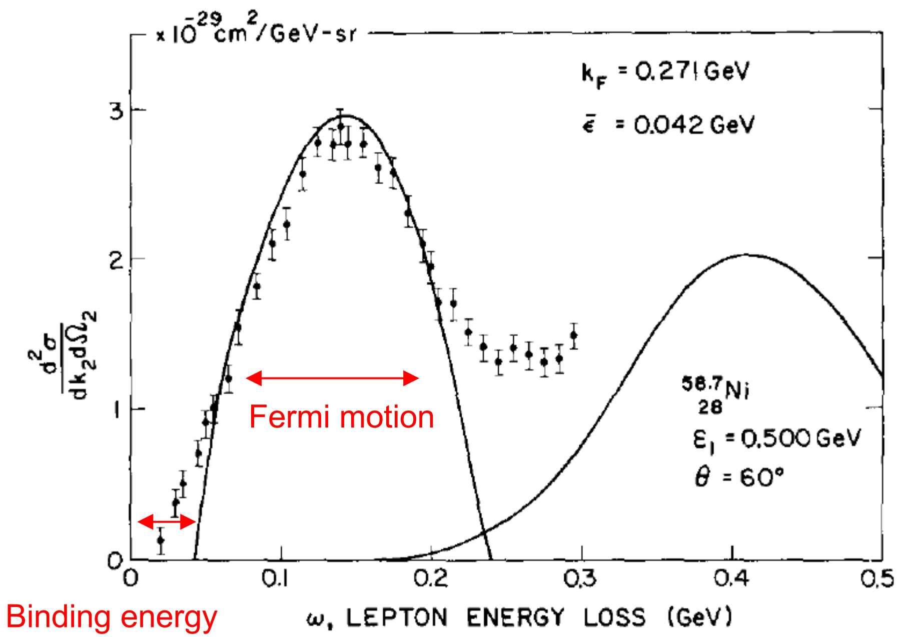
2. Pauli blocking

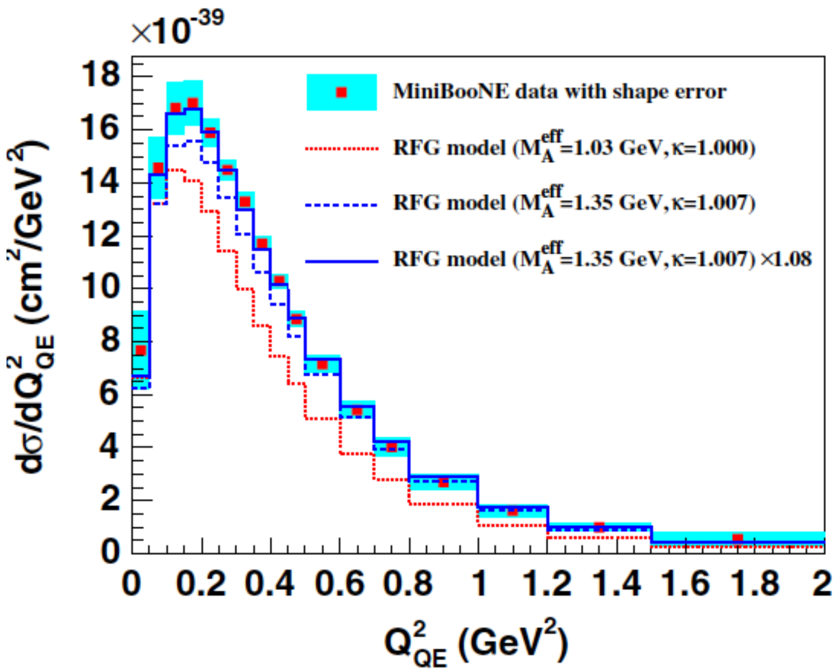
Quasi Elastic



Pauli blocking

- Low momentum transfer reaction is forbidden.
- data show more suppression than what Pauli blocking can → RPA(?)
- In the global Fermi gas model, Pauli blocking looks unphysical

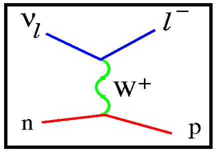




Smith and Moniz, NPB43(1972)605
Bodek and Cai, EPJC79(2019)293

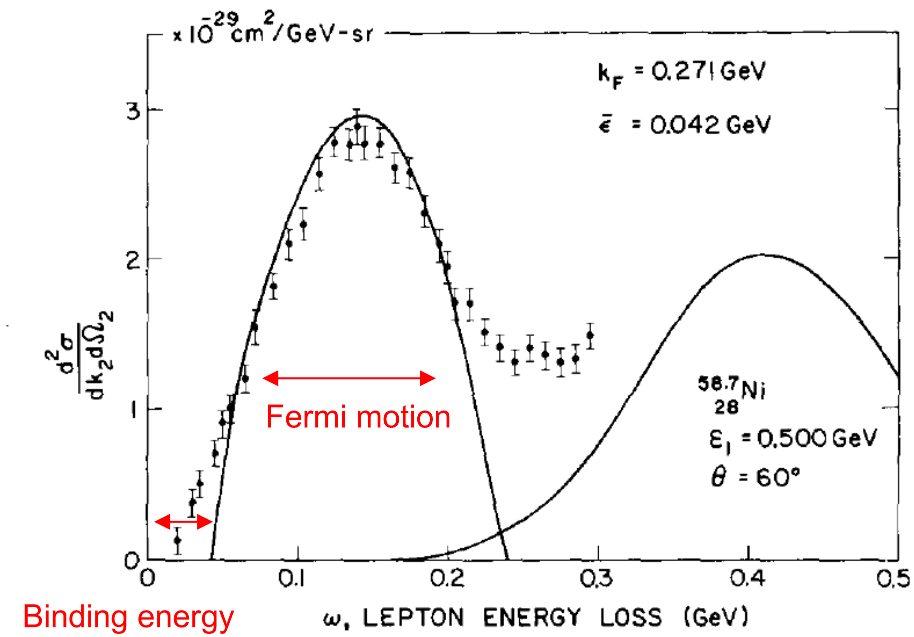
2. Nuclear Shell structure and binding energy

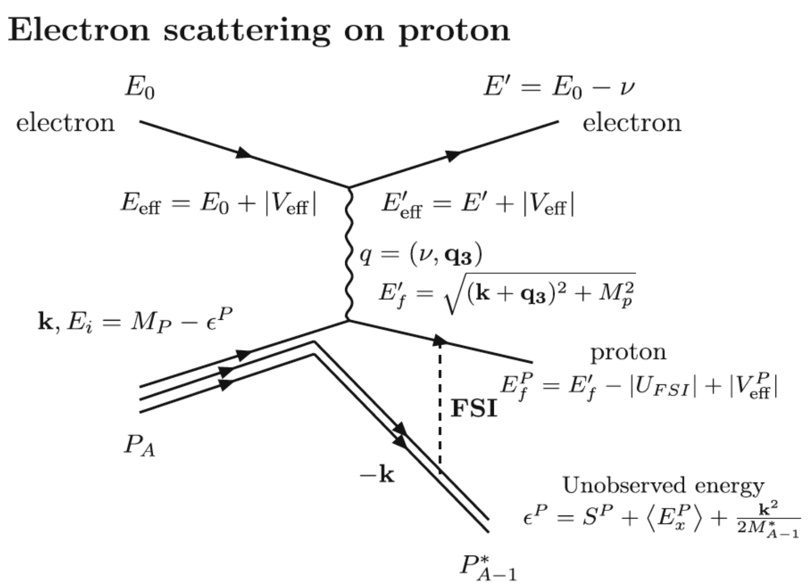
Quasi Elastic



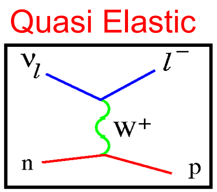
Binding energy ~ unobserved energy

- Energy to cost to release 1 nucleon, not constant
- Separation energy + excitation energy + recoil energy
  - Separation energy: energy to release 1 nucleon from the shell (~15 MeV, depends)
  - Excitation energy: energy used to excite leftover target nucleus (~1 MeV)
  - Recoil energy: kinetic energy of recoil target nucleus (~2-3 MeV)



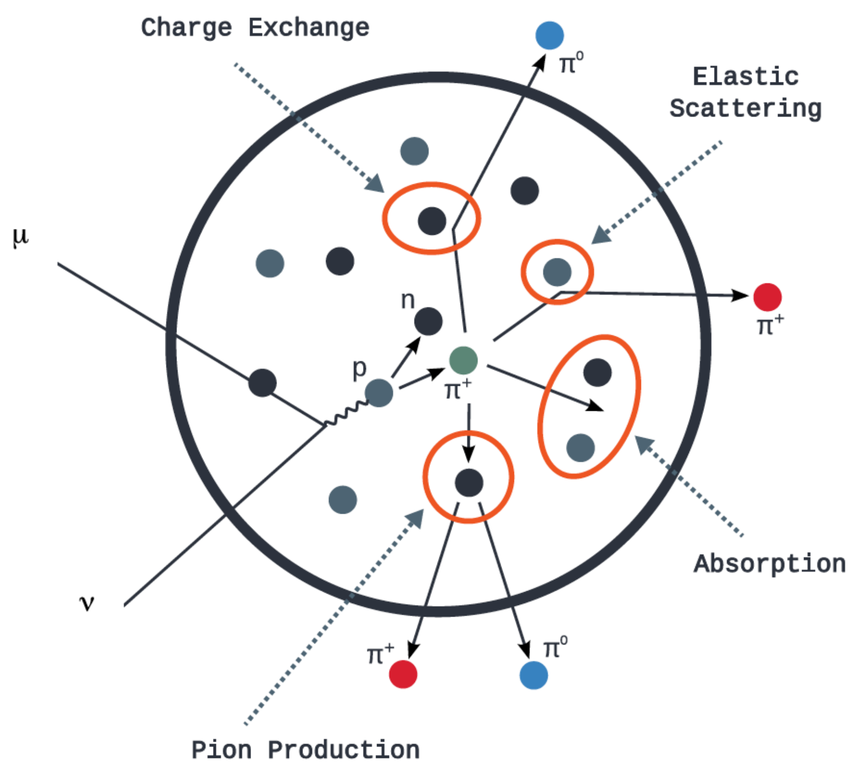


2. Final state interaction



Cascade model

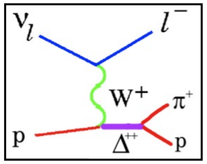
- Elastic scattering: Nucleon elastic scattering, pion elastic scattering
- Inelastic scattering: Nucleon inelastic scattering, pion inelastic scattering
- Charge exchange: Nucleon charge exchange, pion charge exchange
- Absorption: Nucleon absorption, pion absorption



MiniBooNE,PRD83(2011)052009
Lalakulich et al,PRC87(2013)014602

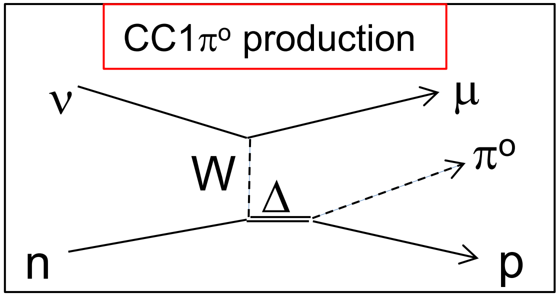
3. Neutrino Baryonic resonance data

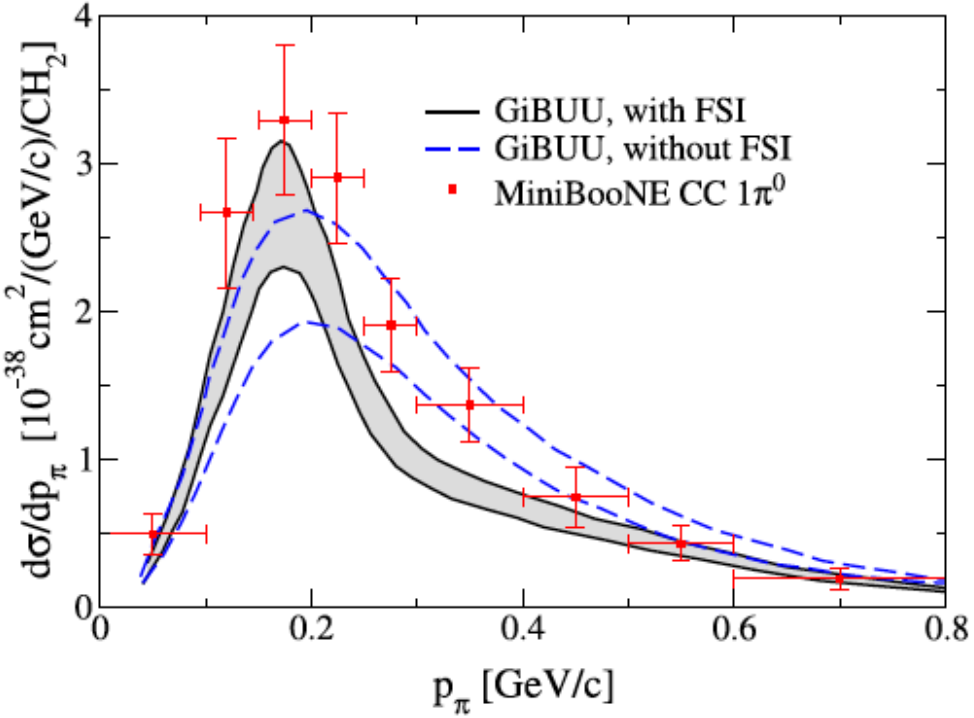
RESonance



Final state interaction

- Cascade model as a standard of the community
- Advanced models are not available for event-by-event simulation





MiniBooNE π° momentum vs simulation

MiniBooNE,PRD83(2011)052009
Lalakulich et al,PRC87(2013)014602

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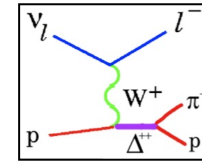

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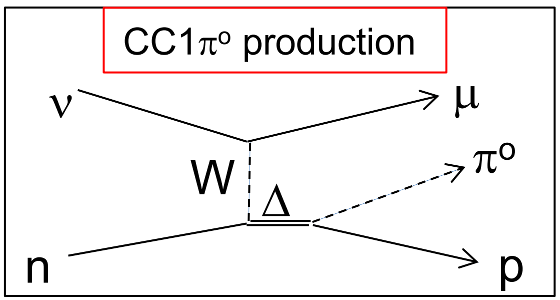

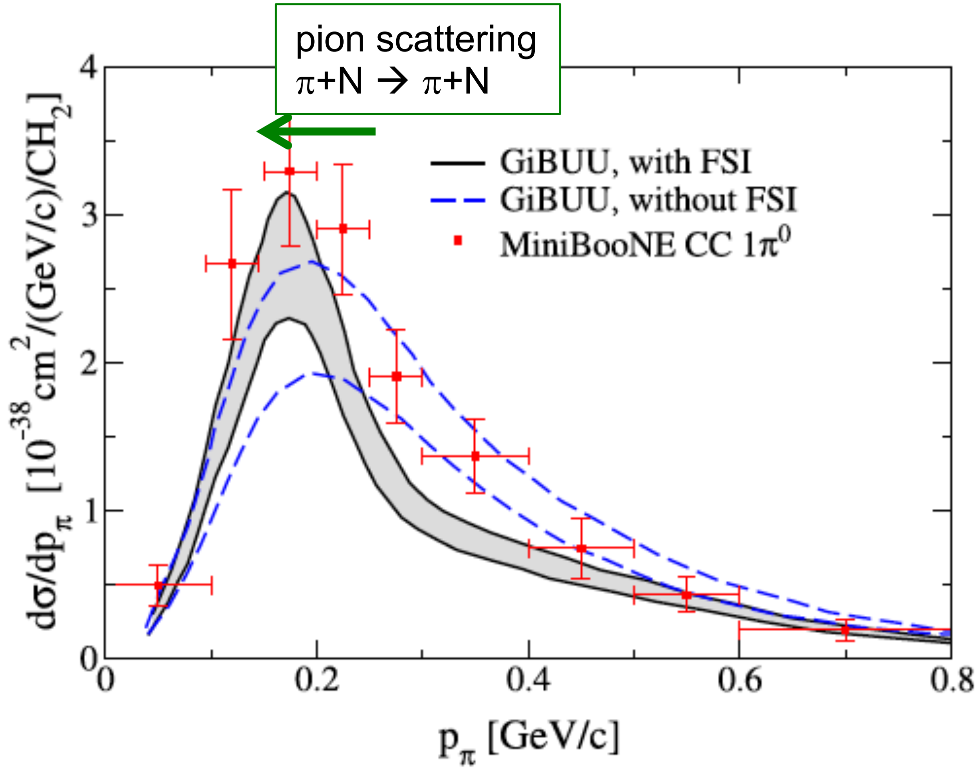

MiniBooNE π° momentum vs simulation

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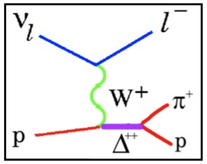

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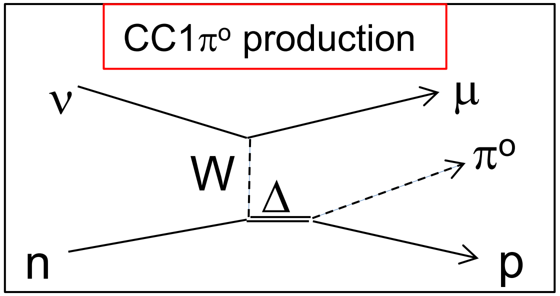

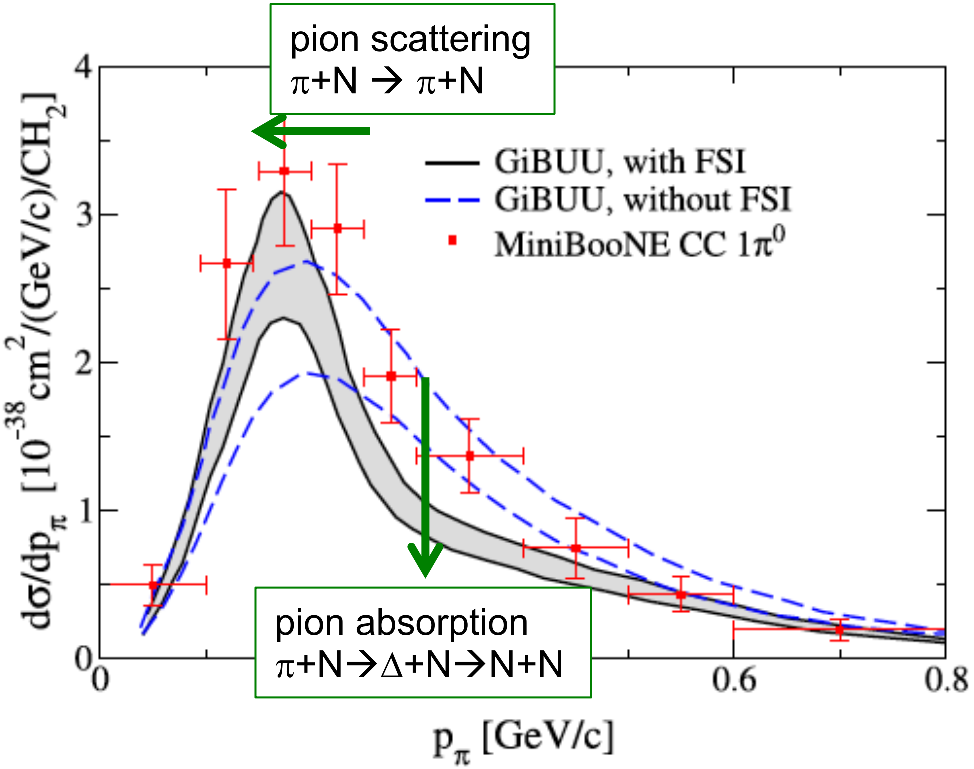

MiniBooNE πº momentum vs simulation

MiniBooNE,PRD83(2011)052009
Lalakulich et al,PRC87(2013)014602

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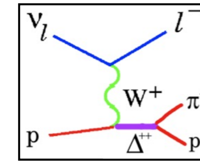

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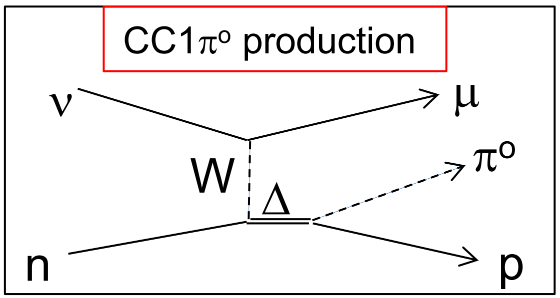
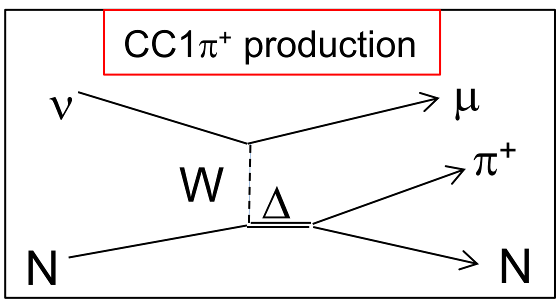
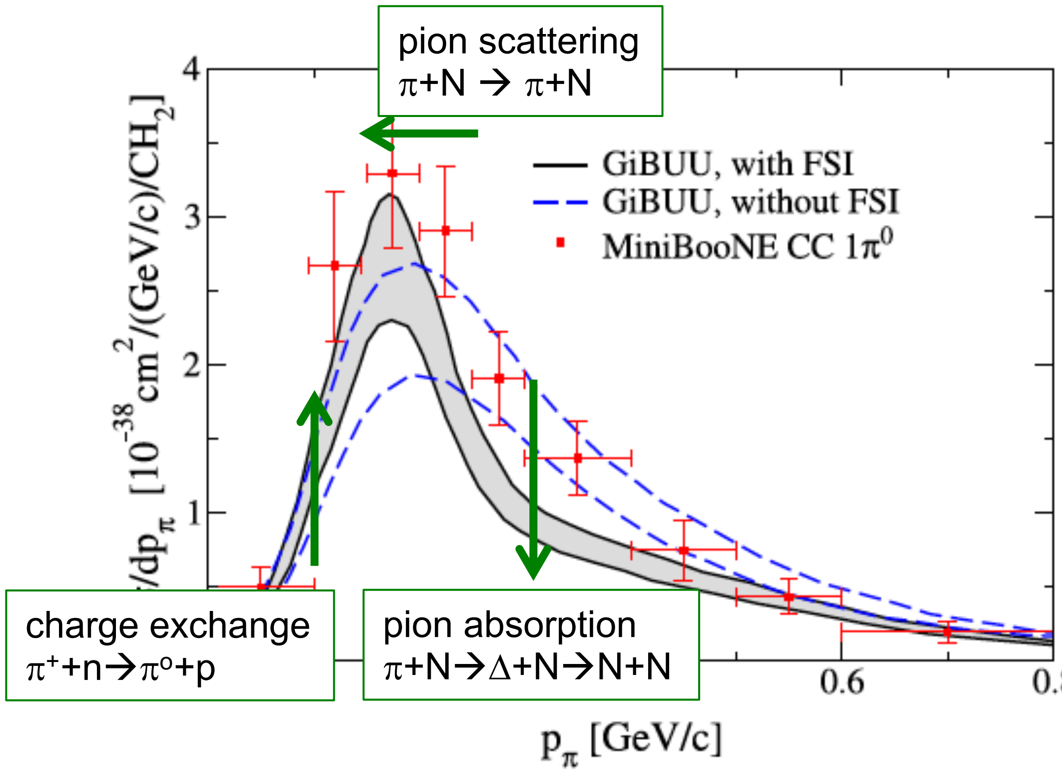




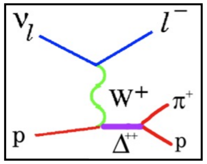


MiniBooNE π^o momentum vs simulation

All neutrino baryonic resonance processes have ~30% errors

MINERvA, PRD100(2019)072005, EPJST230(2021)4243

RESonance



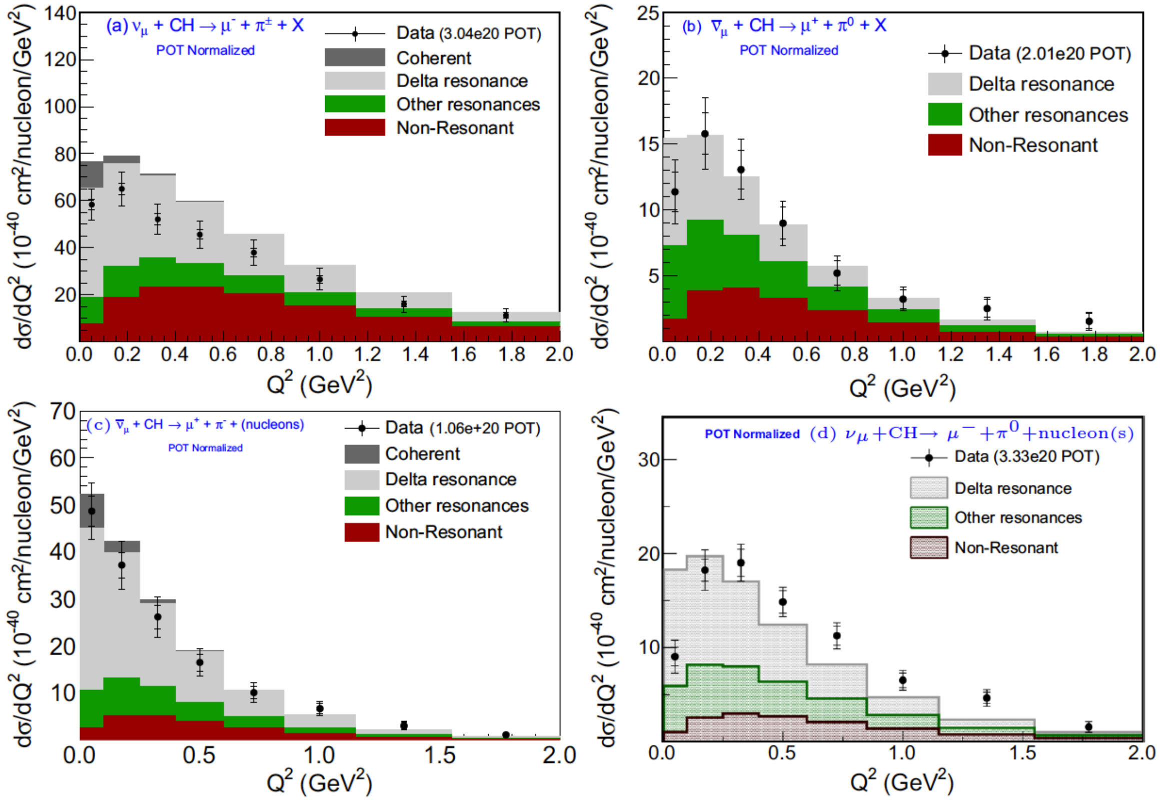
3. pion production global fit

MINERvA pion data

- It is extremely difficult to tune pion and/or FSI parameters to fit all pion data
  - $\nu_\mu CC\pi^\pm$, low Q2 suppression, over-predicted
  - $\nu_\mu CC\pi^0$, strong low Q2 suppression
  - $\bar{\nu}_\mu CC\pi^-$, no low Q2 suppression
  - $\bar{\nu}_\mu CC\pi^0$, low Q2 suppression, under-predicted

The study relies of available knobs in the simulation

It looks the simulation doesn't have good knobs to tune or missing



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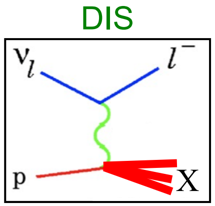
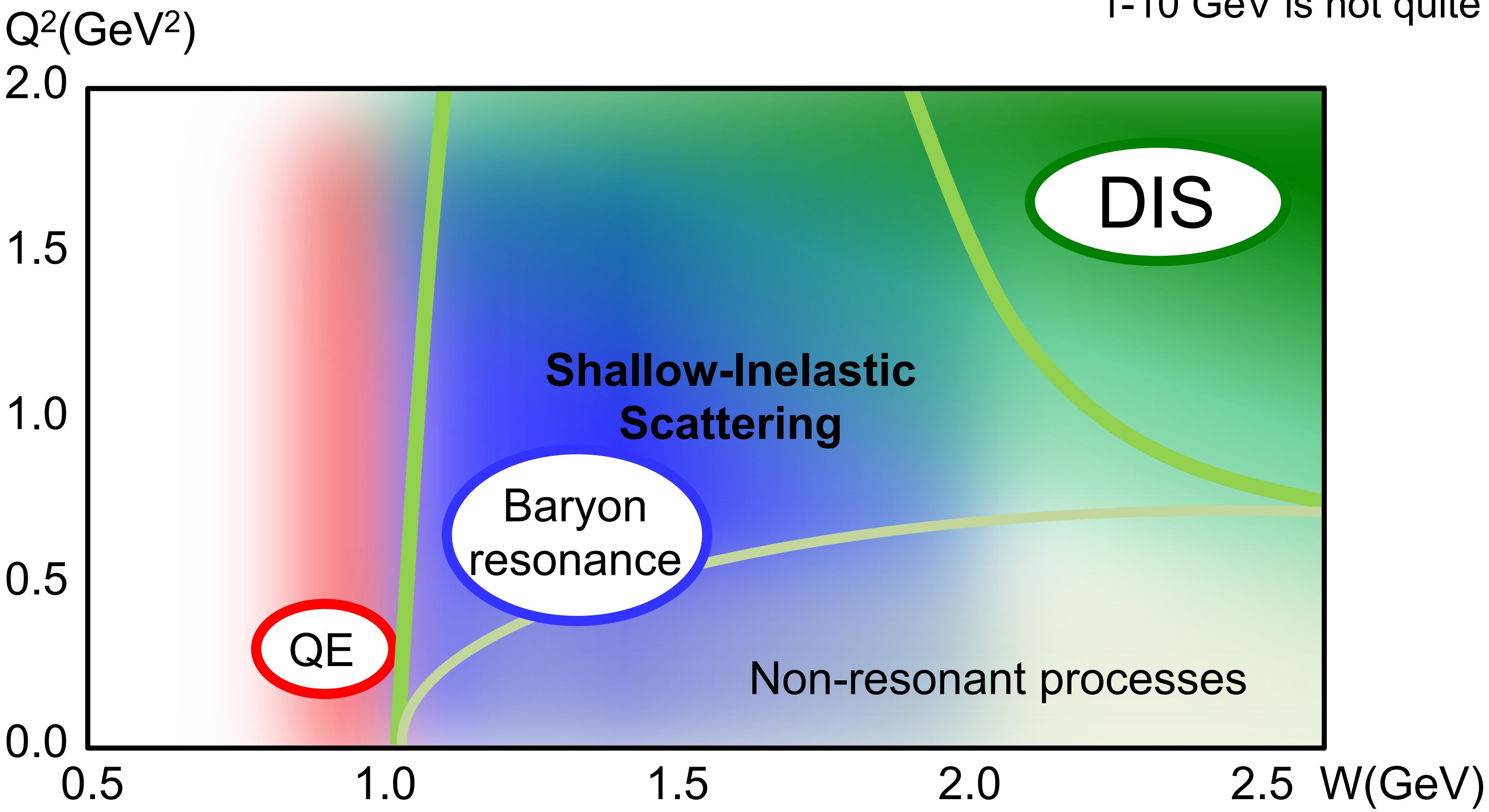
4. Shallow- and Deep-Inelastic Scattering (SIS and DIS)



Cross section

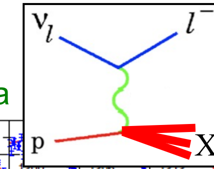
- Higher resonances and hadron dynamics
- Quark-Hadron duality (low Q^2, low W DIS)
- Nuclear dependent DIS

Neutrino experiment around 1-10 GeV is not quite DIS ye



Nakamura et al,Rep.Prog.Phys.80(2017)056301

4. Higher baryonic resonances



Cross section
- Higher resonances and hadron dynamics
- Quark-Hadron duality (low Q^2, low W DIS)
- Nuclear dependent DIS

DCC model
- Total amplitude is conserved
- Channels are coupled (πN, $\pi\pi N$, etc)
- 2 pion productions ~10% at 2 GeV

DCC model vs. electro-pionproduction data

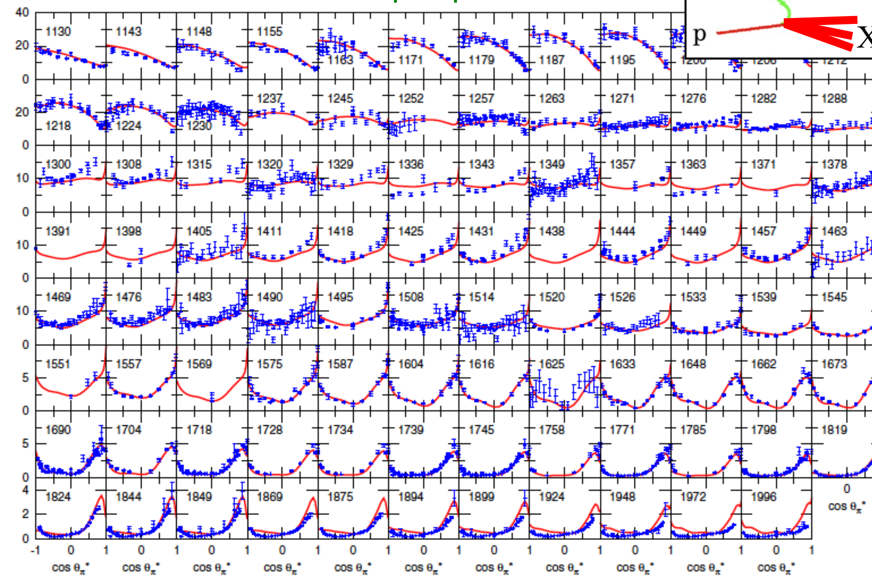
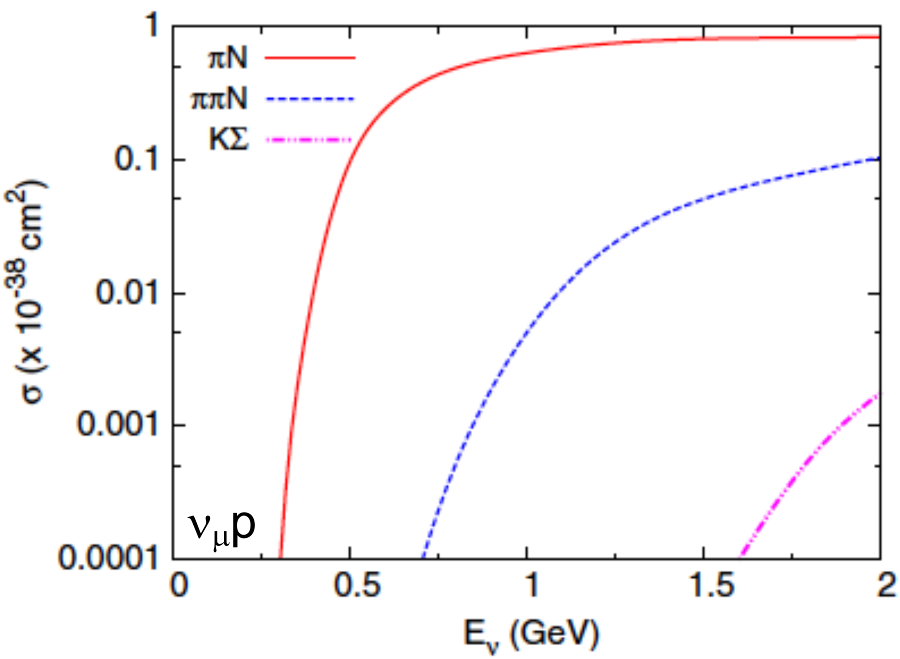
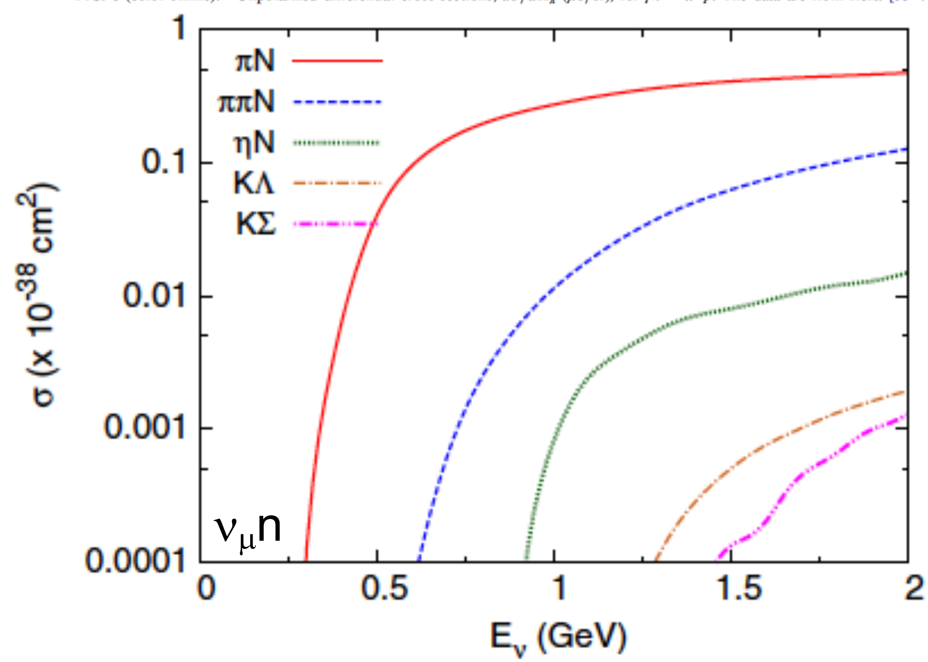


FIG. 8 (color online). Unpolarized differential cross sections, $d\sigma/d\Omega_\pi^*$ (μb/sr), for $\gamma n \to \pi^- p$. The data are from Refs. [55–78].

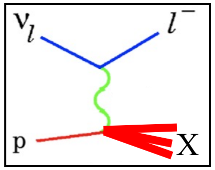




4. Quark-Hadron duality

Nachtmann variable $\xi = \dfrac{2x}{\left(1+\sqrt{1+\dfrac{4x^2M^2}{Q^2}}\right)}$

DIS



Cross section

- Higher resonances and hadron dynamics
- Quark-Hadron duality (low Q^2, low W DIS)
- Nuclear dependent DIS

GRV98 LO PDF + Bodek-Yang correction

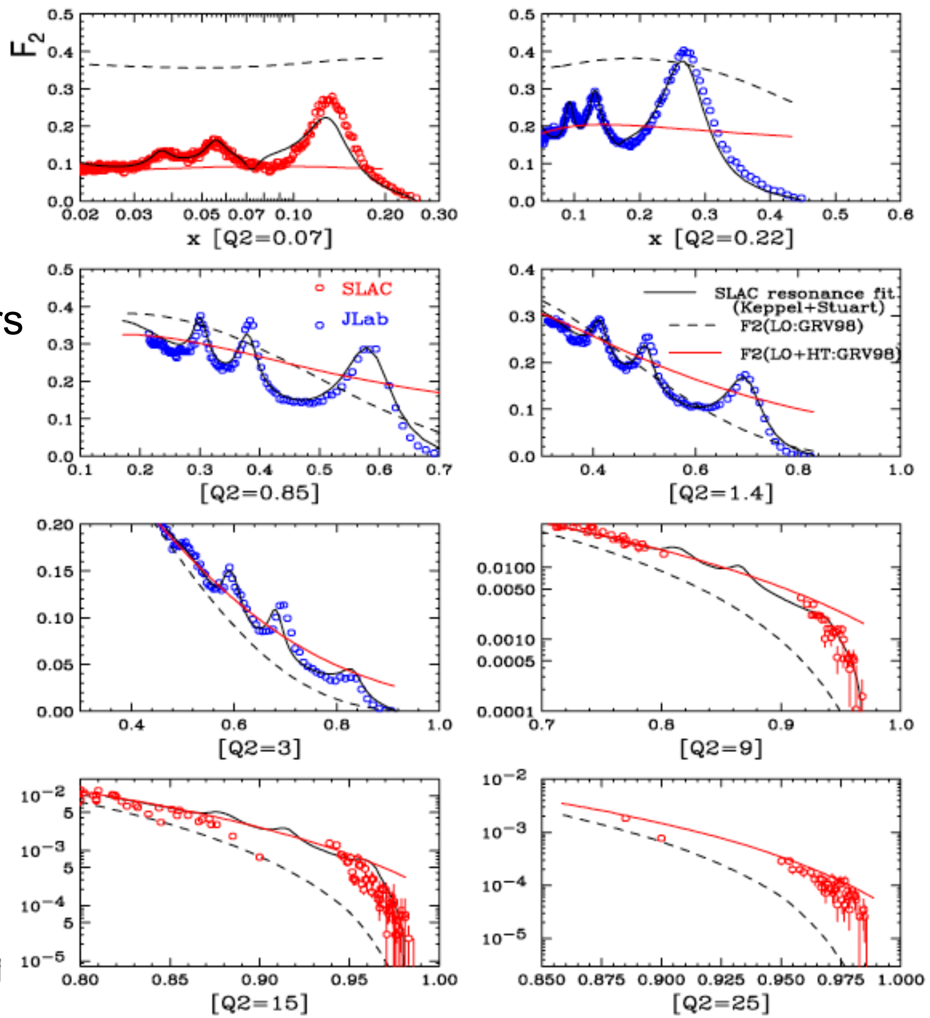
- GRV98 for low Q^2 DIS
- Bodek-Yang correction for QH-duality
- 20 years old, out-of-dated
- not sure how to implement systematic errors

$$\xi \rightarrow \xi_\omega = \frac{2x\left(1+\dfrac{M_f^2+B}{Q^2}\right)}{\left(1+\sqrt{1+\dfrac{4x^2M^2}{Q^2}}\right)+\dfrac{2Ax}{Q^2}}$$

$$K_{valence}(Q^2) = [1-G_D^2(Q^2)]\cdot\left(\frac{Q^2+C_{v2}}{Q^2+C_{v1}}\right)$$

$$K_{sea}(Q^2) = \frac{Q^2}{Q^2+C_{s1}}$$

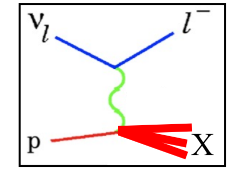
Proton F2 function GRV98-BY correction vs. data



HKN,PRC76(2007)065207, EPS,JHEP04(2009)065, FSSZ,PRD85(2012)074028
nCTEQ, PRD80(2009)094004

4. Nuclear dependent DIS

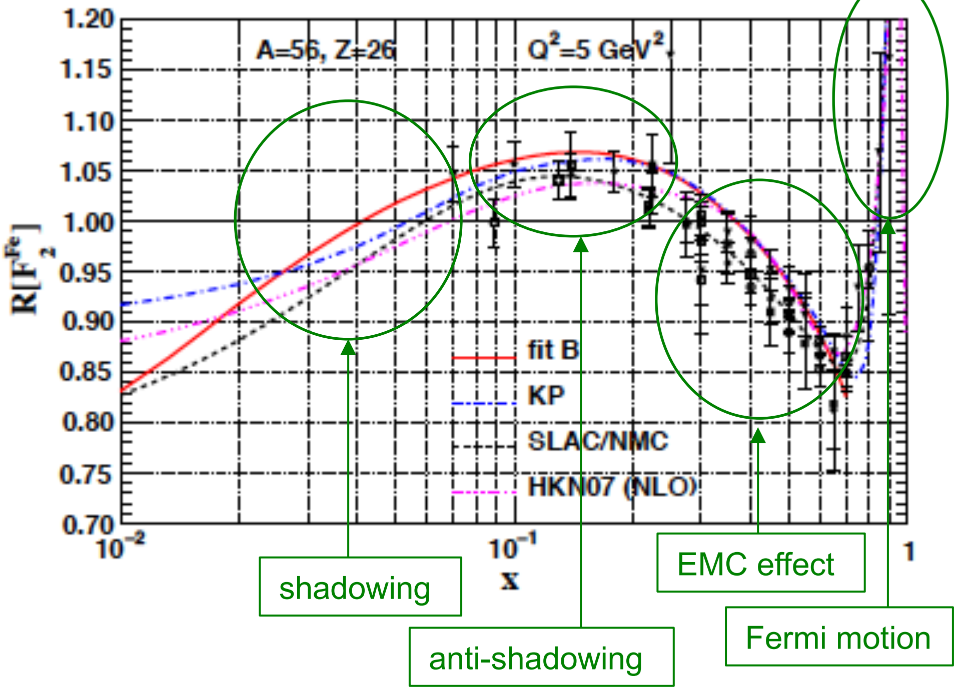
DIS



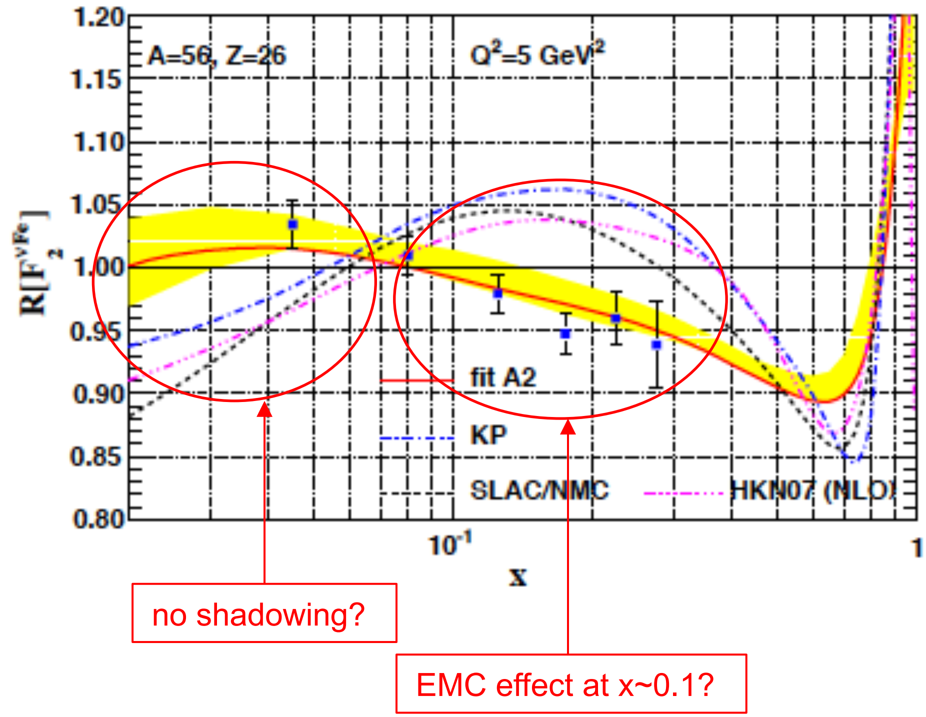
Cross section
- Higher resonances and hadron dynamics
- Quark-Hadron duality (low Q^2, low W DIS)
- Nuclear dependent DIS

Nuclear PDF
- Shadowing, EMC effect, Fermi motion
- Likely due to nucleon dynamics in nucleus
- Various models describe charged lepton data
- Neutrino data look very different

$\ell^{\pm}$-Fe nuclear correction factor



ν-Fe nuclear correction factor



Conclusion

ν-N scattering : important reactions for long baseline neutrino oscillation experiment
(T2K, NOvA, DUNE, Hyper-Kamiokande, etc)

CCQE: charged-current quasi-elastic, around 1 GeV
RES: baryonic resonance, around 2 GeV
DIS: deep inelastic scattering, 3 GeV to higher

Nuclear physics sucks
- Fermi motion: nucleon motion smears kinematic reconstruction
- Pauli blocking: It limits low momentum transfer reaction
- Nuclear shell structure: separation energy (missing energy) for different nucleons
- Final state interaction: RES looks like CCQE, DIS looks like RES, etc
- Nucleon correlation: Physics between ν-N and ν-A interaction
- Quark-Hadron duality: Physics between ν-q and ν-N interaction
- Nuclear dependent PDF: Physics between ν-q and ν-A interaction

Currently, ~30% error is acceptable for many processes

References (books)

Quarks and Leptons (Halzen and Martin)
- show many calculations
- solutions for all exercises

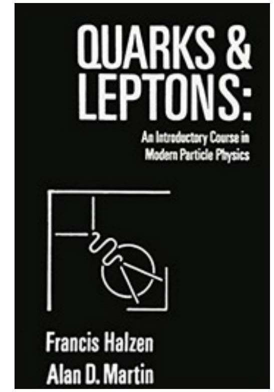
Weak interactions of Leptons and Quarks (Commins and Bucksbaum)
- show details of weak interaction calculations
- too many typos

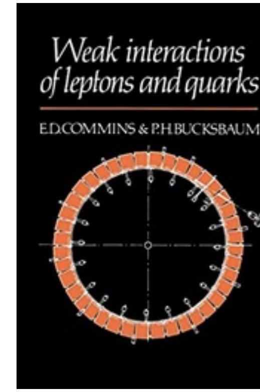
Physics of Neutrinos (Fukugita and Yanagida)
- very intense
- from solar neutrinos to SUSY

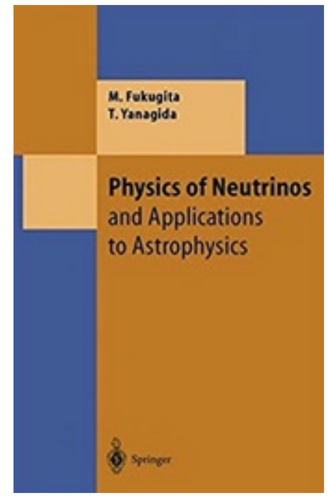
Neutrino astrophysics (Bahcall)
- good book to read

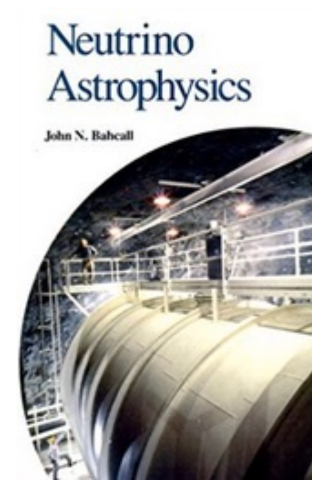
Foundation of Nuclear and Particle Physics (2017)
- Authors: Donnelly, Formaggio, Holstein, Milner, Surrow
- buy if your PhD thesis topic is about neutrino
cross section measurements in T2K, NOvA, SBN, etc

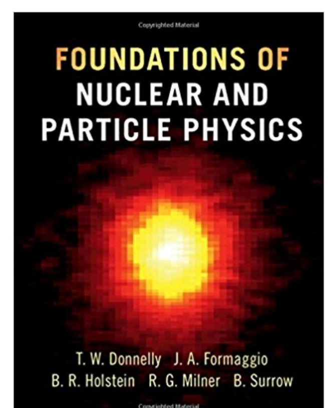
The Physics of Neutrino interactions (2020)
- Authors: Sajjad Athar, Singh
- The newest book in this kind (970 pages!)

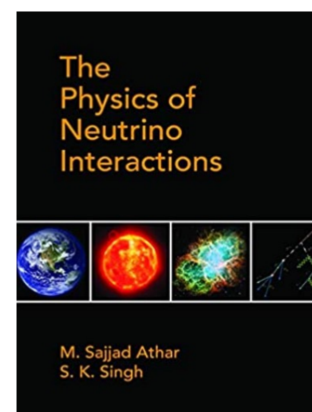












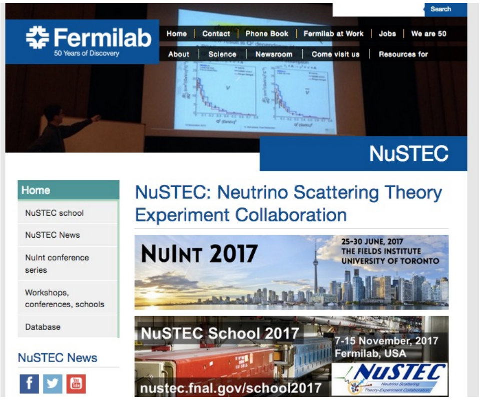
References (papers)

"From eV to EeV: Neutrino cross sections across energy scales"
- Authors: Formaggio and Zeller (MicroBooNE spokesperson)
- Rev.Mod.Phys.84(2012)1307, https://arxiv.org/abs/1305.7513
- very good summary of neutrino cross sections

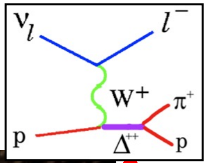
"Neutrino-Nucleus Cross Sections for Oscillation Experiments"
- Authors: Katori (me) and Martini (Martini model)
- J.Phys. G45 (2018) no.1, 013001, https://arxiv.org/abs/1611.07770
- A review both theoretical and experimental views

"NuSTEC White Paper: Status and challenges of neutrino–nucleus scattering"
- NuSTEC (Neutrino Scattering Theory-Experiment Collaboration)
- Prog.Part.Nucl.Phys. 100 (2018) 1-68, https://arxiv.org/abs/1706.03621
- Cover all open issues in the community

"NuSTEC News"
- http://nustec.fnal.gov/
- subscribe mailing list, "like" facebook page, use #nuxsec

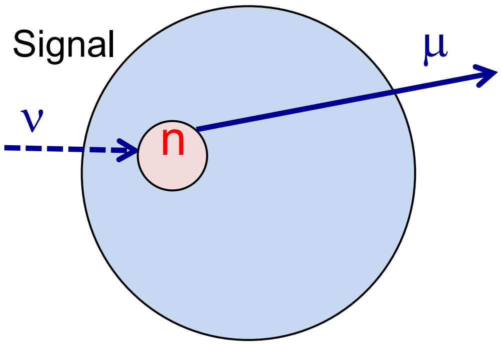


Backup



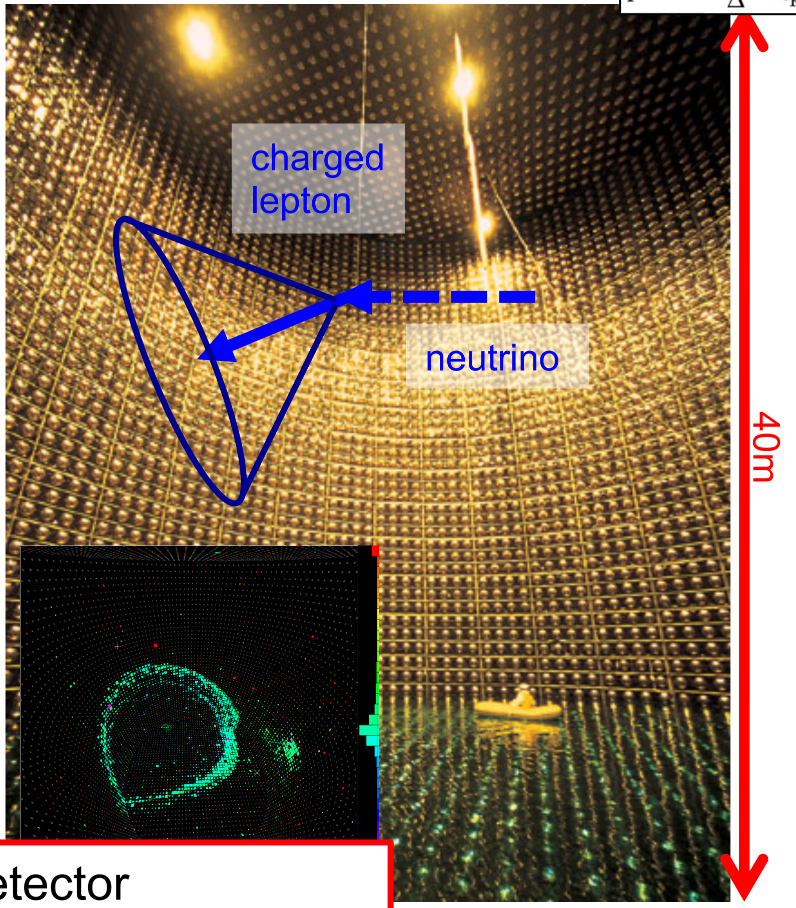
3. non-QE background (resonance pion production)

non-QE background → shift spectrum



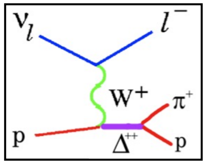
Neutrino energy is reconstructed from the observed lepton kinematics
"QE assumption"

1. assuming neutron at rest
2. assuming interaction is CCQE



Typical neutrino oscillation detector
- Big and dense, to maximize interaction rate
- Coarsely instrumented, to minimize cost

(not great detector to measure hadrons)



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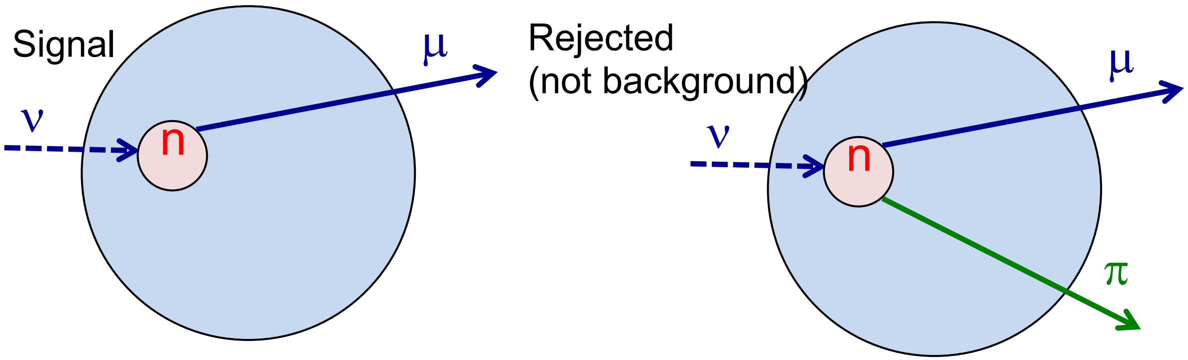

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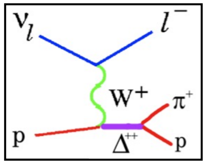
RESonance



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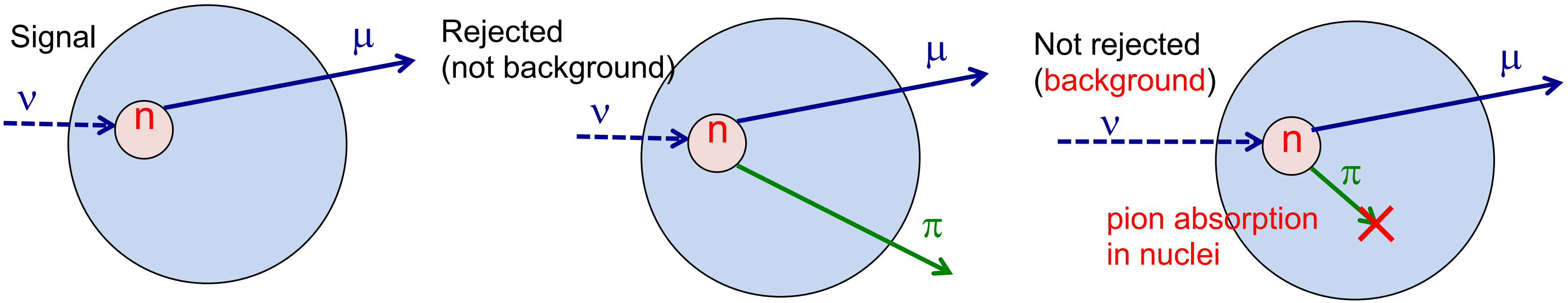

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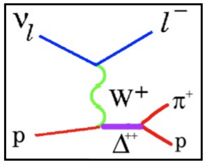
RESonance



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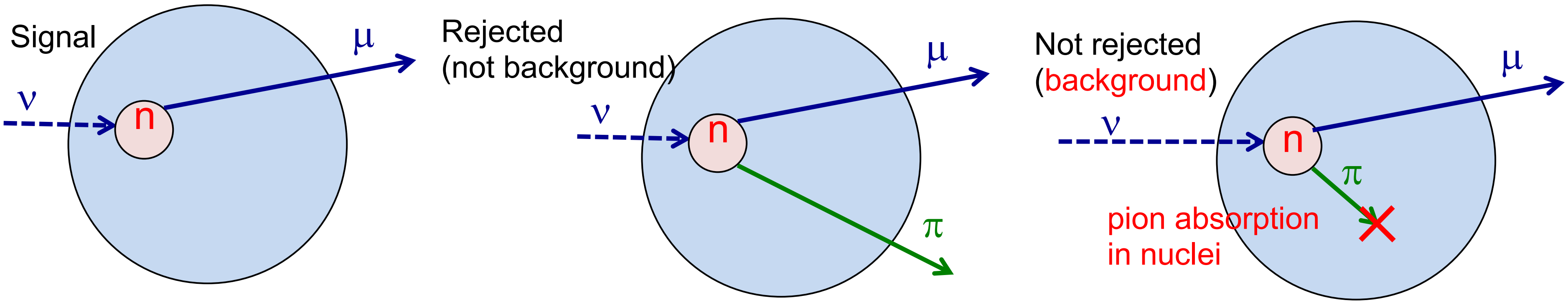
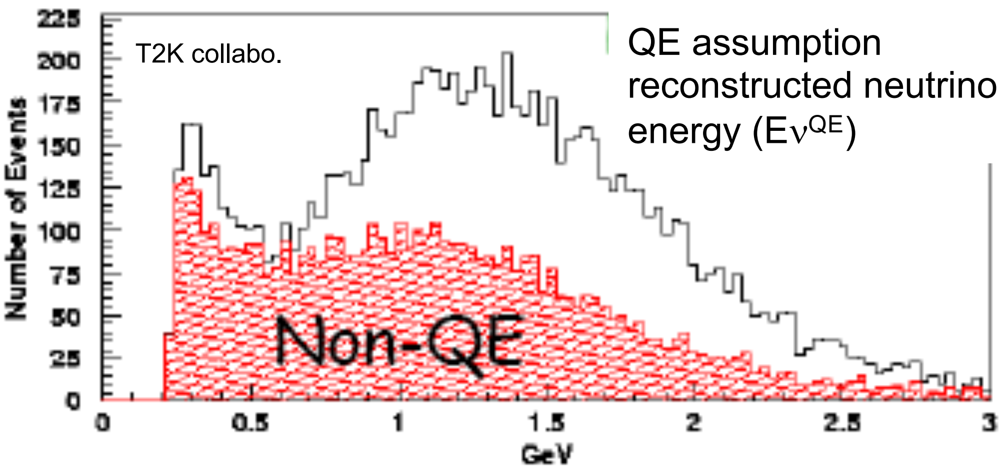
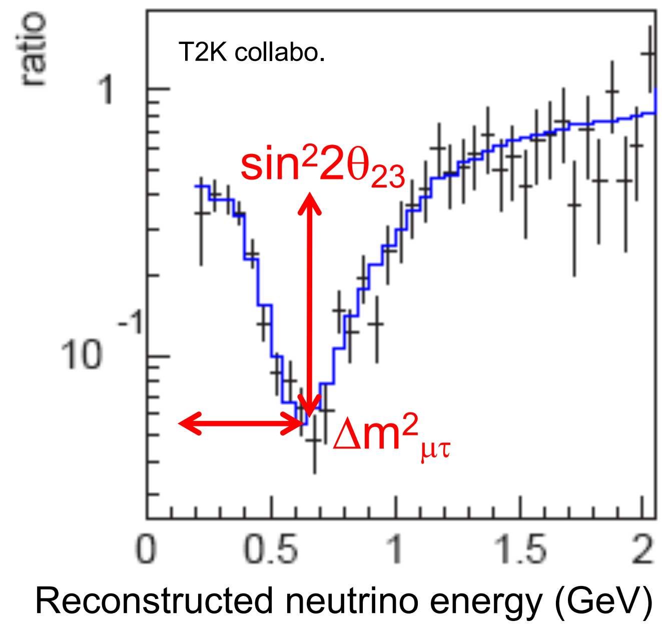




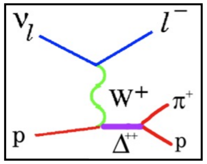
QE assumption reconstructed neutrino energy ($E\nu^{QE}$)



katori@fnal.gov

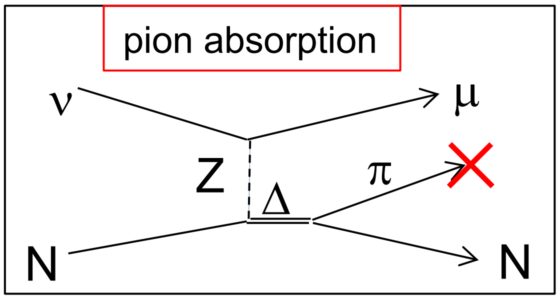
Coloma et al,PRL111(2013)221802
Mosel et al,PRL112(2014)151802

RESonance

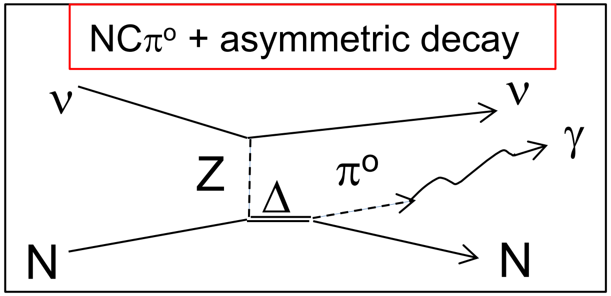


3. non-QE background (resonance pion production)

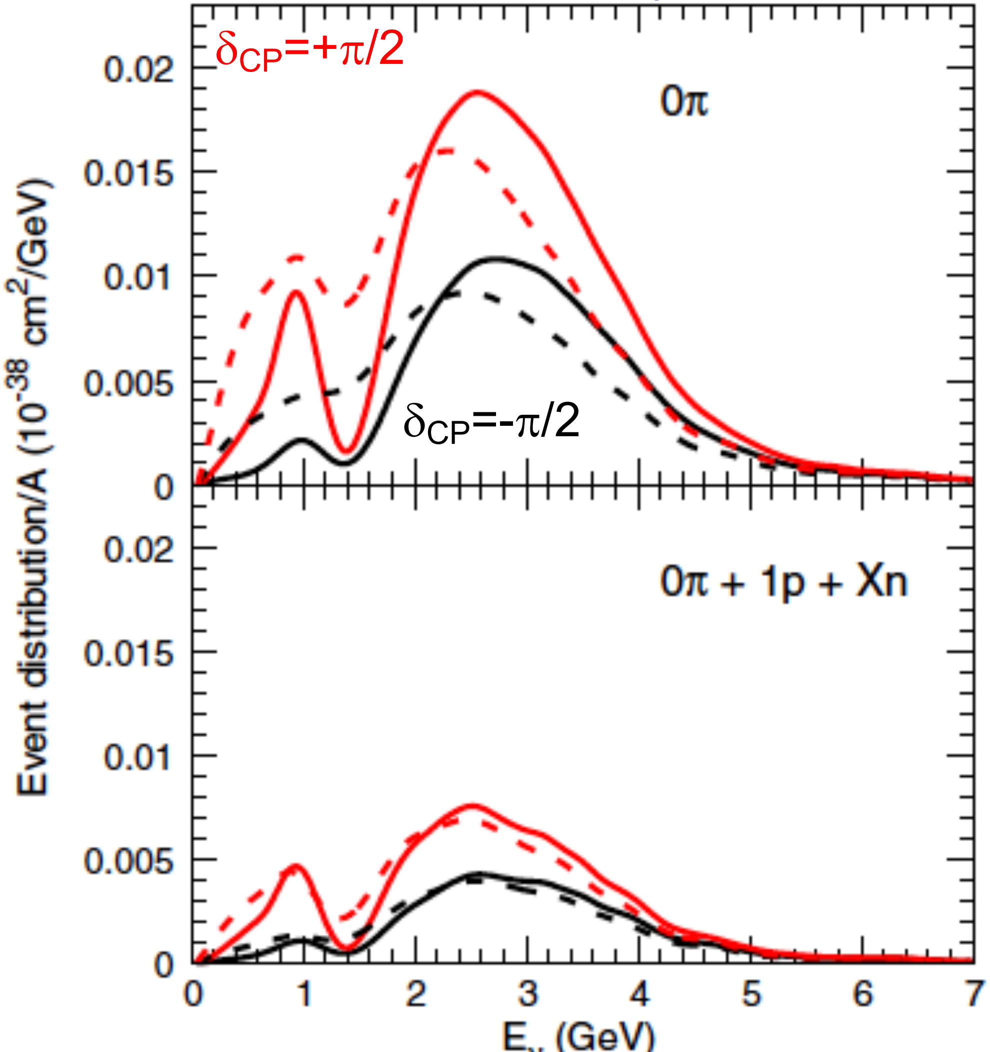
Pion production for ν_μ disappearance search
- Source of mis-reconstruction of neutrino energy



Neutral pion production in ν_e appearance search
- Source of misID of electron



DUNE true vs. reconstructed Eν spectrum



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Understanding of neutrino baryonic resonance meson production is important for oscillation experiments