Neutrino Interaction Physics

Lecture 1: Introduction of neutrino interactions
1. Overview
2. Neutrino lepton scattering
3. Neutrino quark scattering (DIS)
4. Neutrino nucleus reactions

Lecture 2: Neutrino interactions for long baseline oscillation experiments
1. Overview
2. CCQE interaction
3. Baryonic resonances
4. Shallow inelastic scattering (SIS)

Teppei Katori
King's College London
Oct. 13, 2020

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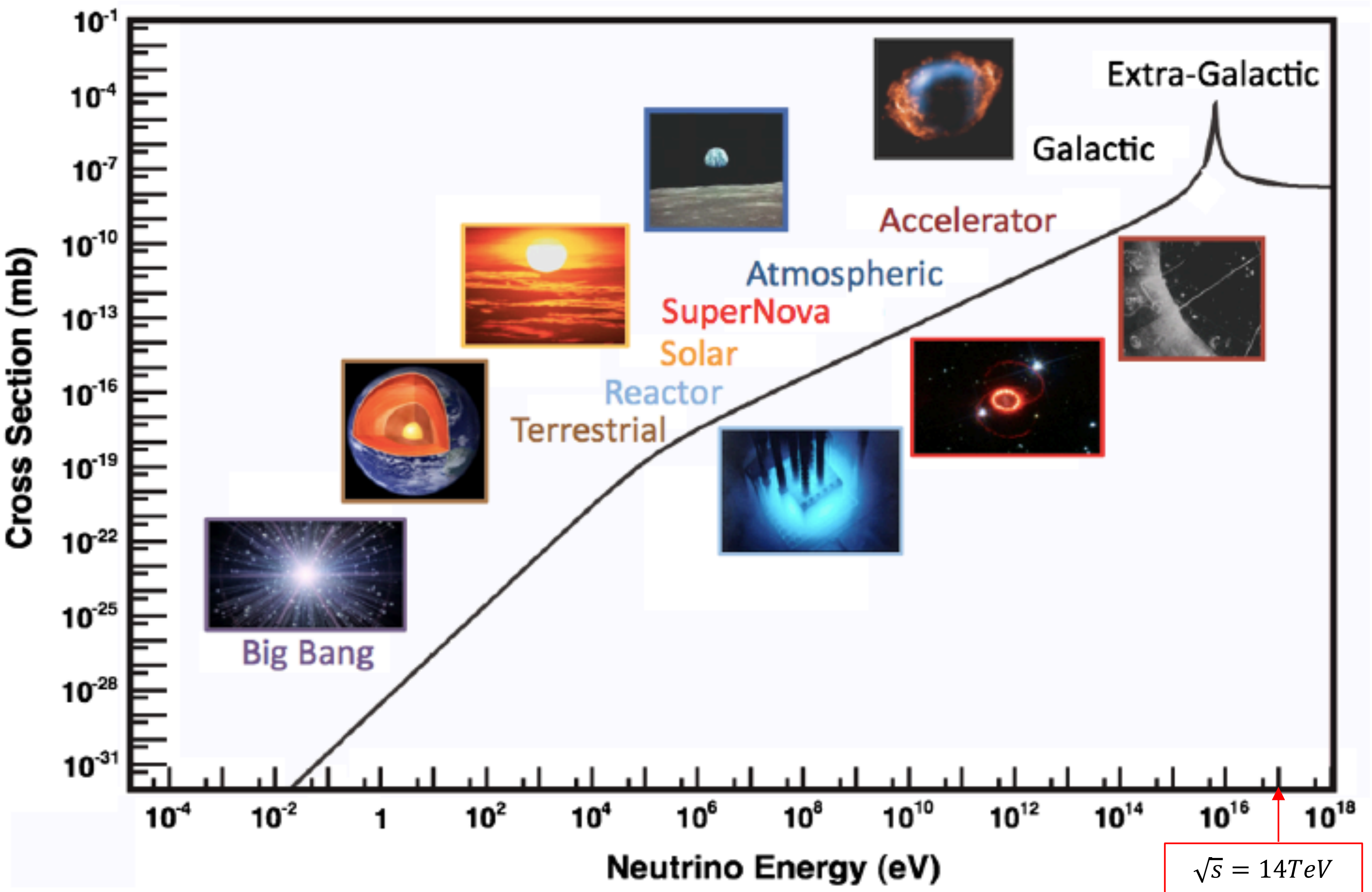
Formaggio and Zeller, Rev.Mod.Phys.,84 (2012) 1307

1. From eV to EeV: Neutrino cross sections across energy scales

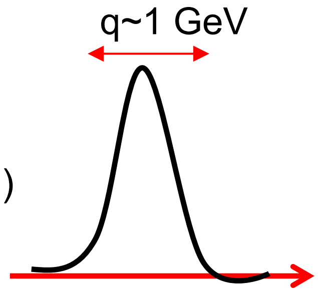


$\sqrt{s} = 14TeV$

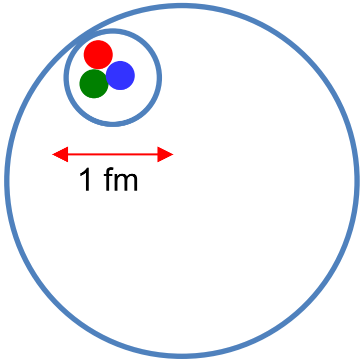
1. Scattering measurements

Size of wave packet ~ momentum transfer (~energy)

$\hbar c = 197\ MeV \cdot fm \rightarrow$ 200 MeV ~ 1 fm (size of nucleon)



~1 GeV neutrino beam
(T2K, NOvA, HyperK DUNE)



1. Scattering measurements

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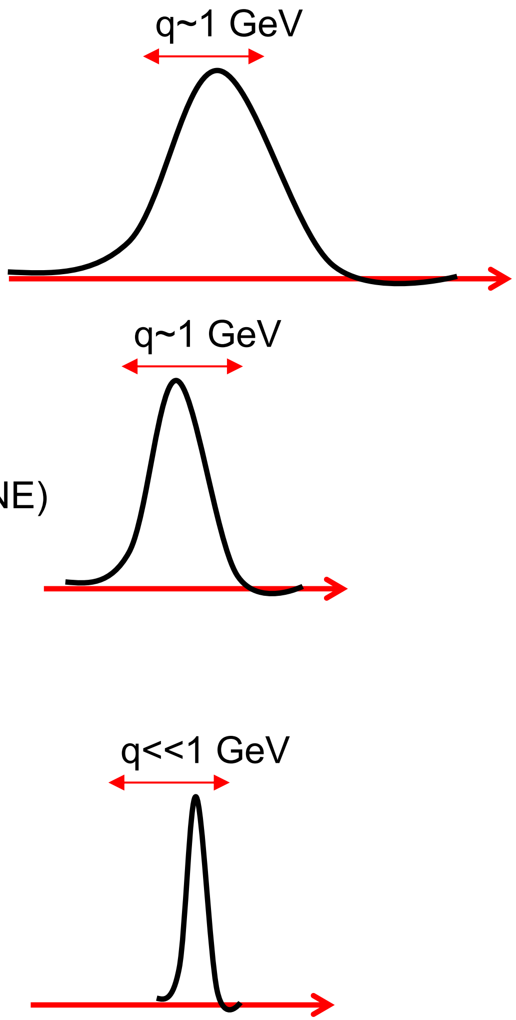
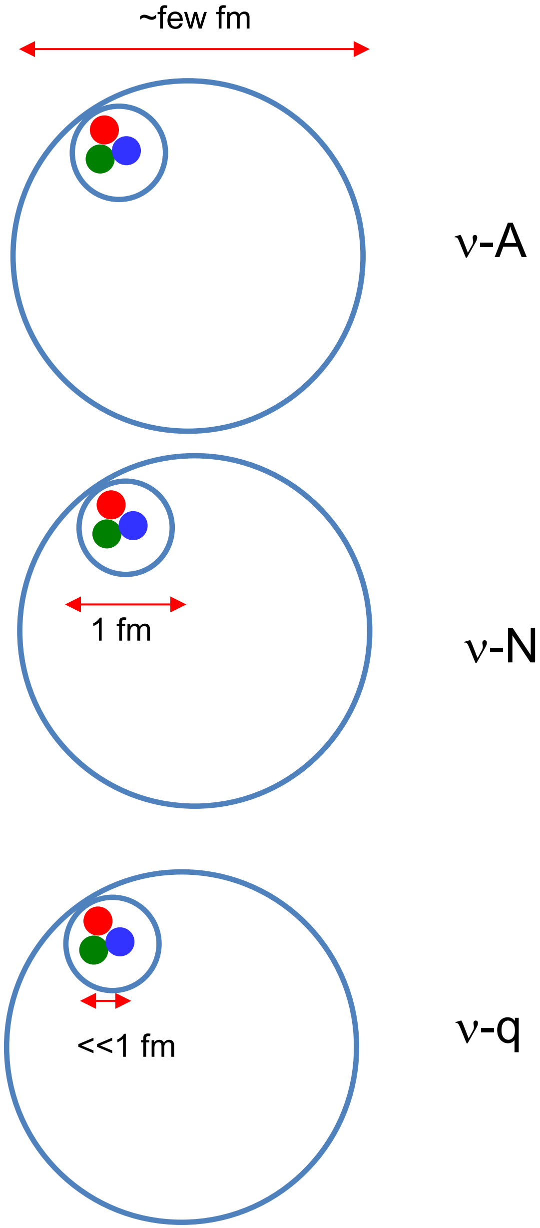
$\hbar c = 197\ MeV \cdot fm \rightarrow$ 200 MeV ~ 1 fm (size of nucleon)



<<1 GeV neutrino beam
(solar neutrinos, etc)

~1 GeV neutrino beam
(T2K, NOvA, HyperK DUNE)

>>1 GeV neutrino beam
(Tevatron, LHC)



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1. Scattering measurements

Size of wave packet ~ momentum transfer (~energy)

$\hbar c = 197\ MeV \cdot fm \rightarrow$ 200 MeV ~ 1 fm (size of nucleon)

Lecture 1: Introduction of neutrino interactions
1. Overview
2. Neutrino lepton scattering (Standard Model)
3. Neutrino quark scattering (ν-q scattering)
4. Neutrino nucleus reactions (ν-A scattering)

Lecture 2: Neutrino interactions for long baseline oscillation experiments (ν-N scattering)
1. Overview
2. CCQE interaction
3. Baryonic resonances
4. Shallow inelastic scattering

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Alexander 2016 J.Phys.:Conf. Ser. 718(2016)062076

2. Neutrino-electron scattering

Neutrino – electron differential cross section

T=recoil electron kinetic energy
E=neutrino energy

$$\frac{d\sigma}{dT} = \frac{2G_F^2 m_e}{\pi}\left[c_L^2 + c_R^2\left(\frac{E-T}{E}\right)^2 - C_L C_R \frac{m_e T}{E^2}\right]$$

Neutrino – electron differential cross section with neutrino magnetic moment ($\mu_\nu < 3\text{x}10^{-11}\ \mu_B$)

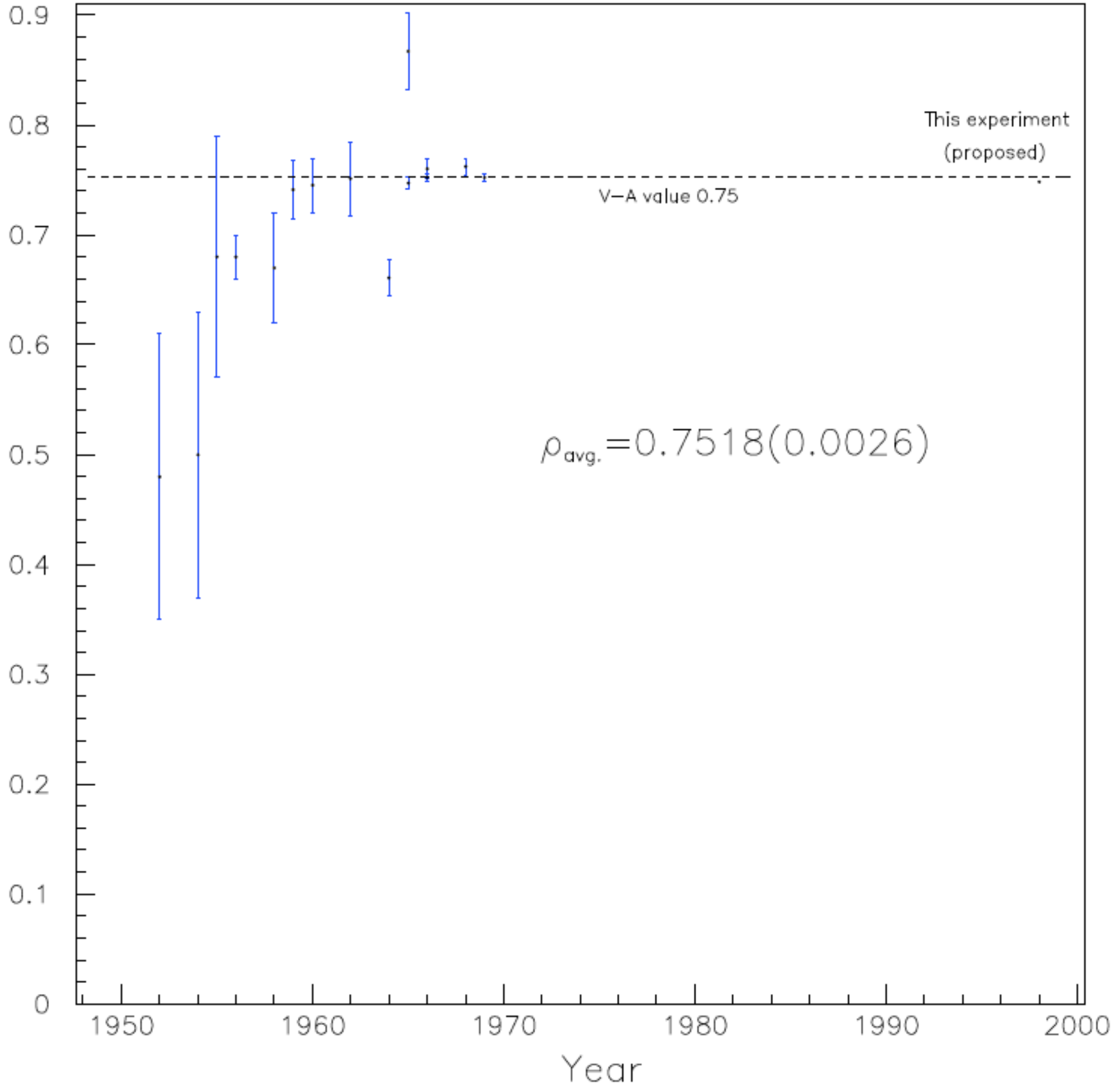
$$\frac{d\sigma}{dT} = \frac{2G_F^2 m_e}{\pi}\left[c_L^2 + c_R^2\left(\frac{E-T}{E}\right)^2 - C_L C_R \frac{m_e T}{E^2}\right] + \frac{\pi\alpha\mu_\nu^2}{m_e^2}\left(\frac{1}{T} - \frac{1}{E}\right)$$

Lepton-only process (pure Standard Model) is often used to test new physics

	C_L	C_R
ν_e - e^-	$½+\sin^2\theta_W$	$\sin^2\theta_W$
$\overline{\nu_e}$ - e^-	$\sin^2\theta_W$	$½+\sin^2\theta_W$
ν_μ - e^-	$-½+\sin^2\theta_W$	$\sin^2\theta_W$
$\overline{\nu_\mu}$ - e^-	$\sin^2\theta_W$	$-½+\sin^2\theta_W$

Lee, PhD thesis, Virginia Tech (2001)

2. Time dependence of muon decay Michel parameter ρ



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3. Neutrino-DIS cross section

Neutrino – single d-quark cross section

$$\frac{d\sigma}{dy}(\nu d \rightarrow \mu u) = \frac{G_F^2 xs}{\pi}$$

Neutrino – d-quark cross section

$$\frac{d\sigma}{dy}(\nu d \rightarrow \mu u) = \int_0^1 \frac{G_F^2 xs}{\pi} d(x)dx$$

Neutrino-nucleon DIS cross section

$$\frac{d\sigma}{dy}(\nu N \rightarrow \mu X) = \int_0^1 \frac{G_F^2 xs}{\pi} [(d(x) + s(x) \ldots) + [\bar{u}(x) + \bar{c}(x) \ldots](1-y)^2]dx$$

Neutrino-nucleus DIS cross section with isoscalar assumption

$$\frac{d\sigma}{dy}(\nu A \rightarrow \mu X) = A\int_0^1 \frac{G_F^2 xs}{\pi} [Q(x) + \bar{Q}(x)(1-y)^2]dx$$

$$u^p(x) + u^n(x) = d^n(x) + d^p(x) = u(x) + d(x) \equiv Q(x)$$
$$\bar{u}^p(x) + \bar{u}^n(x) = \bar{u}^n(x) + \bar{u}^p(x) = \bar{u}(x) + \bar{d}(x) \equiv \bar{Q}(x)$$

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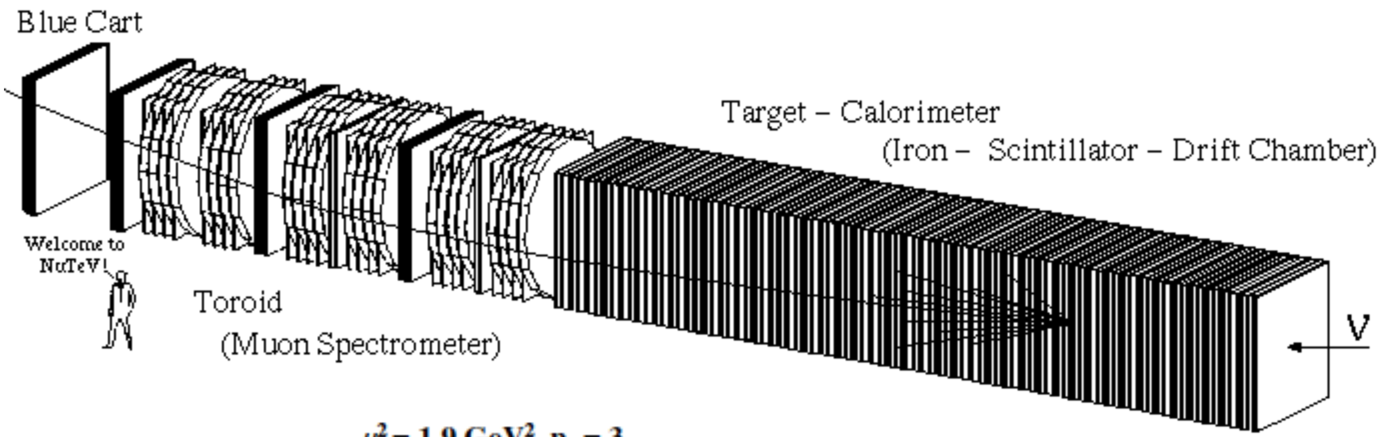
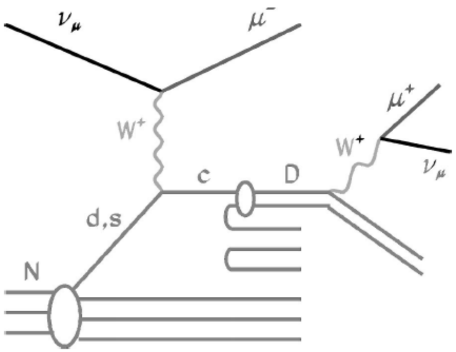
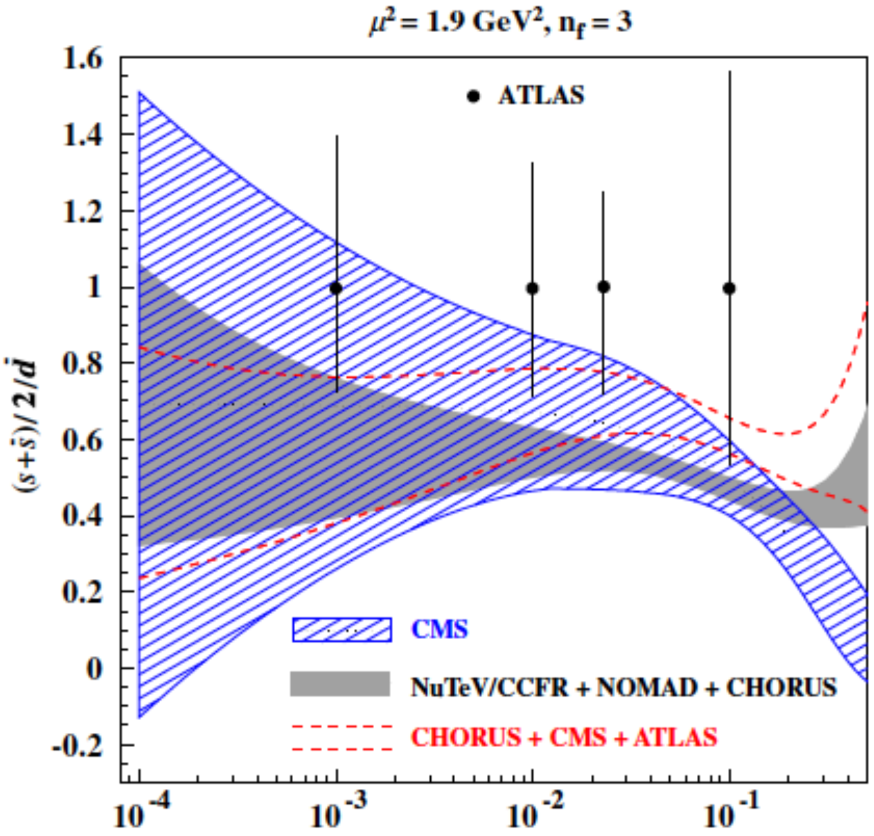
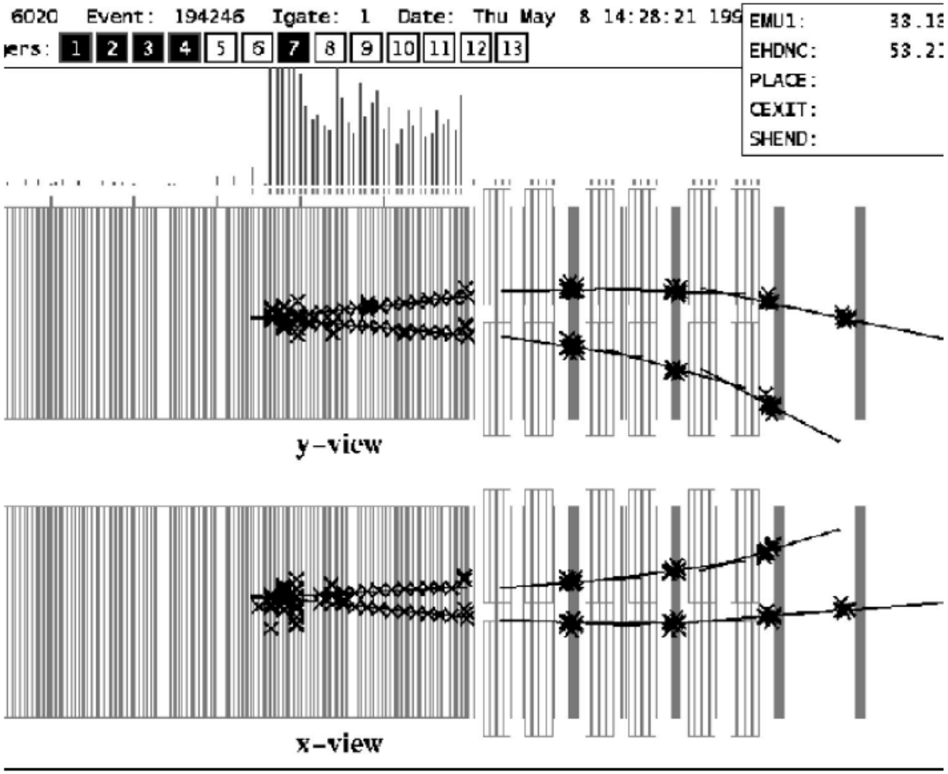
NuTeV, PRD64(2001)112006
Alekhin et al.,PRD91(2015)094002

3. Di-muon production



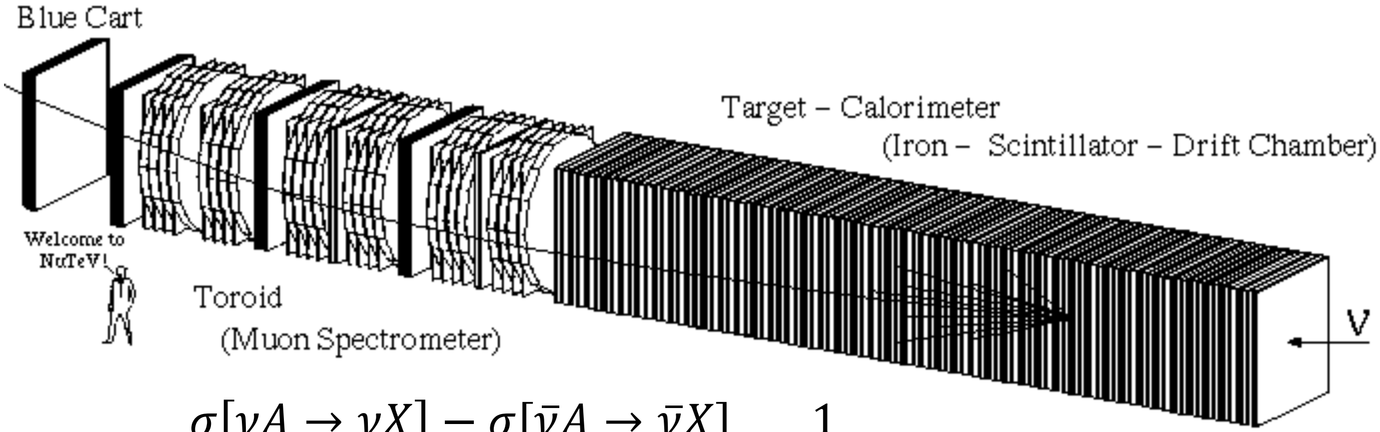






NuTeV, PRL88(2002)091802

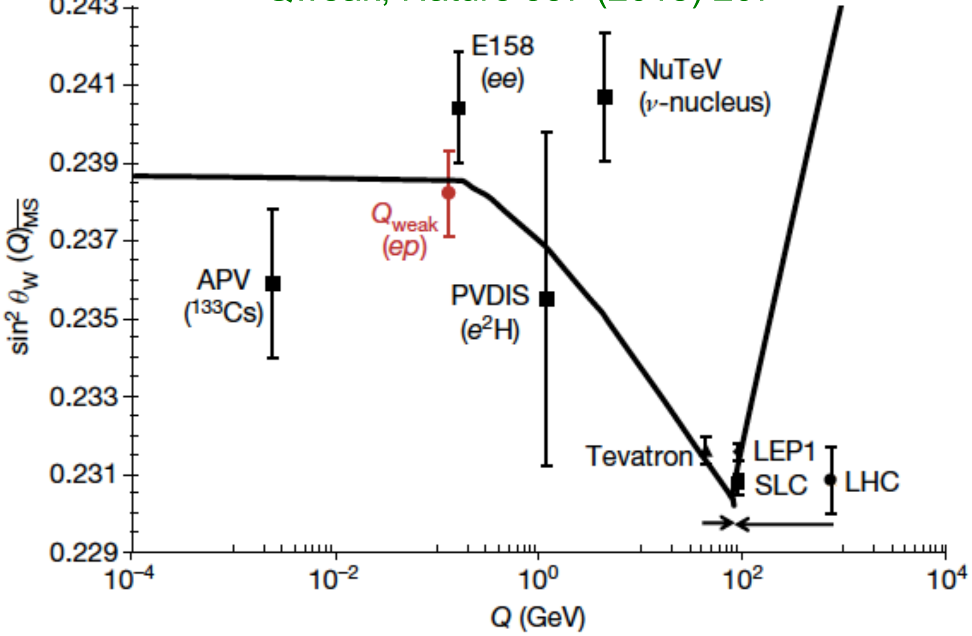
3. Paschos-Wolfenstein ratio and NuTeV anomaly



Manny Paschos
(Dortmund)

$$R_{PW} = \frac{\sigma[\nu A \to \nu X] - \sigma[\bar{\nu} A \to \bar{\nu} X]}{\sigma[\nu A \to \mu X] - \sigma[\bar{\nu} A \to \mu^{+} X]} = \frac{1}{2} - sin^{2}\theta_{W}$$

Qweak, Nature 557 (2018) 207



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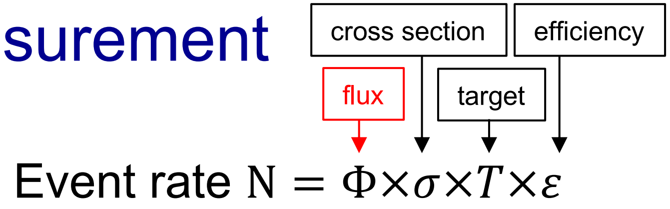
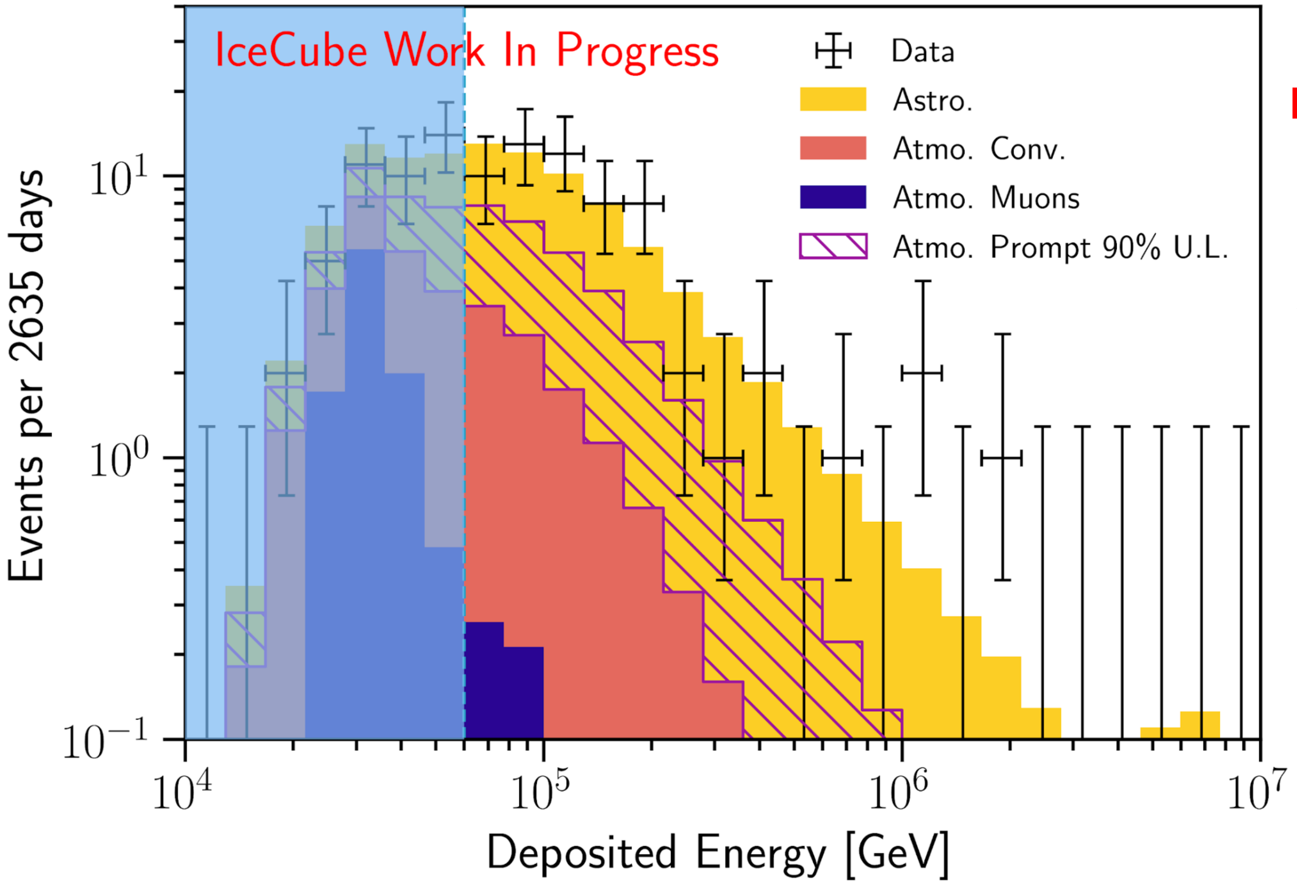
Farrag et al, ArXiv:1906.09623

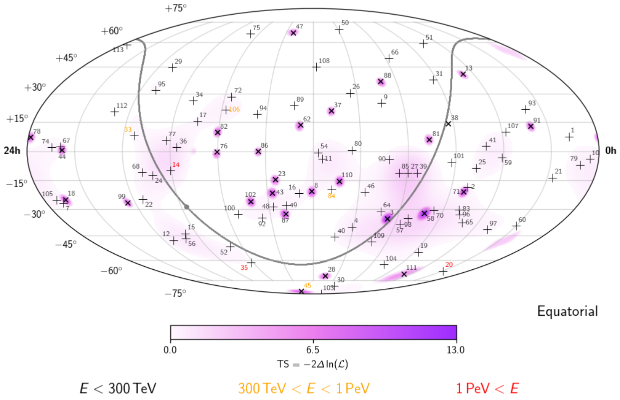
3. Astrophysical high-energy neutrino measurement



Event rate N $= \Phi \times \sigma \times T \times \varepsilon$

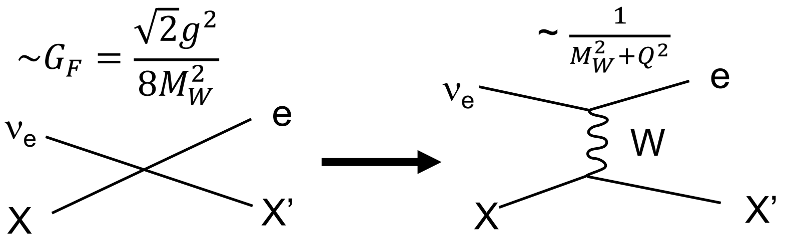


IceCube preliminary 2018



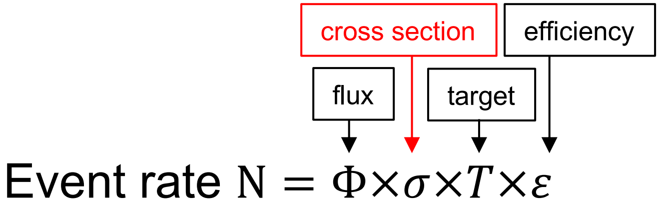
IceCube, Nature551(2017)596

3. Neutrino DIS saturation



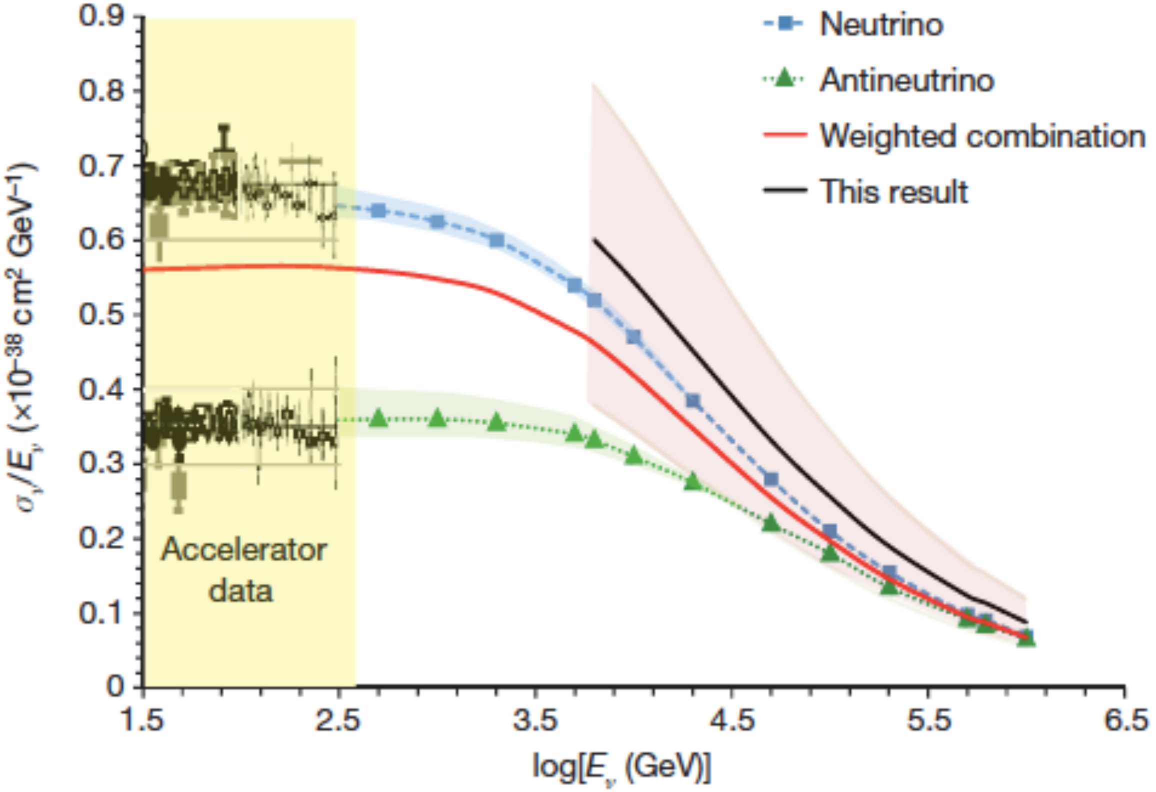
$$\sim G_F = \frac{\sqrt{2}g^2}{8M_W^2}$$

$$\sim \frac{1}{M_W^2+Q^2}$$



cross section
efficiency
flux
target

Event rate $N = \Phi\times\sigma\times T\times\varepsilon$



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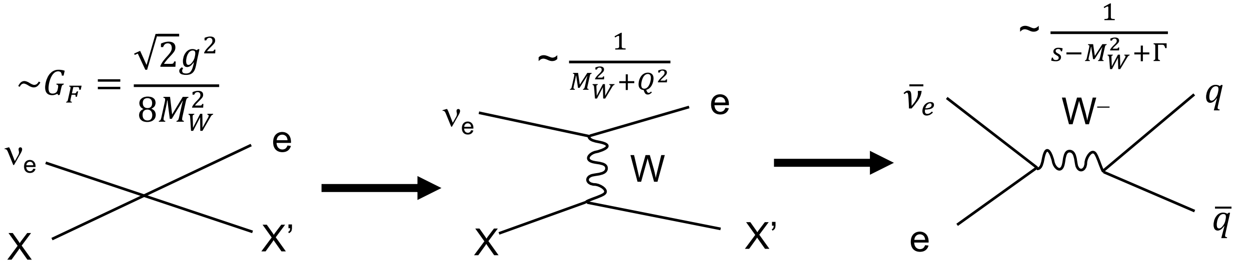
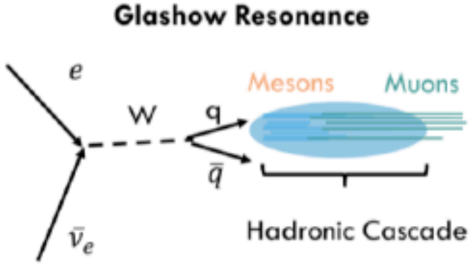
Taboada, Neutrino 2018, IceCube, arXiv:1710.01191

IceCube preliminary

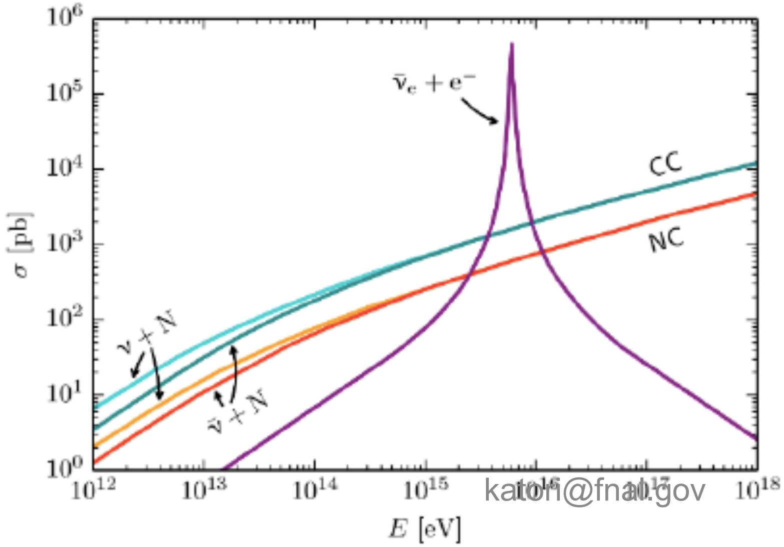
3. Glashow resonance



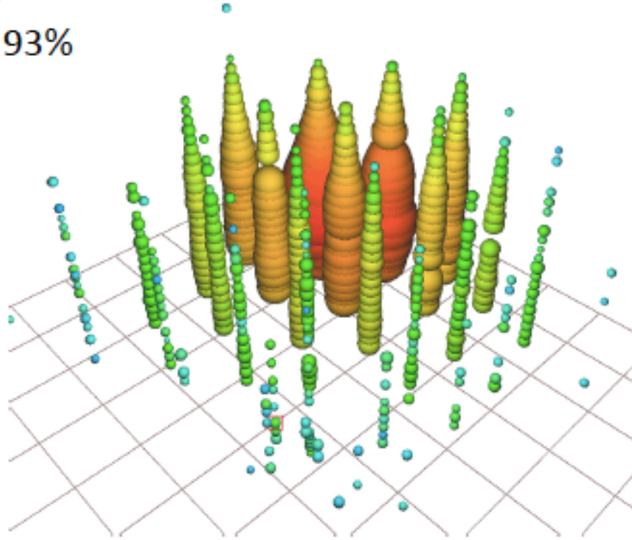
A 5.9 PeV event in IceCube



Resonance: E_ν = 6.3 PeV
Typical visible energy is 93%



Work in progress



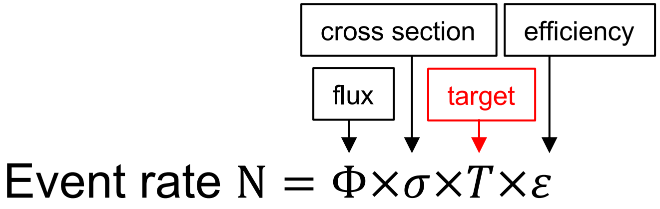
Event identified in a partially-contained PeV search (PEPE)

Deposited energy: 5.9±0.18 PeV (stat only)

ICRC 2017 arXiv:1710.01191

Dziewonski, Anderson (PREM), Phys. Earth Planet.Inter.25,(1981)297
Donini, Palomares-Ruiz, Salvado, Nature Physics 15(2019)37

3. Earth tomography

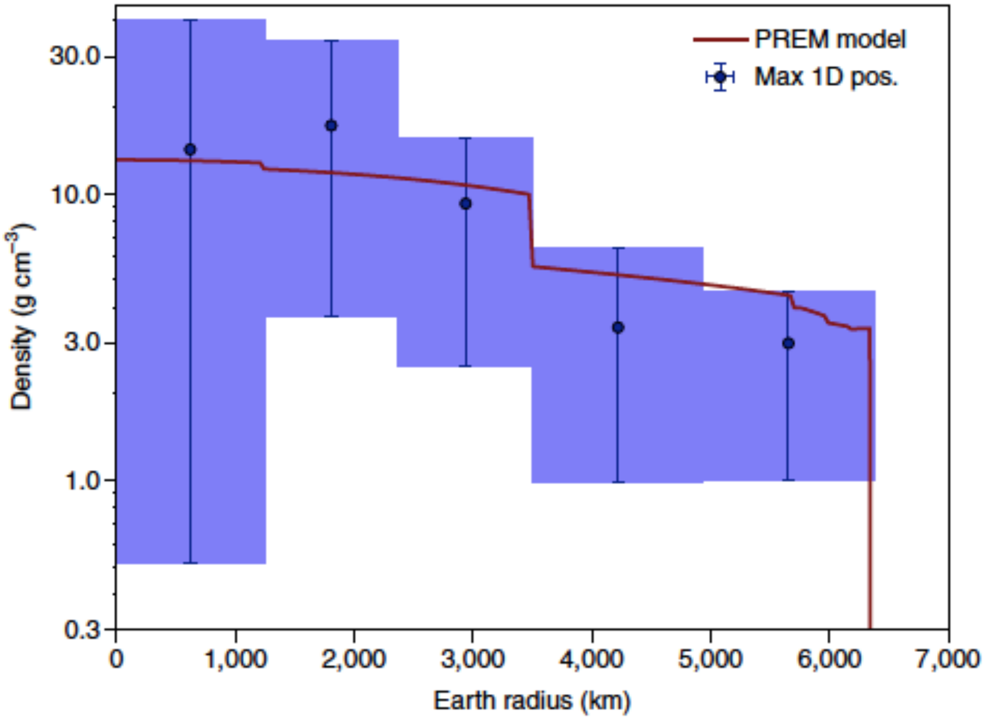


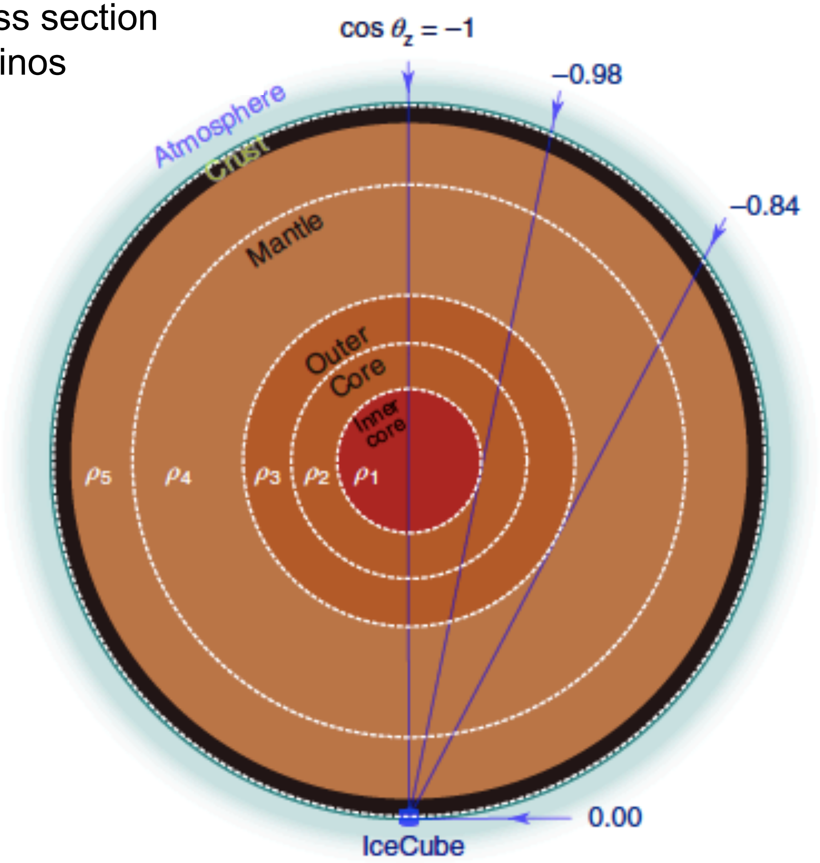
cross section
efficiency
flux
target

$$\text{Event rate N} = \Phi\times\sigma\times T\times\varepsilon$$

Earth absorption for Earth density measurement

- PREM (Preliminary reference Earth model)
- Standard earth density model used by T2K, NOvA, etc
- Earth density profile is extracted by assuming flux and cross section
- Measure Earth moment of inertia and Earth mass by neutrinos





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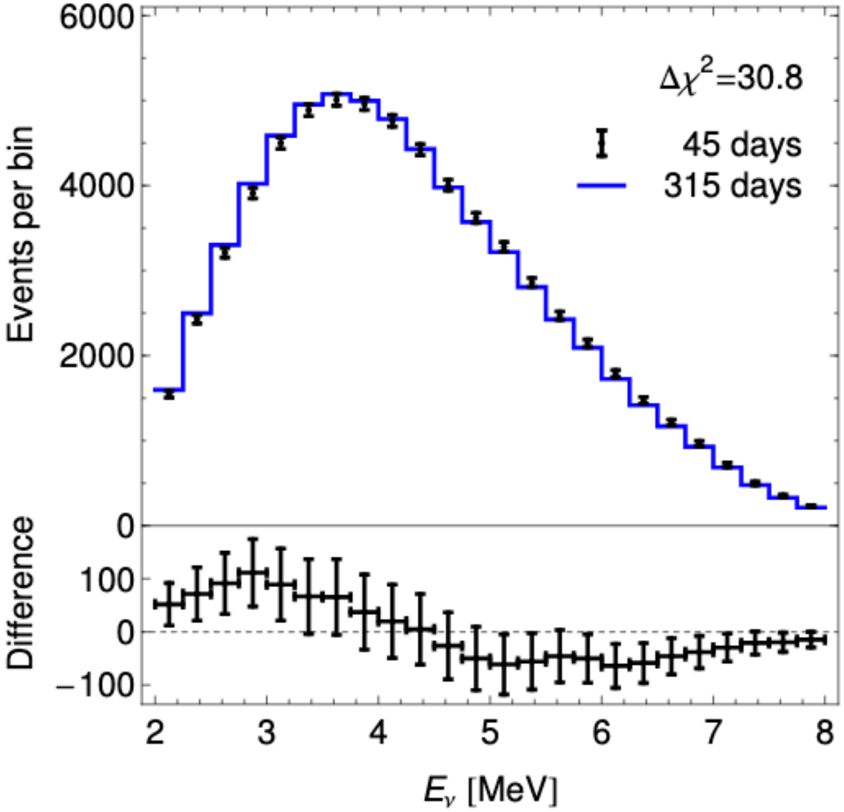
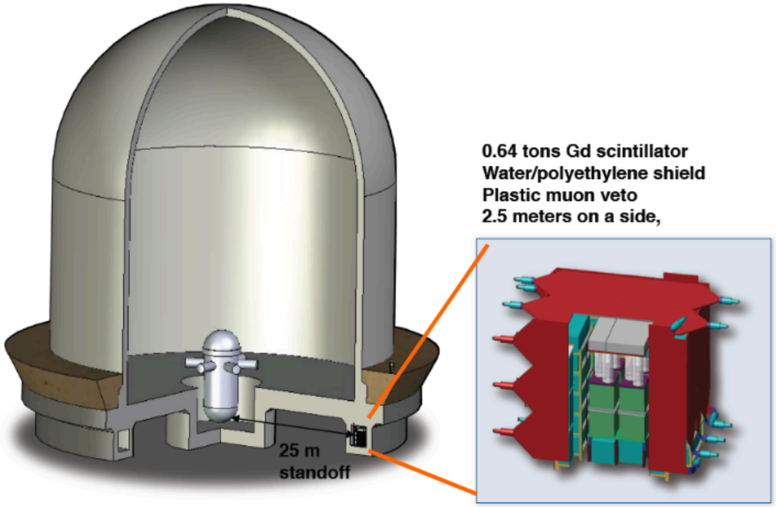
Akindele (Neutrieno2020), https://zenodo.org/record/3959532
SONGS, Journal of Applied Physics, 105(6), 064902 (2009)

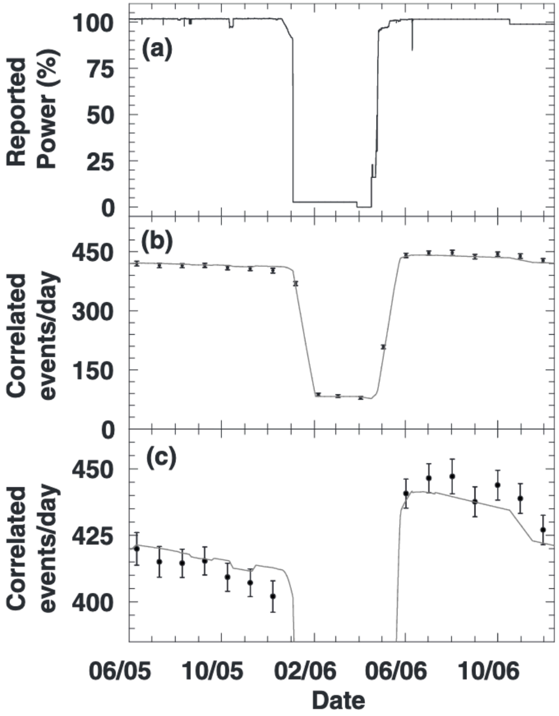
4. Reactor neutrino

Low energy electron anti-neutrinos
- High-precision spectrum prediction
- Monitoring fission reactor





0.64 tons Gd scintillator
Water/polyethylene shield
Plastic muon veto
2.5 meters on a side,



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Cleveland et al.,Astrophys.J.496(1998)505
Pontecorvo,Phys.Lett.28B(1969)493

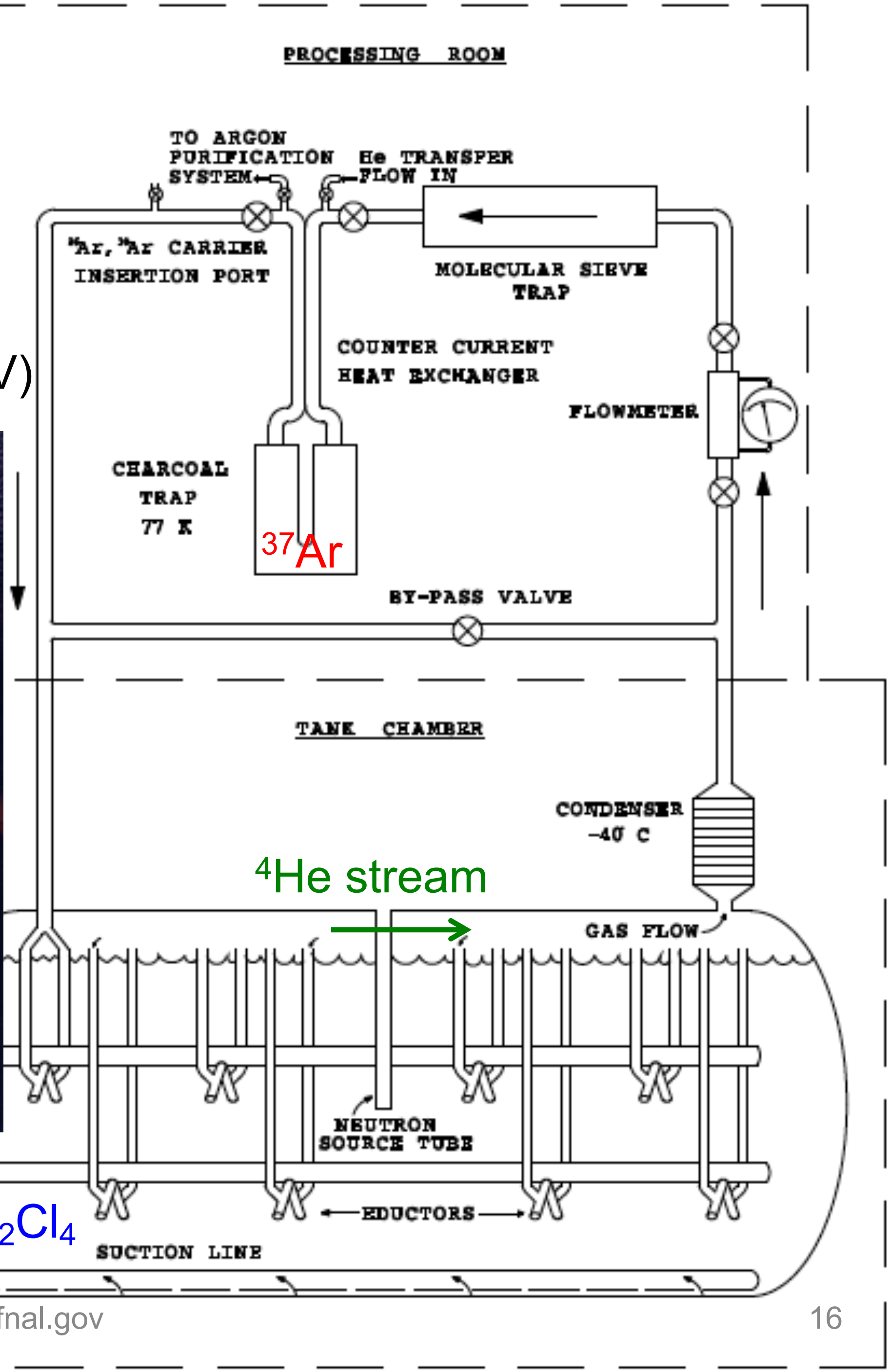
4. Solar neutrino

Homestake experiment

$$\nu_e + {}^{37}\mathrm{Cl} \rightarrow e^- + {}^{37}\mathrm{Ar}$$

(proposed by Pontecorvo)

- mainly sensitive to ^{8}B neutrino (~10 MeV)



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GALLEX, PLB490(2000)16;SAGE, J.Expt.Theor.Phys.95(2002)181
Borexino, PRL 108(2012)051302

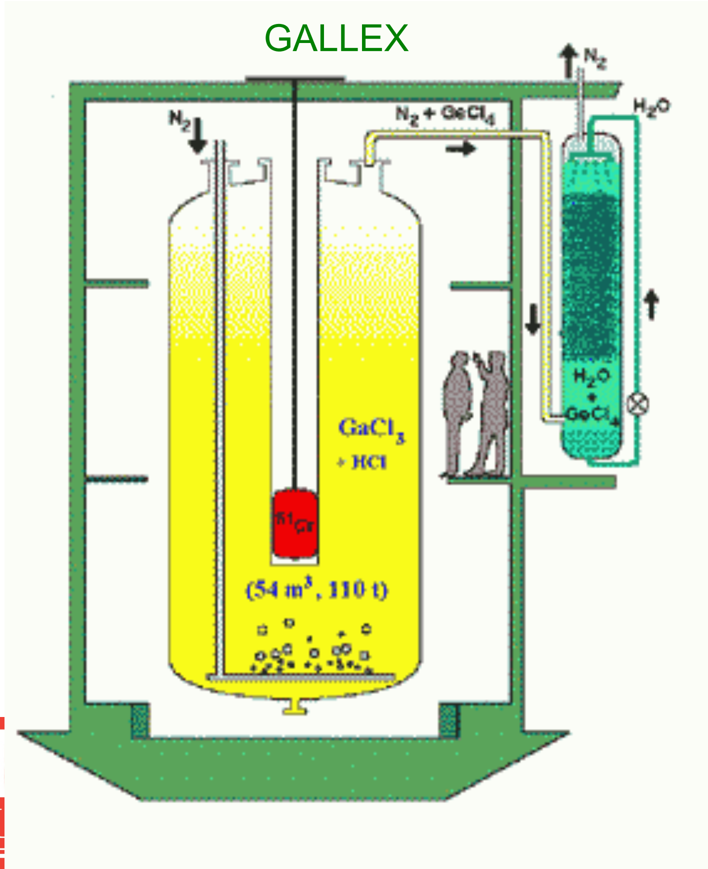
4. Solar neutrino

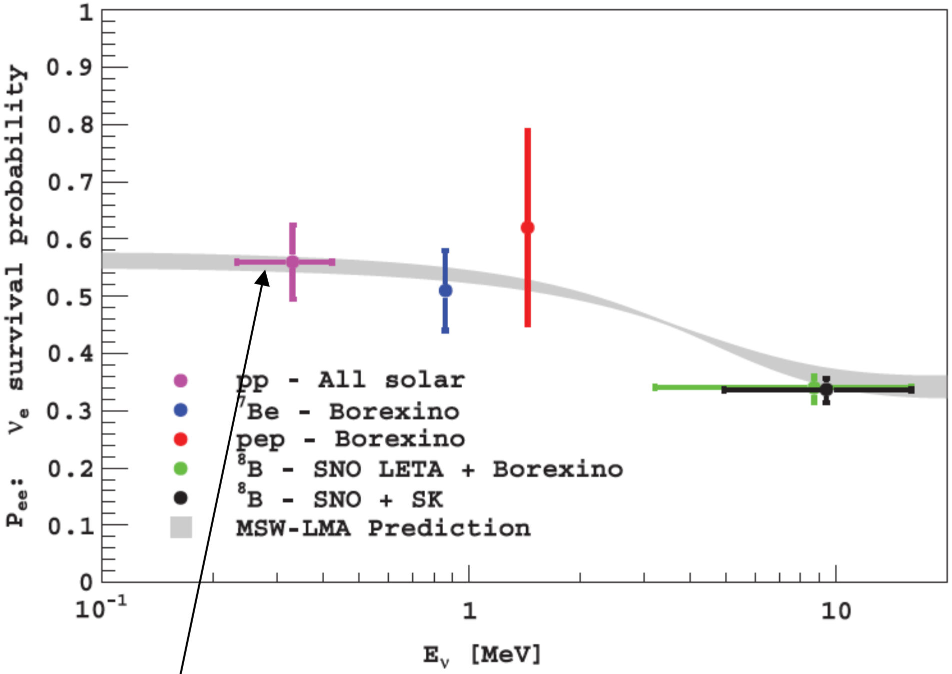
Gallium experiment

$$\nu_e + {}^{71}\mathrm{Ga} \rightarrow e^- + {}^{71}\mathrm{Ge}$$

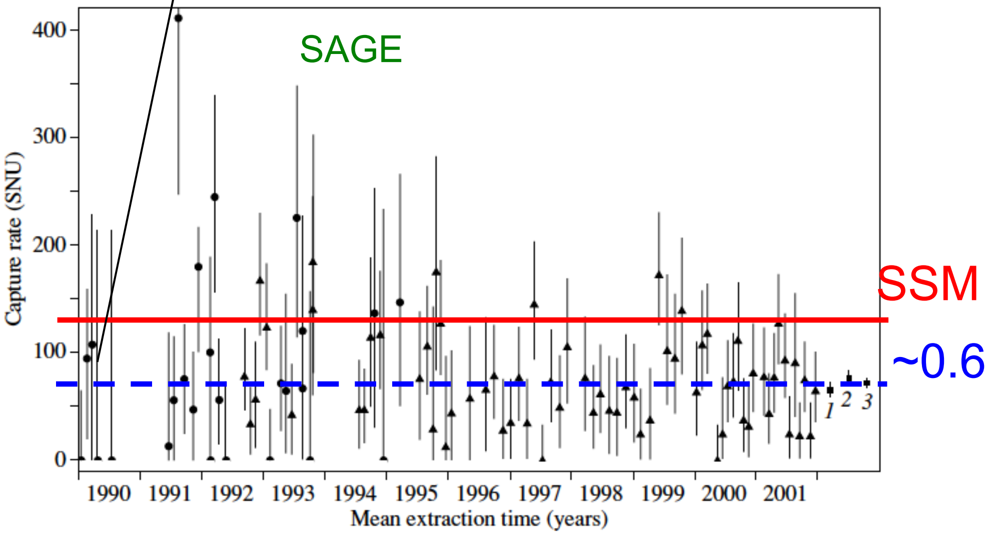
- Sensitive to pp-neutrino (0.42 MeV), 90% of total solar neutrino flux.
- Both experiments observed deficit, but higher than Homestake result

GALLEX





SAGE



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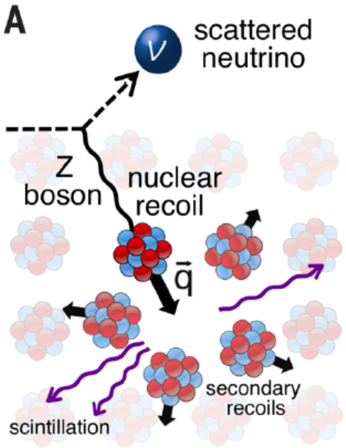
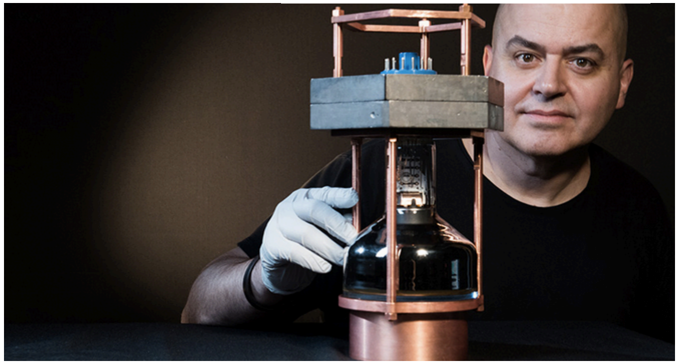
COHERENT, Science10.1126/science.aao0990 (2017).

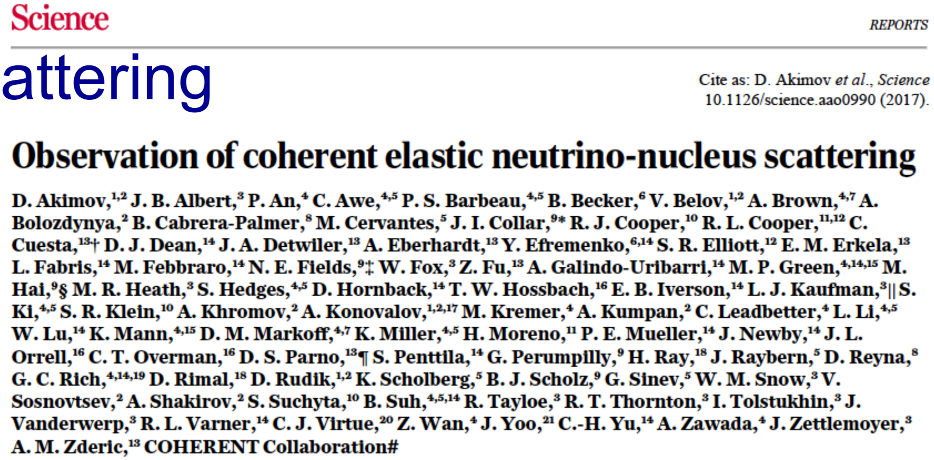
4. Neutrino-Nucleus coherent scattering

Low energy neutrinos from neutron sources at SNS (spallation neutron source), ORNL (Oak Ridge National Lab)

$\nu + A \rightarrow \nu + A$





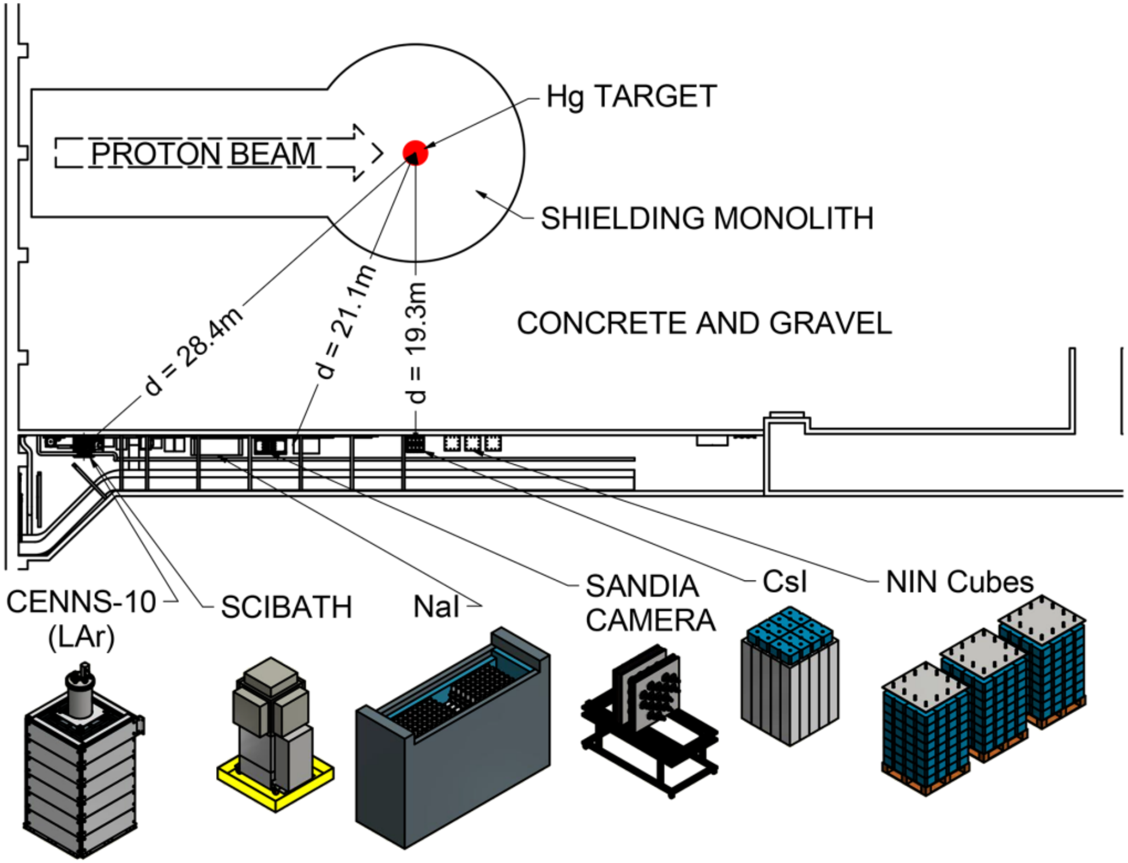


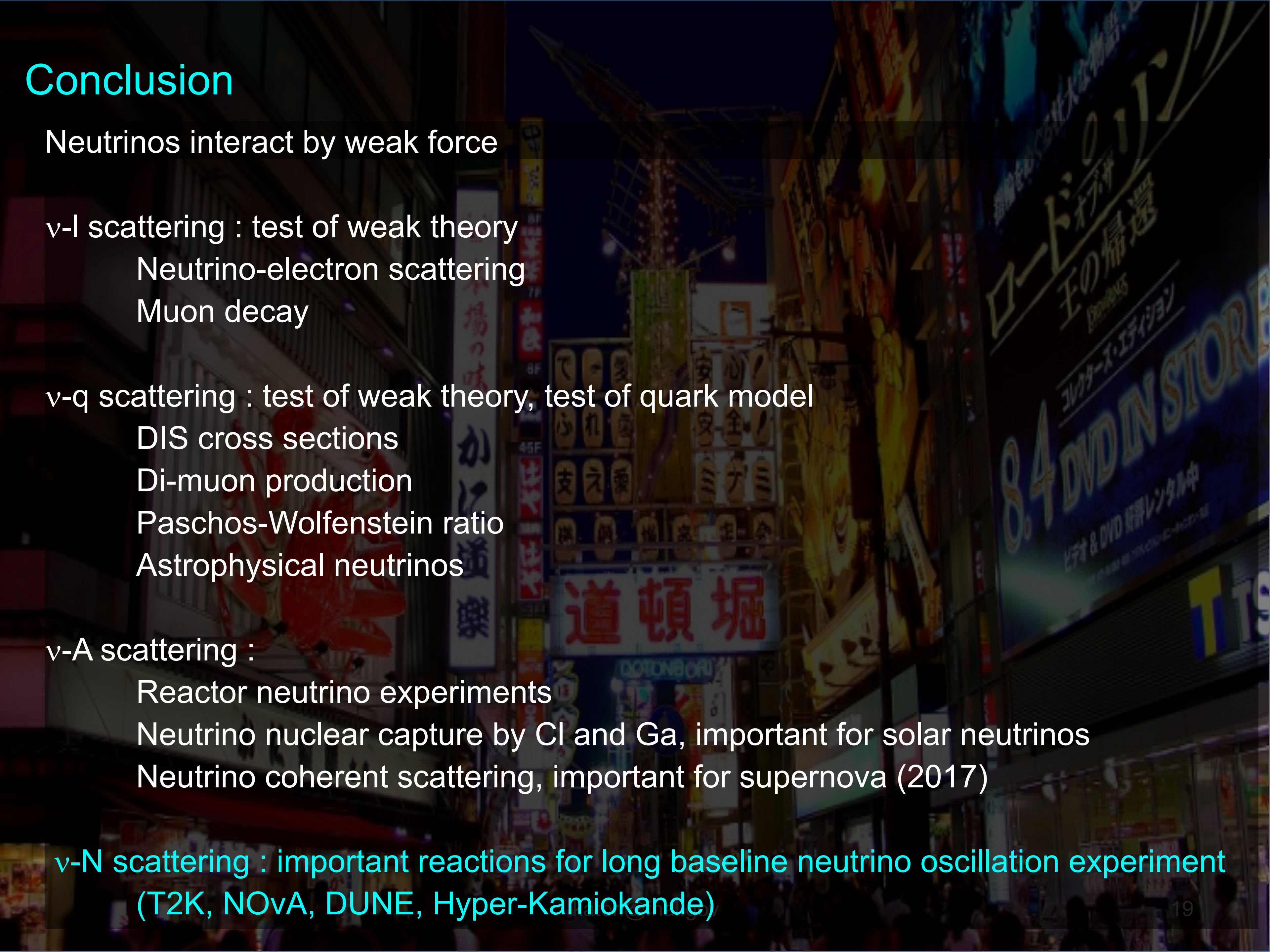
Science REPORTS

Cite as: D. Akimov *et al.*, *Science* 10.1126/science.aao0990 (2017).

Observation of coherent elastic neutrino-nucleus scattering

D. Akimov,[1,2] J. B. Albert,[3] P. An,[4] C. Awe,[4,5] P. S. Barbeau,[4,5] B. Becker,[6] V. Belov,[1,2] A. Brown,[4,7] A. Bolozdynya,[2] B. Cabrera-Palmer,[8] M. Cervantes,[5] J. I. Collar,[9]* R. J. Cooper,[10] R. L. Cooper,[11,12] C. Cuesta,[13]† D. J. Dean,[14] J. A. Detwiler,[13] A. Eberhardt,[13] Y. Efremenko,[6,14] S. R. Elliott,[12] E. M. Erkela,[13] L. Fabris,[14] M. Febbraro,[14] N. E. Fields,[9]‡ W. Fox,[3] Z. Fu,[13] A. Galindo-Uribarri,[14] M. P. Green,[4,14,15] M. Hai,[9]§ M. R. Heath,[3] S. Hedges,[4,5] D. Hornback,[14] T. W. Hossbach,[16] E. B. Iverson,[14] L. J. Kaufman,[3]|| S. Ki,[4,5] S. R. Klein,[10] A. Khromov,[2] A. Konovalov,[1,2,17] M. Kremer,[4] A. Kumpan,[2] C. Leadbetter,[4] L. Li,[4,5] W. Lu,[14] K. Mann,[4,15] D. M. Markoff,[4,7] K. Miller,[4,5] H. Moreno,[11] P. E. Mueller,[14] J. Newby,[14] J. L. Orrell,[16] C. T. Overman,[16] D. S. Parno,[13]¶ S. Penttila,[14] G. Perumpilly,[9] H. Ray,[18] J. Raybern,[5] D. Reyna,[8] G. C. Rich,[4,14,19] D. Rimal,[18] D. Rudik,[1,2] K. Scholberg,[5] B. J. Scholz,[9] G. Sinev,[5] W. M. Snow,[3] V. Sosnovtsev,[2] A. Shakirov,[2] S. Suchyta,[10] B. Suh,[4,5,14] R. Tayloe,[3] R. T. Thornton,[3] I. Tolstukhin,[3] J. Vanderwerp,[3] R. L. Varner,[14] C. J. Virtue,[20] Z. Wan,[4] J. Yoo,[21] C.-H. Yu,[14] A. Zawada,[4] J. Zettlemoyer,[3] A. M. Zderic,[13] COHERENT Collaboration#





Conclusion

Neutrinos interact by weak force

ν-l scattering : test of weak theory
- Neutrino-electron scattering
- Muon decay

ν-q scattering : test of weak theory, test of quark model
- DIS cross sections
- Di-muon production
- Paschos-Wolfenstein ratio
- Astrophysical neutrinos

ν-A scattering :
- Reactor neutrino experiments
- Neutrino nuclear capture by Cl and Ga, important for solar neutrinos
- Neutrino coherent scattering, important for supernova (2017)

ν-N scattering : important reactions for long baseline neutrino oscillation experiment
- (T2K, NOvA, DUNE, Hyper-Kamiokande)

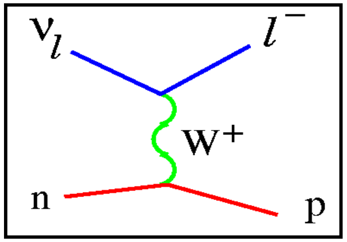
Formaggio and Zeller, Rev.Mod.Phys.84(2012)1307

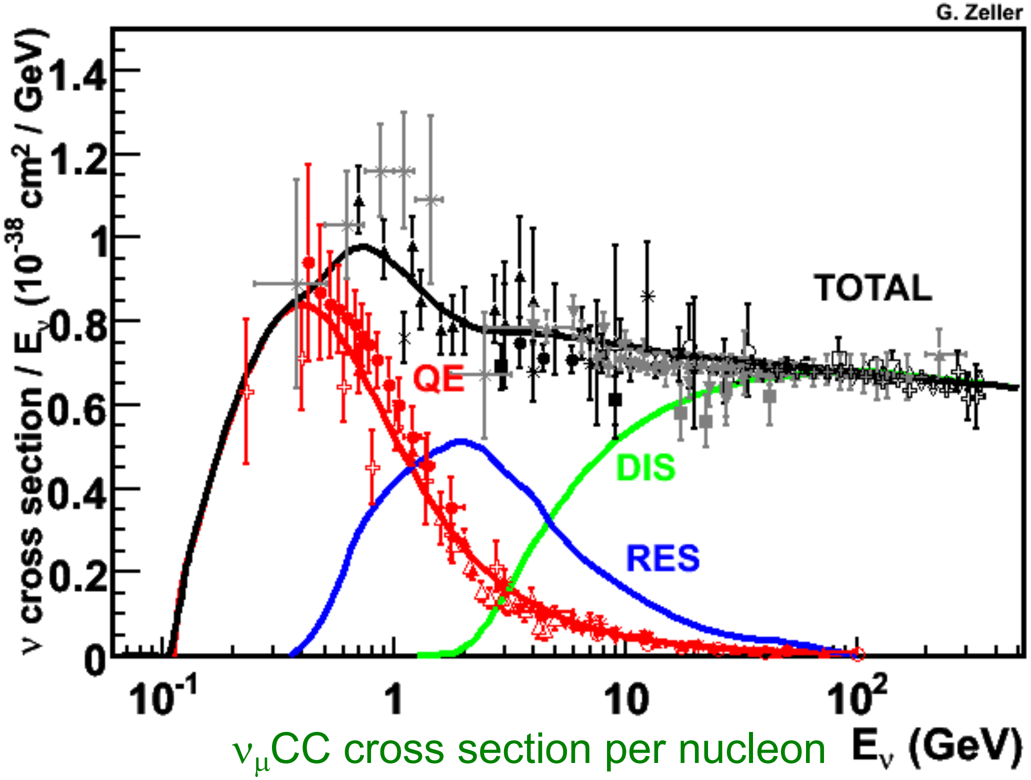
1. Next generation neutrino oscillation experiments

Neutrino oscillation experiments

- Present to Future: T2K, NOvA, PINGU, ORCA, Hyper-Kamiokande, DUNE

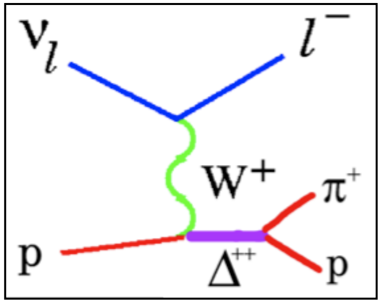
Quasi Elastic

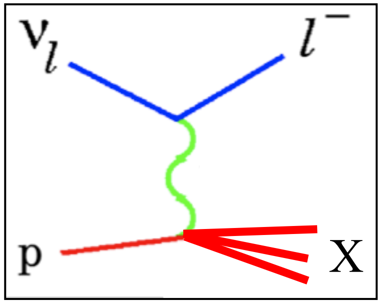




ν_μCC cross section per nucleon

baryonic RESonance





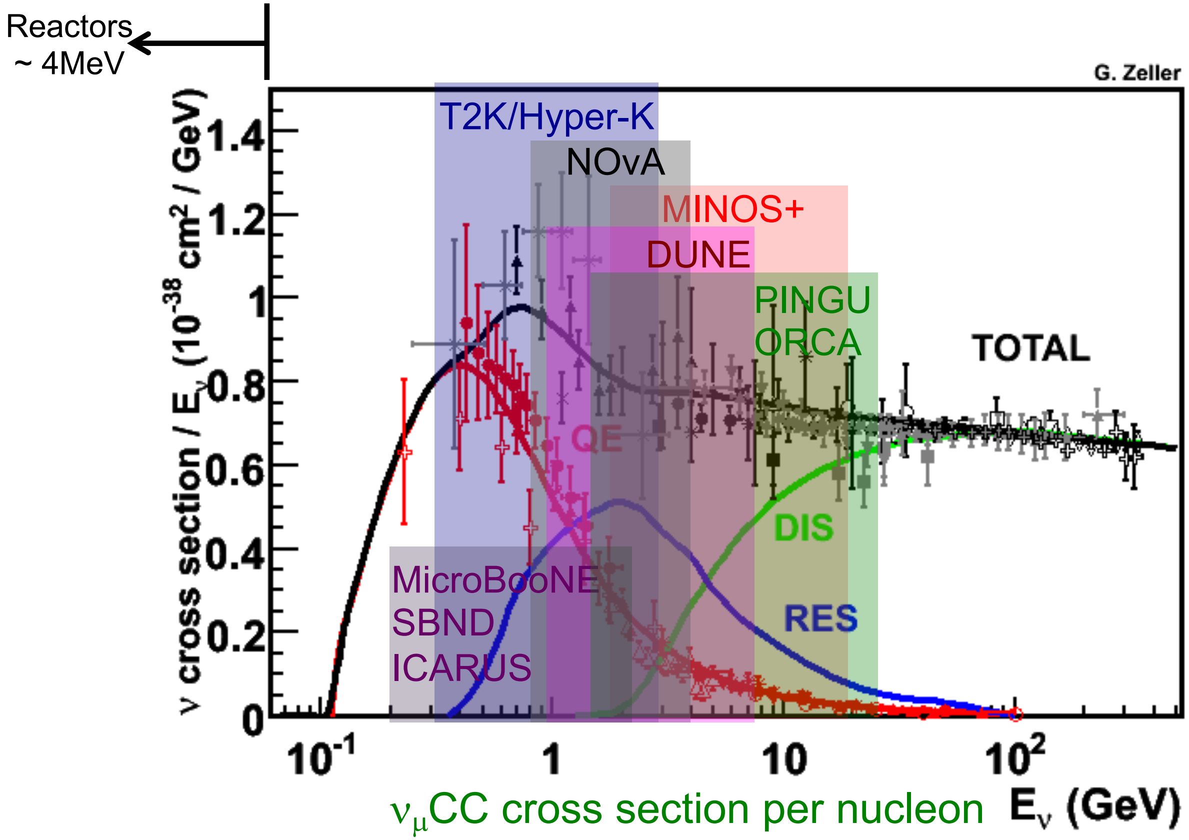
Deep Inelastic Scattering

Formaggio and Zeller, Rev.Mod.Phys.84(2012)1307

1. Next generation neutrino oscillation experiments

Neutrino oscillation experiments

- Present to Future: T2K, NOvA, PINGU, ORCA, Hyper-Kamiokande, DUNE…



ν_μCC cross section per nucleon

$$P_{\mu\to e}(L/E)=\sin^2 2\theta \sin^2\left(1.27\Delta m^2(eV^2)\frac{L(km)}{E(GeV)}\right)$$

1. Next generation neutrino oscillation experiments

Neutrino oscillation experiments

- Present to Future: T2K, NOvA, PINGU, ORCA, Hyper-Kamiokande, DUNE…

Most of data are from muon neutrino beam

- create by π-DIF (pion decay-in-flight)
- $\Phi(\nu_\mu) > \Phi(\bar{\nu}_\mu)$: more π^+ than π^- (because they are made by protons)
- δ_{CP} study need electro-neutrino cross-section (ν_e appearance) and anti-neutrino cross-section (CP violation)

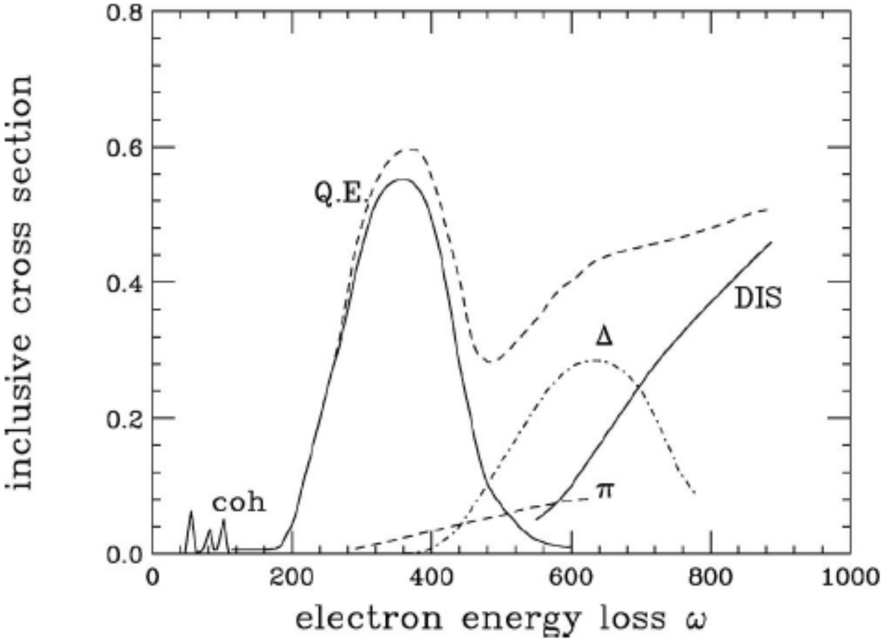
Nuclear physics sucks

- Simple extrapolation may be broken due to nuclear physics
- We are not good at nuclear physics because we are not nuclear physicists
- Nuclear physics = non-perturbative QCD (many models, no theory)
- Particle physics is developed by avoiding nuclear physics…

$$\boxed{P_{\mu\to e}(L/E)=\sin^2 2\theta\sin^2\left(1.27\Delta m^2(eV^2)\frac{L(km)}{E(GeV)}\right)}$$

Benhar,Day,Sick,Rev.Mod.Phys.80(2008)189
Nakamura et a;,Rep.Prog.Phys.80(2017)056301

1. Particle Physics vs. Nuclear Physics

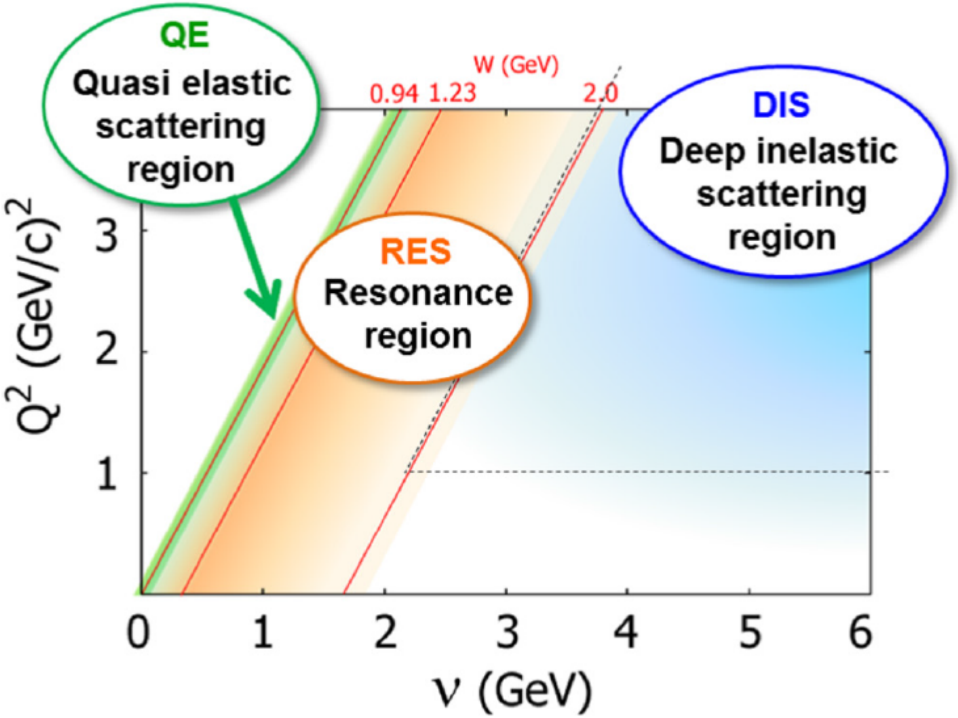


Nuclear physics

Interactions are classified in q (3-momnetum transfer) and ω (energy transfer)

Particle physics (neutrino physics)

Interactions are classified in Q^2 (4-momentum transfer) and ν (energy transfer) or W^2 (invariant mass)



1. Neutrino cross-section formula

Cross-section

- product of Leptonic and Hadronic tensor

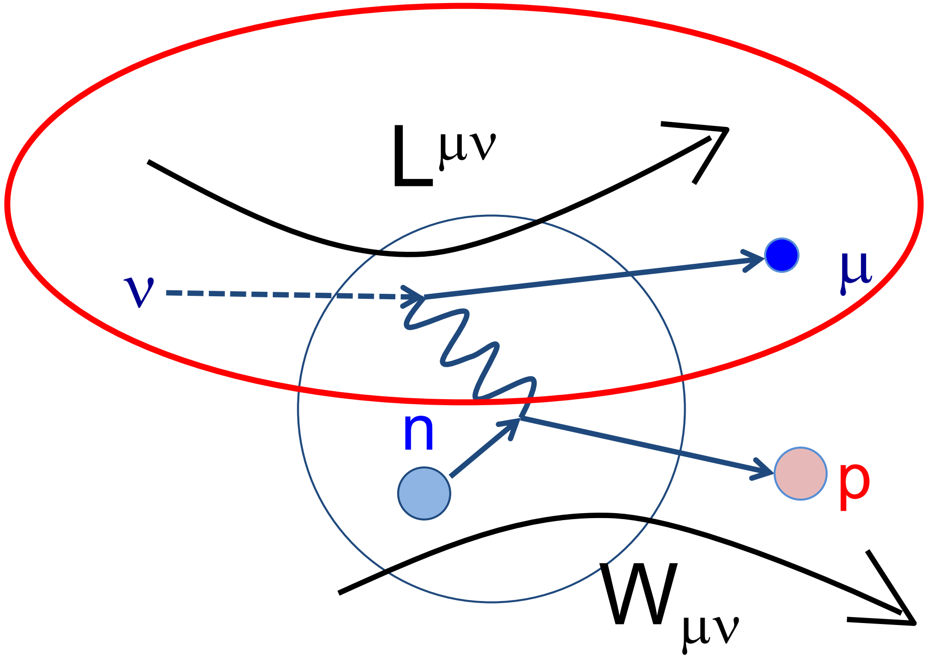
$$d\sigma \sim L^{\mu\nu}W_{\mu\nu}$$

Leptonic tensor

→ the Standard Model (easy)

Hadronic tensor

→ nuclear physics (hard)



katori@fnal.gov

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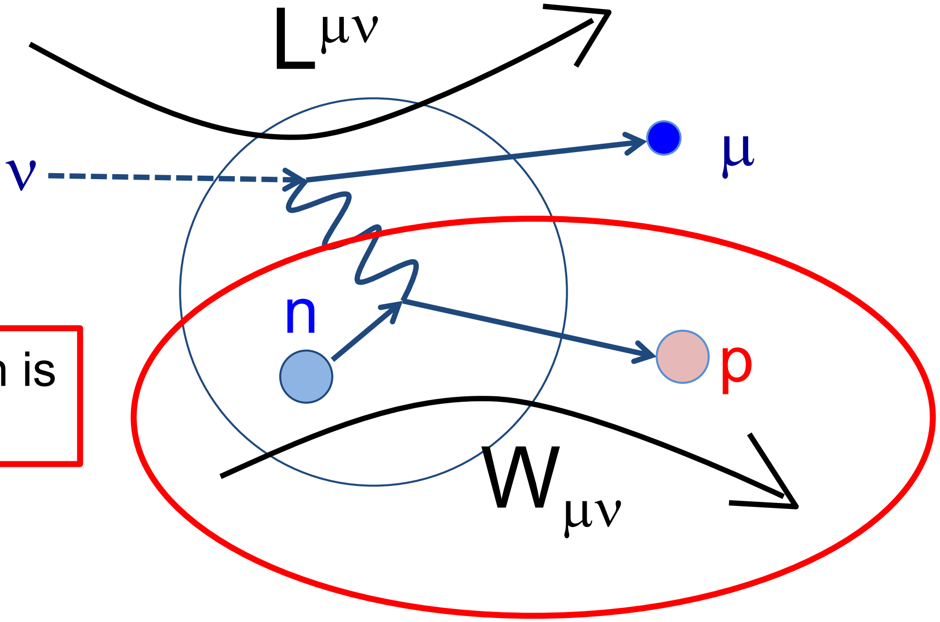

All complication of neutrino cross-section is how to model the hadronic tensor part

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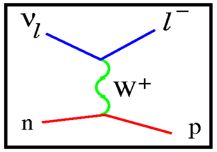
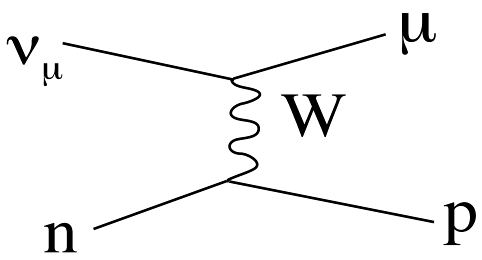
Quasi Elastic



2. Charged Current Quasi-Elastic scattering (CCQE)

The simplest and the most abundant interaction around ~1 GeV.

$$\nu_\mu + n \rightarrow p + \mu^- \quad (\nu_\mu + X \rightarrow X' + \mu^-)$$



$d\sigma \sim L_{\mu\nu} T^{\mu\nu}$

$L_{\mu\nu} \sim J_\mu J_\nu$: Lepton tensor

$W_{\mu\nu} = \int f(\vec{k}, \vec{q}, \omega) T_{\mu\nu} dE$: hadronic tensor

$f(\vec{k}, \vec{q}, \omega)$: nucleon phase space

$T_{\mu\nu} = T_{\mu\nu}(F_1, F_2, F_A, F_P)$: form factors

Form factors can be parameterized with dipole form

$$F(Q^2) = \frac{g}{\left(1 + \frac{Q^2}{M^2}\right)^2}$$

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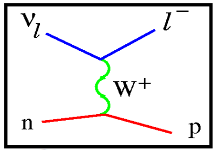
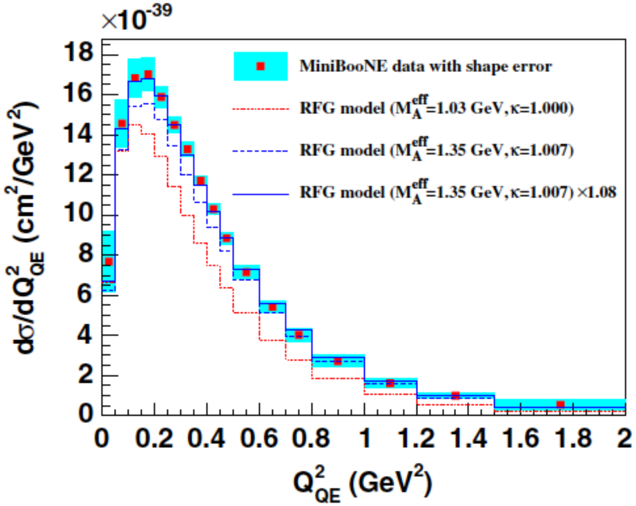
Quasi Elastic



2. Form factors

MiniBooNE CCQE cross section
PRD81(2010)092005



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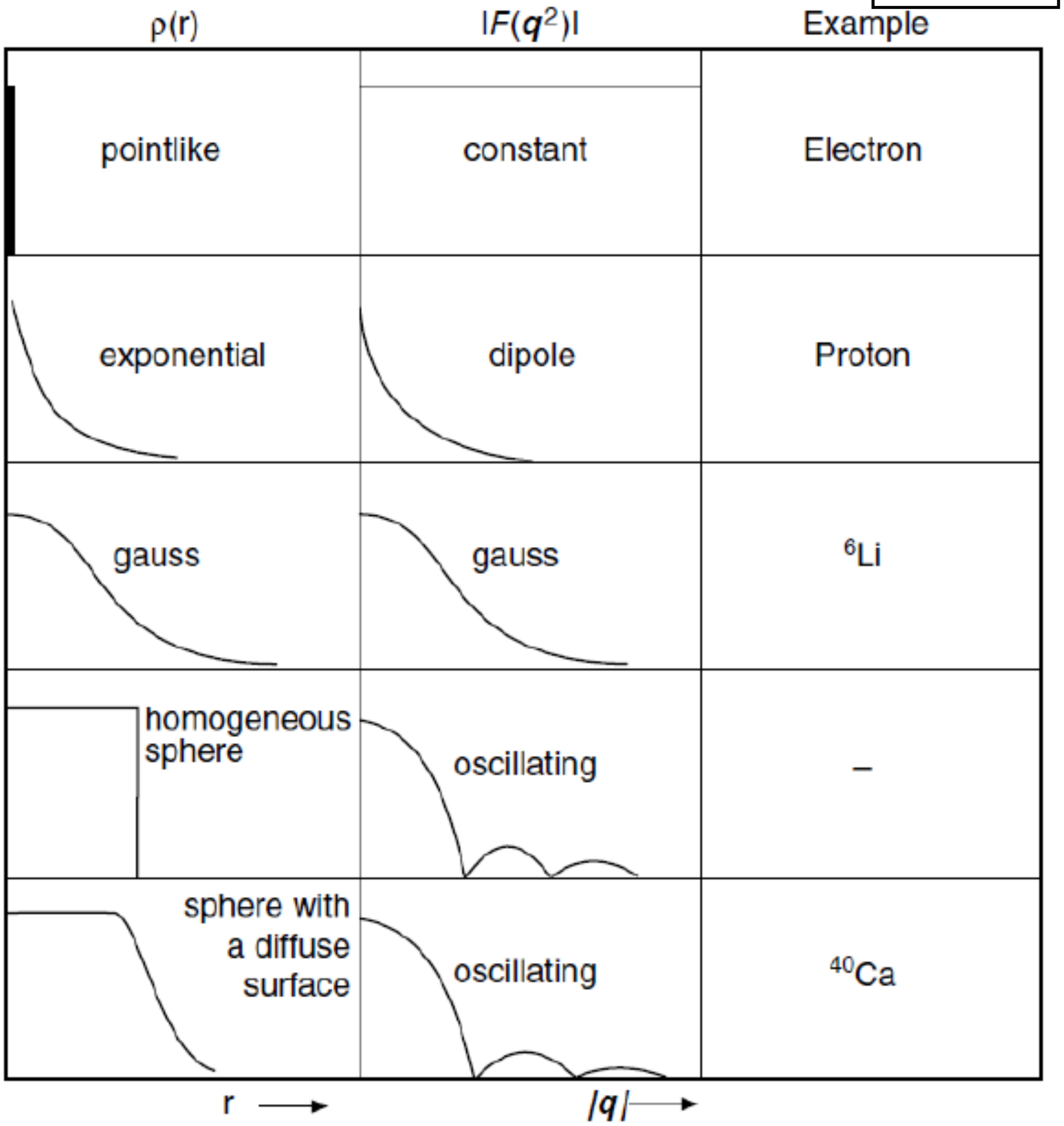
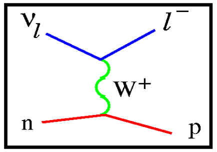


ρ(r)	$\lvert F(q^2) \rvert$	Example
pointlike	constant	Electron
exponential	dipole	Proton
gauss	gauss	^{6}Li
homogeneous sphere	oscillating	–
sphere with a diffuse surface	oscillating	^{40}Ca

2. Charged Current Quasi-Elastic scattering (CCQE)

Quasi Elastic



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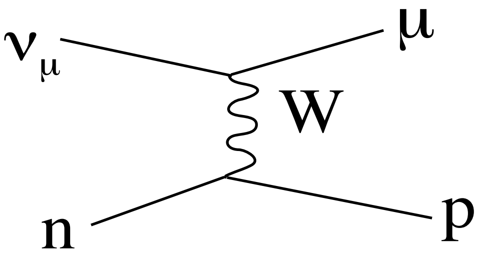
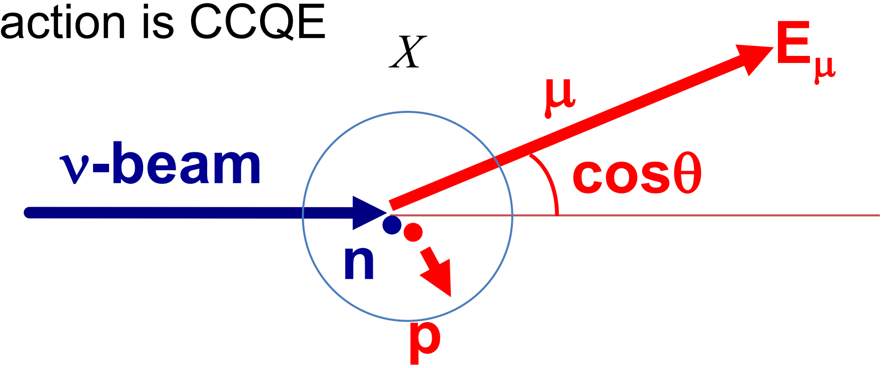


Neutrino energy is reconstructed from the observed lepton kinematics

"QE assumption"

1. assuming neutron at rest
2. assuming interaction is CCQE

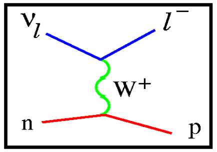


$$E_\nu^{QE} = \frac{ME_\nu - 0.5m_\mu^2}{M - E_\mu + p_\mu cos\theta}$$

Neutrinos hit nucleons inside of nucleus, and the energy reconstruction is possible only with QE assumption

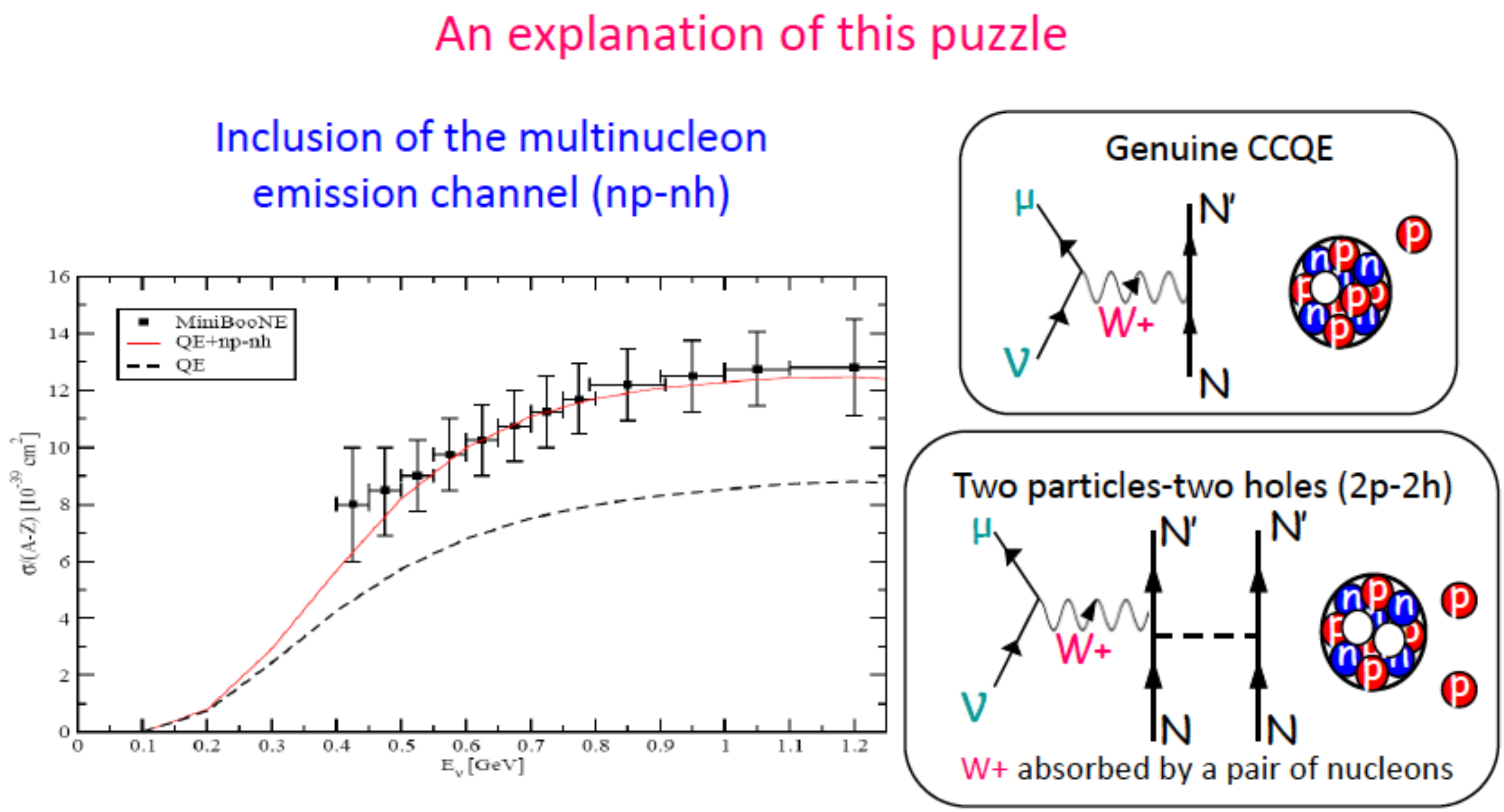
Martini et al, PRC80(2009)065501

Quasi Elastic



2. Nucleon correlations

2-particle 2-hole (2p2h) effect
- Mimic CCQE interaction, significant change cross section (both shape and normalization)
- The biggest topic in nuxsec community (T2K, NOvA, MINERvA, MicroBooNE)



An explanation of this puzzle

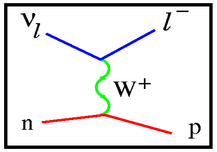
Inclusion of the multinucleon emission channel (np-nh)

W+ absorbed by a pair of nucleons

Martini et al, PRC80(2009)065501
NOvA, arXiv:2006.08727

2. Nucleon correlations

Quasi Elastic



2-particle 2-hole (2p2h) effect

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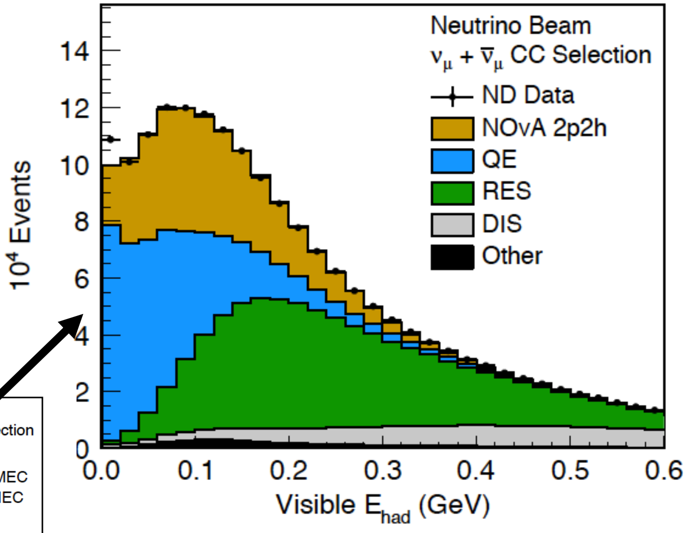
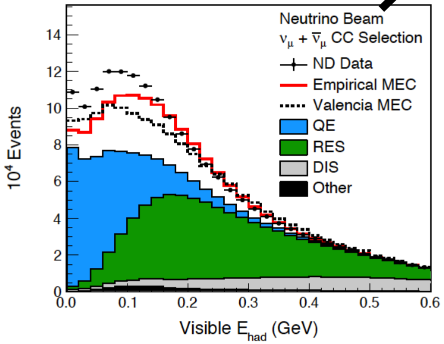
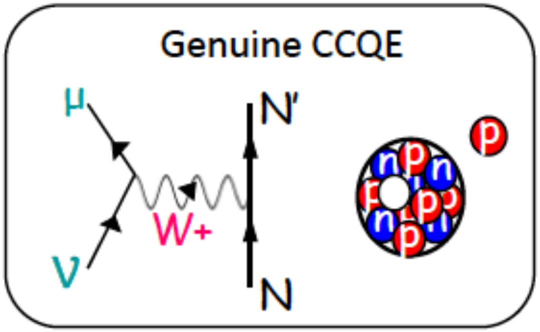
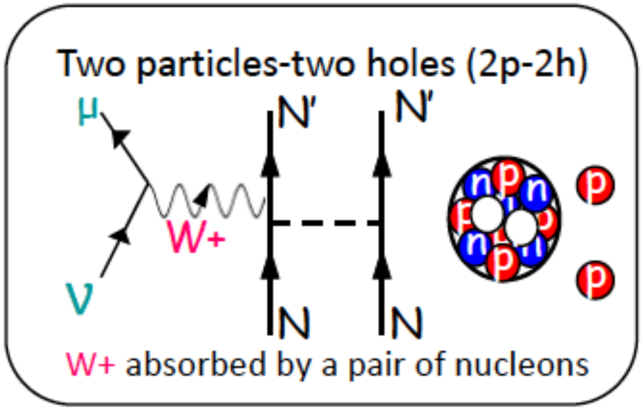
- The biggest topic in nuxsec community (T2K, NOvA, MINERvA, MicroBooNE)
- 2p2h models in generators don't describe data well?
- High resolution detector (LArTPC, emulsion, etc) can find what is going on?

NOvA near detector data-MC comparison after fit





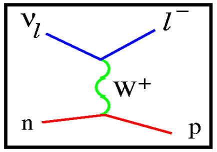




Smith and Moniz, NPB43(1972)605
Ashkenazi (Neutrino 2020), https://zenodo.org/record/3959538

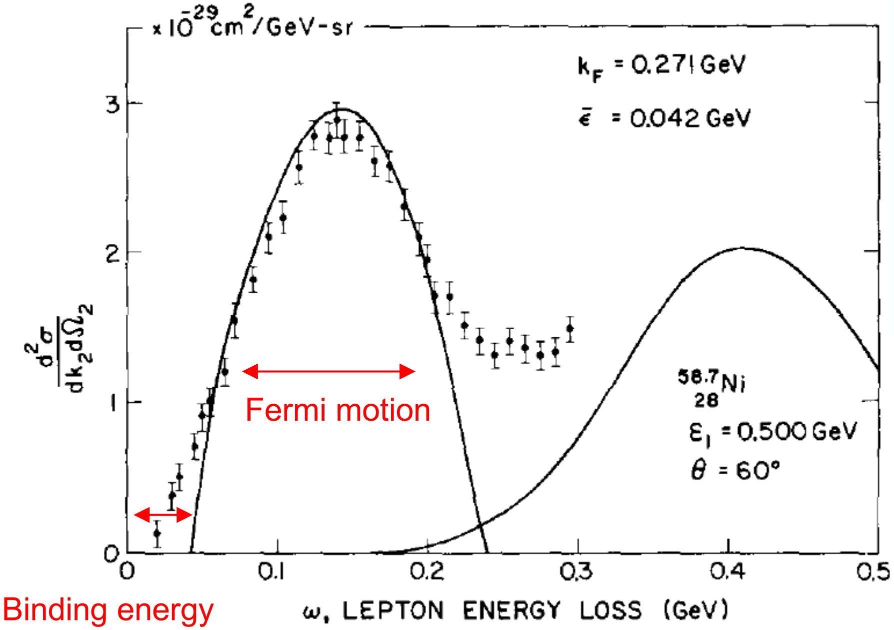
2. Fermi motion

Quasi Elastic

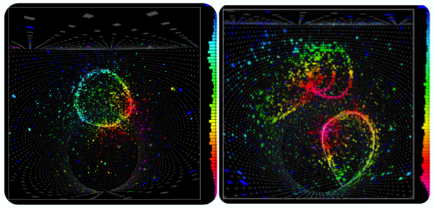


Fermi motion

- Measured energy is smeared from the true energy if you assume nucleon at rest
- High resolution detector can measure all outgoing hadrons
  → initial nucleon momentum can be reconstructed (no Fermi motion smearing)



Binding energy

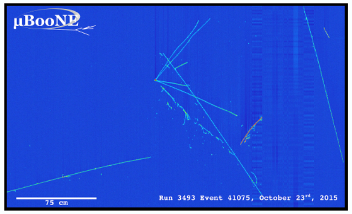


Cherenkov detectors:

Assuming QE interaction

Using lepton only

$$E_{QE} = \frac{2M\epsilon + 2ME_l - m_l^2}{2(M - E_l + |k_l|\cos\theta_l)}$$



Tracking detectors:

Calorimetric sum

Using All detected particles

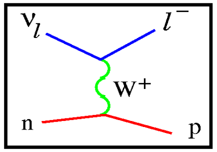
$$E_{\mathrm{cal}} = E_l + E_p^{\mathrm{kin}} + \epsilon$$

[1p0π]

Smith and Moniz, NPB43(1972)605
MiniBooNE, PRD81(2010)092005

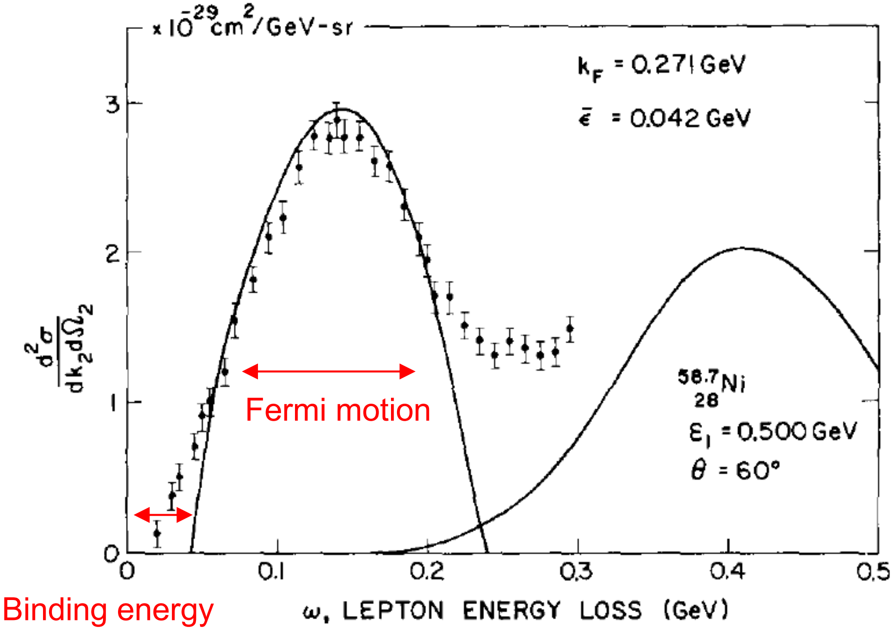
2. Pauli blocking

Quasi Elastic



Pauli blocking

- Low momentum transfer reaction is forbidden.
- data show more suppression than what Pauli blocking can → RPA(?)
- In the global Fermi gas model, Pauli blocking looks unphysical



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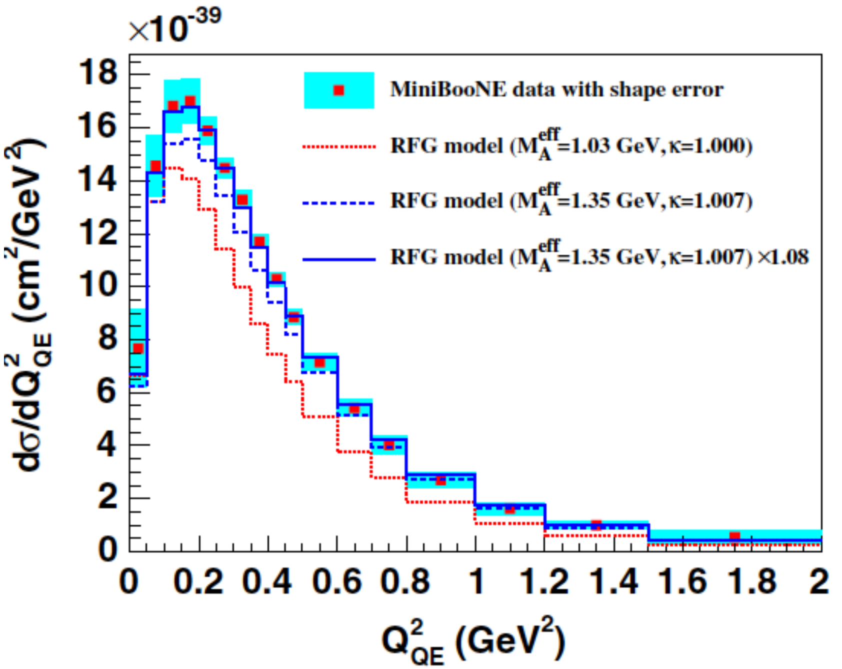
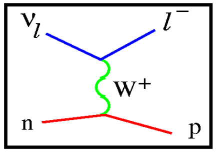


Smith and Moniz, NPB43(1972)605
Bodek and Cai, EPJC79(2019)293

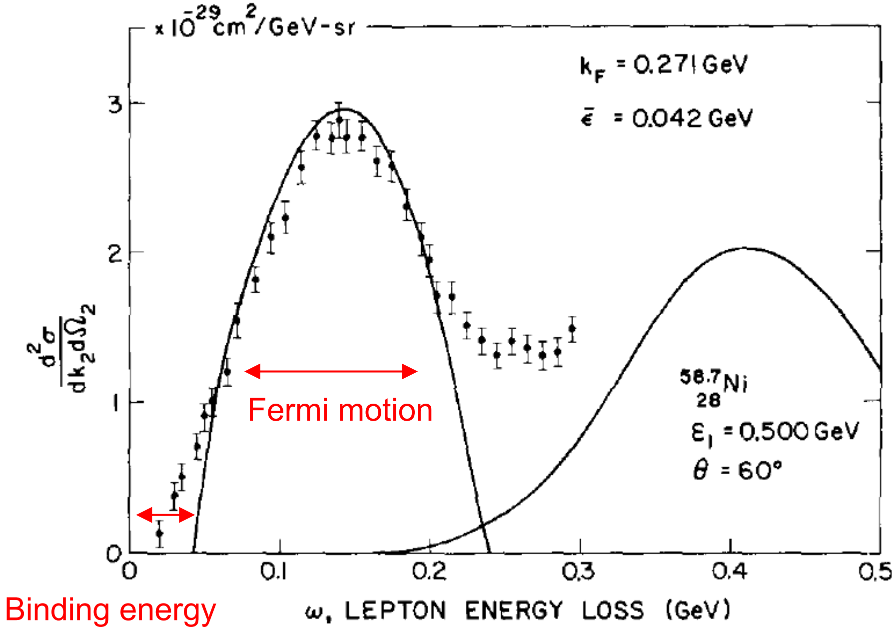
2. Nuclear Shell structure and binding energy

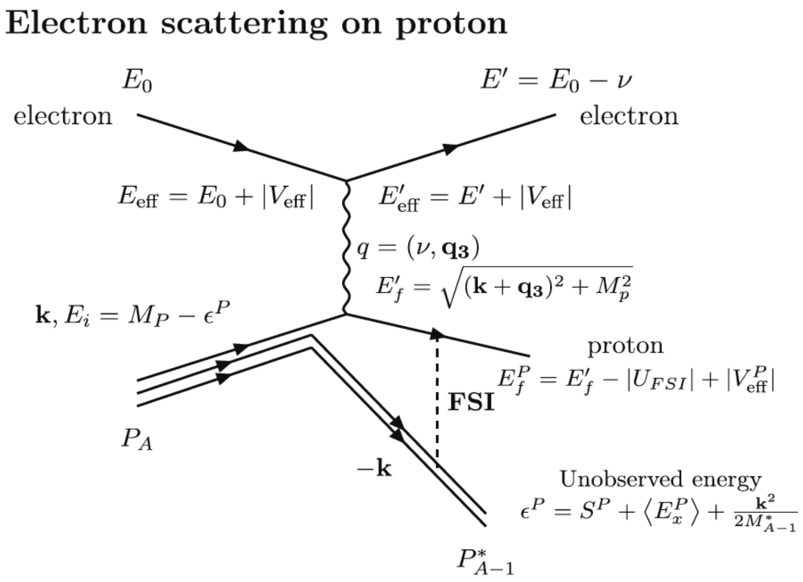
Quasi Elastic



Binding energy ~ unobserved energy

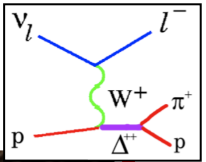
- Energy to cost to release 1 nucleon, not constant
- Separation energy + excitation energy + recoil energy
  - Separation energy: energy to release 1 nucleon from the shell (~15 MeV, depends)
  - Excitation energy: energy used to excite leftover target nucleus (~1 MeV)
  - Recoil energy: kinetic energy of recoil target nucleus (~2-3 MeV)



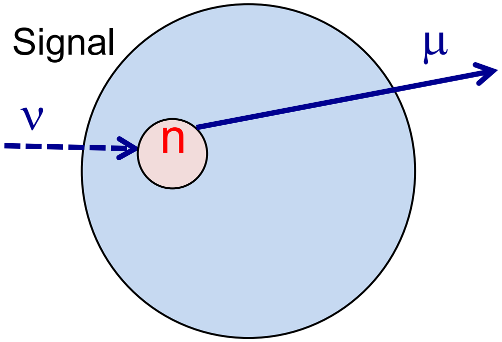


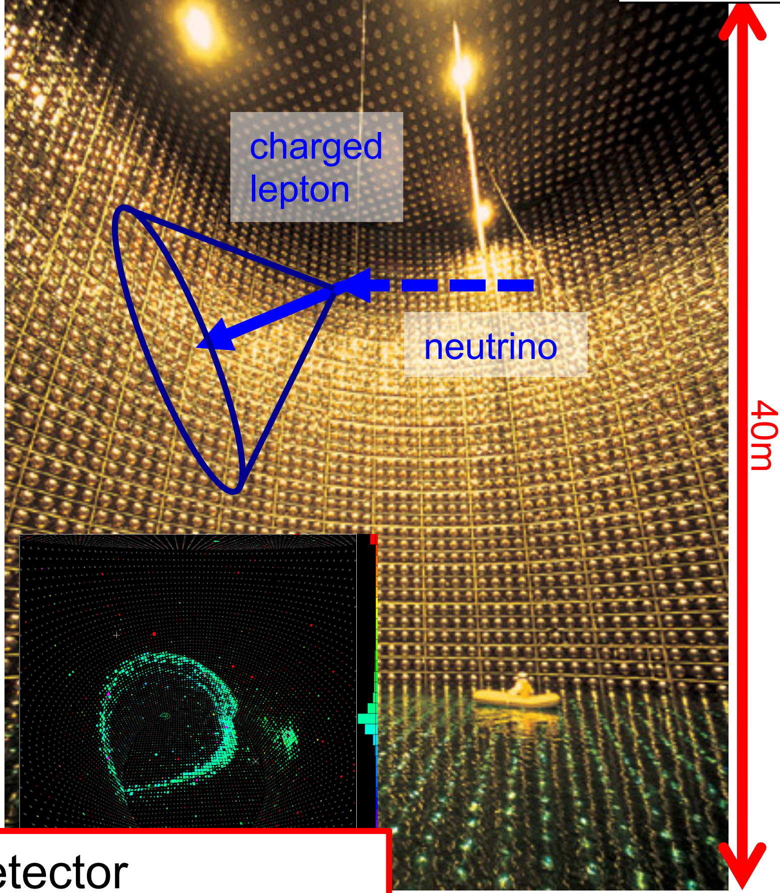
3. non-QE background (resonance pion production)

RESonance



non-QE background → shift spectrum





Neutrino energy is reconstructed from the observed lepton kinematics
"QE assumption"
1. assuming neutron at rest
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Typical neutrino oscillation detector
- Big and dense, to maximize interaction rate
- Coarsely instrumented, to minimize cost

(not great detector to measure hadrons)

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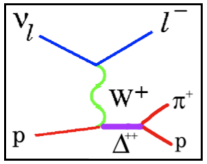

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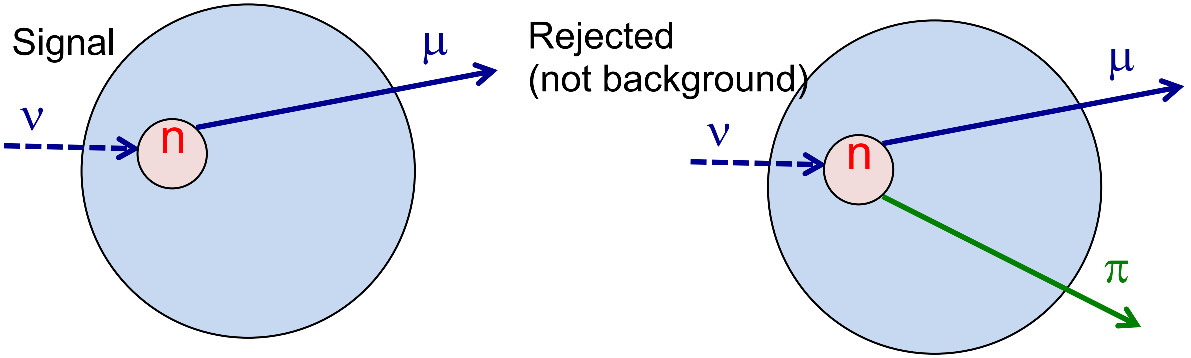

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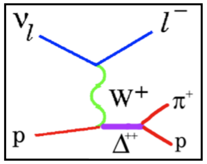

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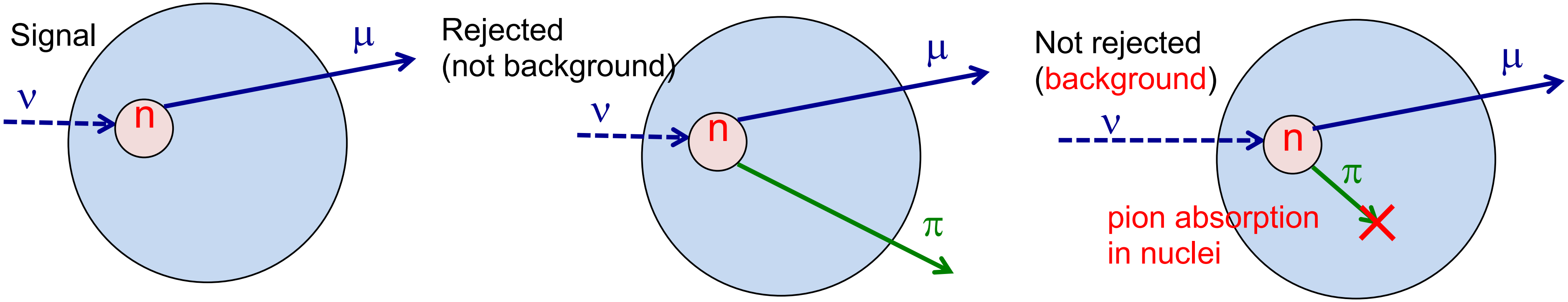

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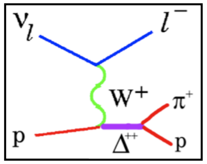

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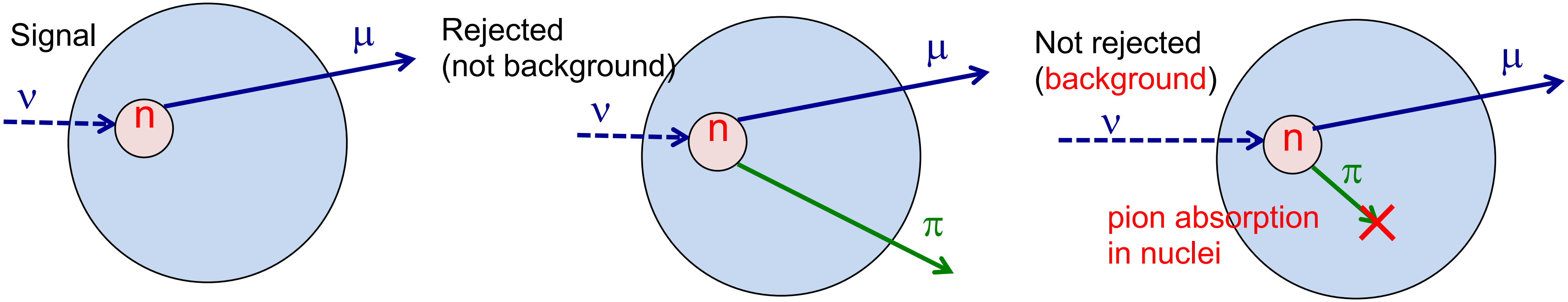
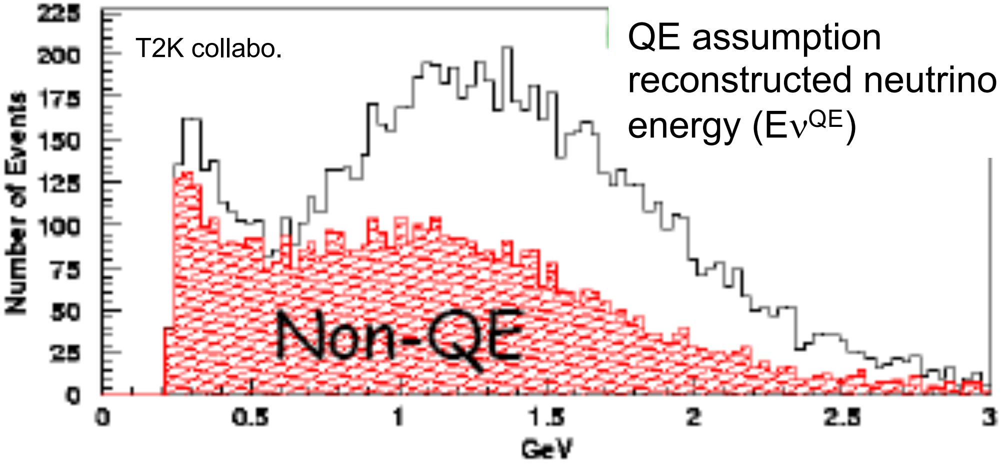

Signal

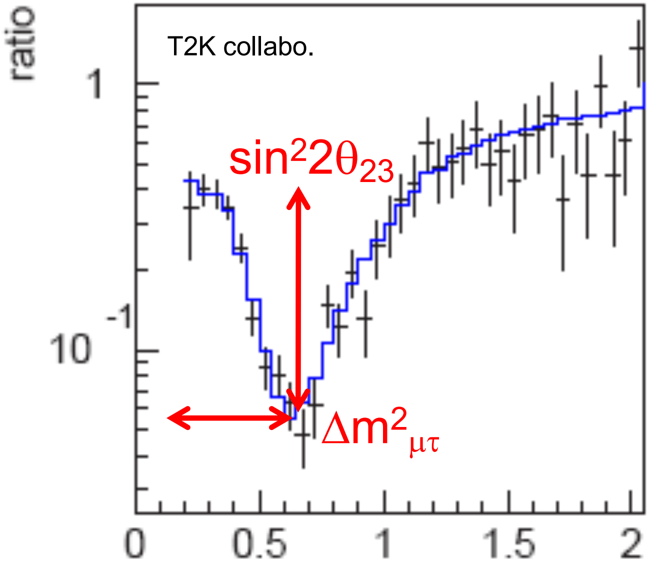
Rejected
(not background)

Not rejected
(background)

pion absorption
in nuclei

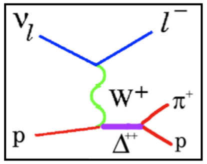


QE assumption reconstructed neutrino energy ($E\nu^{QE}$)



Reconstructed neutrino energy (GeV)

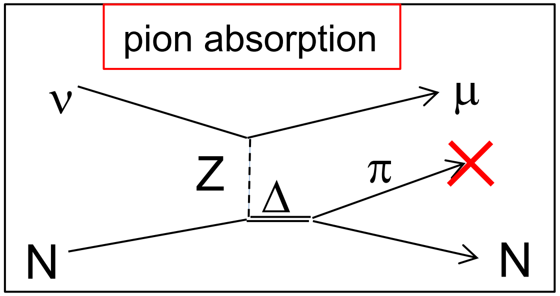
Coloma et al,PRL111(2013)221802
Mosel et al,PRL112(2014)151802



3. non-QE background (resonance pion production)

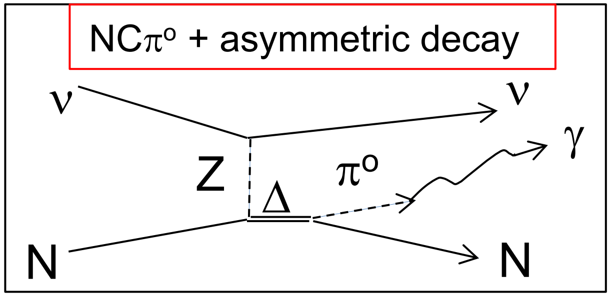
Pion production for ν_μ disappearance search

- Source of mis-reconstruction of neutrino energy

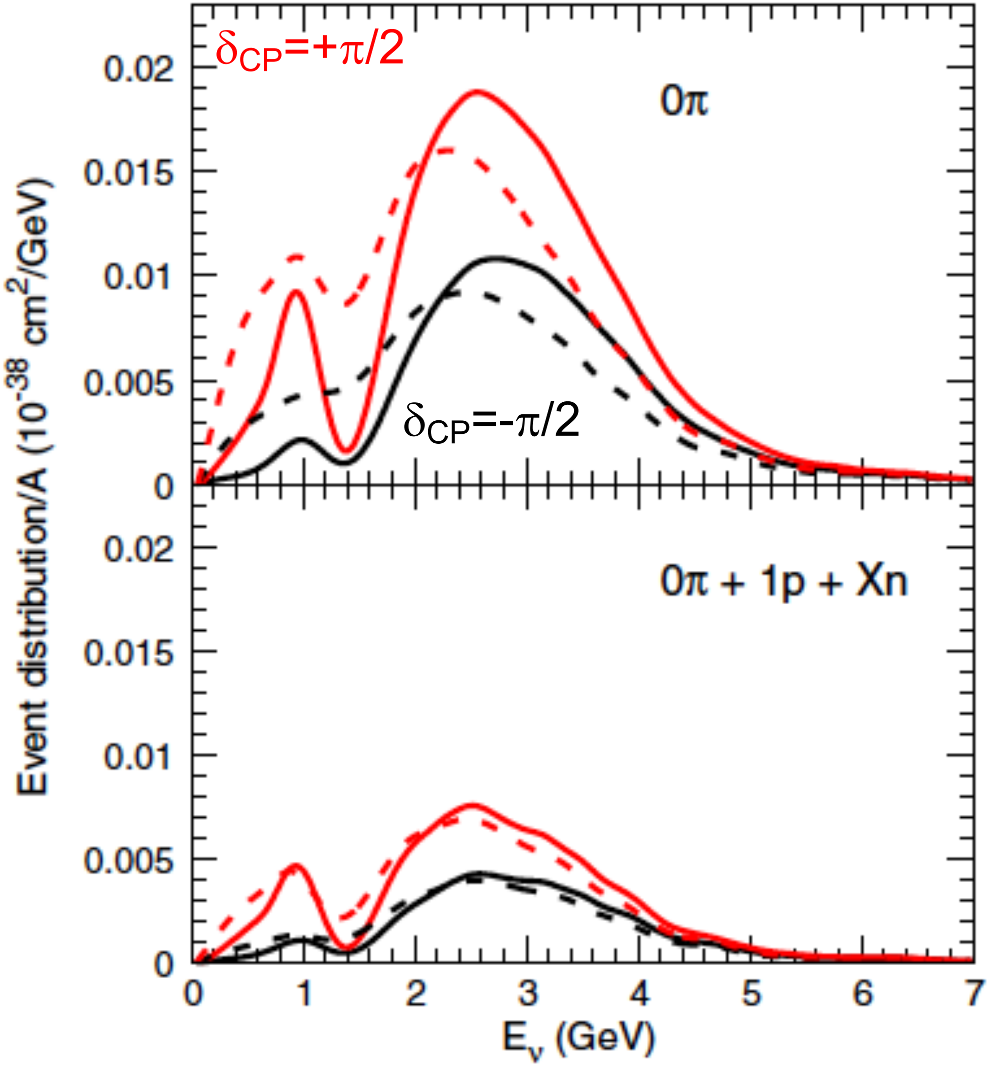


Neutral pion production in ν_e appearance search

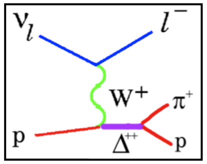
- Source of misID of electron



DUNE true vs. reconstructed Eν spectrum



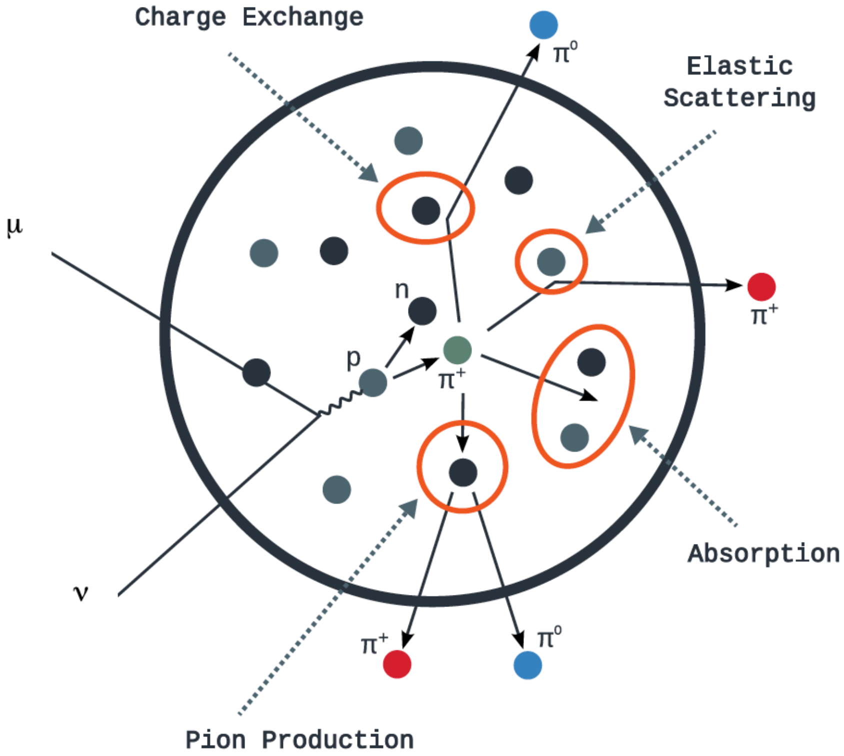
Understanding of neutrino baryonic resonance meson production is important for oscillation experiments



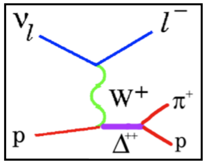
3. Final state interaction

Cascade model

- Elastic scattering: Nucleon elastic scattering, pion elastic scattering
- Inelastic scattering: Nucleon inelastic scattering, pion inelastic scattering
- Charge exchange: Nucleon charge exchange, pion charge exchange
- Absorption: Nucleon absorption, pion absorption



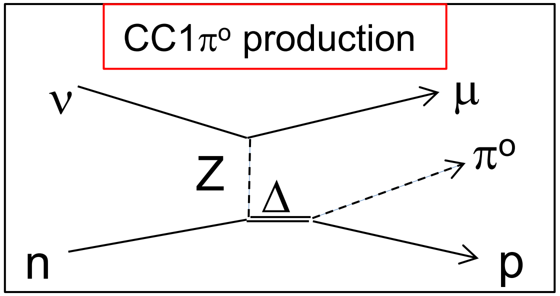
MiniBooNE,PRD83(2011)052009
Lalakulich et al,PRC87(2013)014602

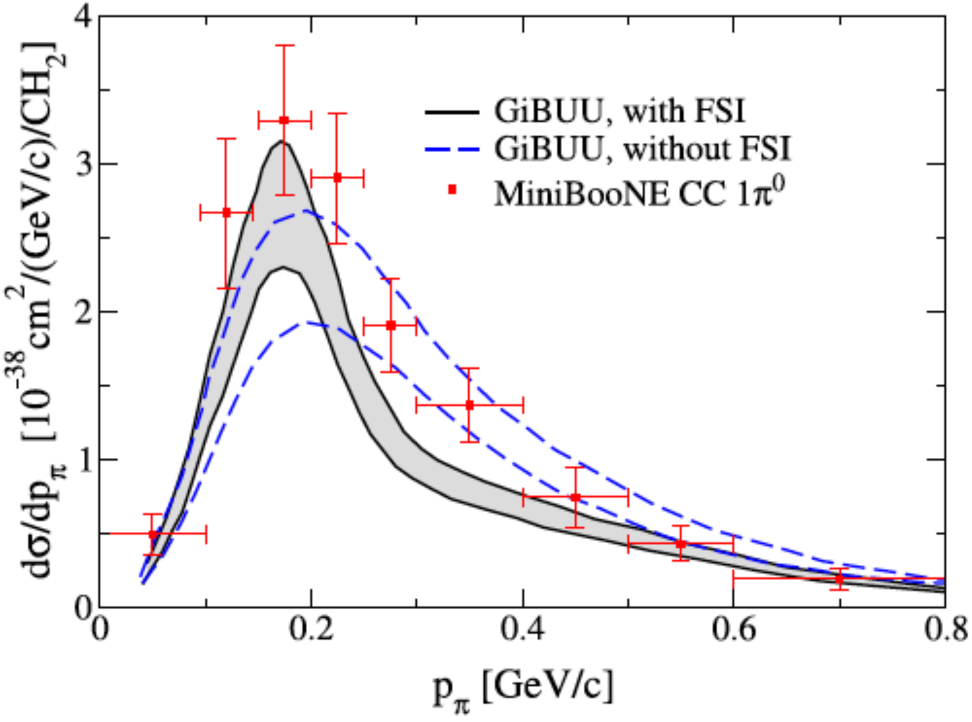


3. Neutrino Baryonic resonance data

Final state interaction

- Cascade model as a standard of the community
- Advanced models are not available for event-by-event simulation

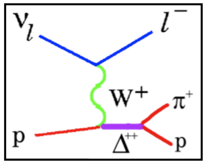




MiniBooNE π° momentum vs simulation

MiniBooNE,PRD83(2011)052009
Lalakulich et al,PRC87(2013)014602

RESonance



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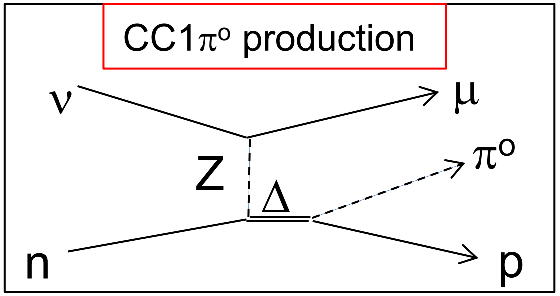
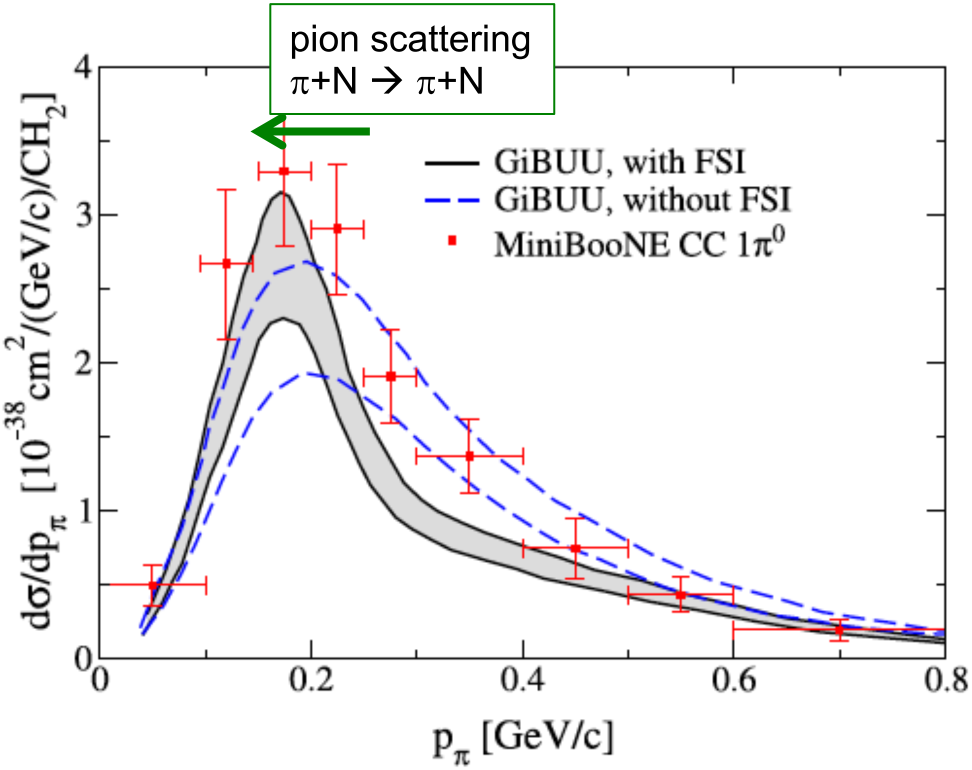
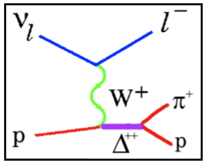



pion scattering
π+N → π+N

MiniBooNE π^o momentum vs simulation

katori@fnal.gov

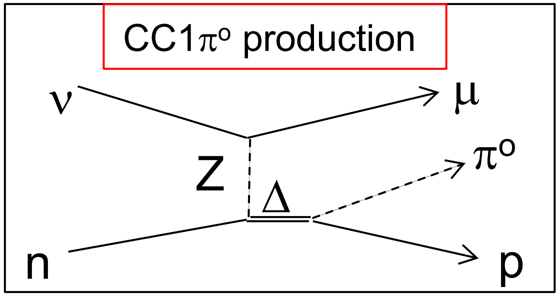
MiniBooNE,PRD83(2011)052009
Lalakulich et al,PRC87(2013)014602

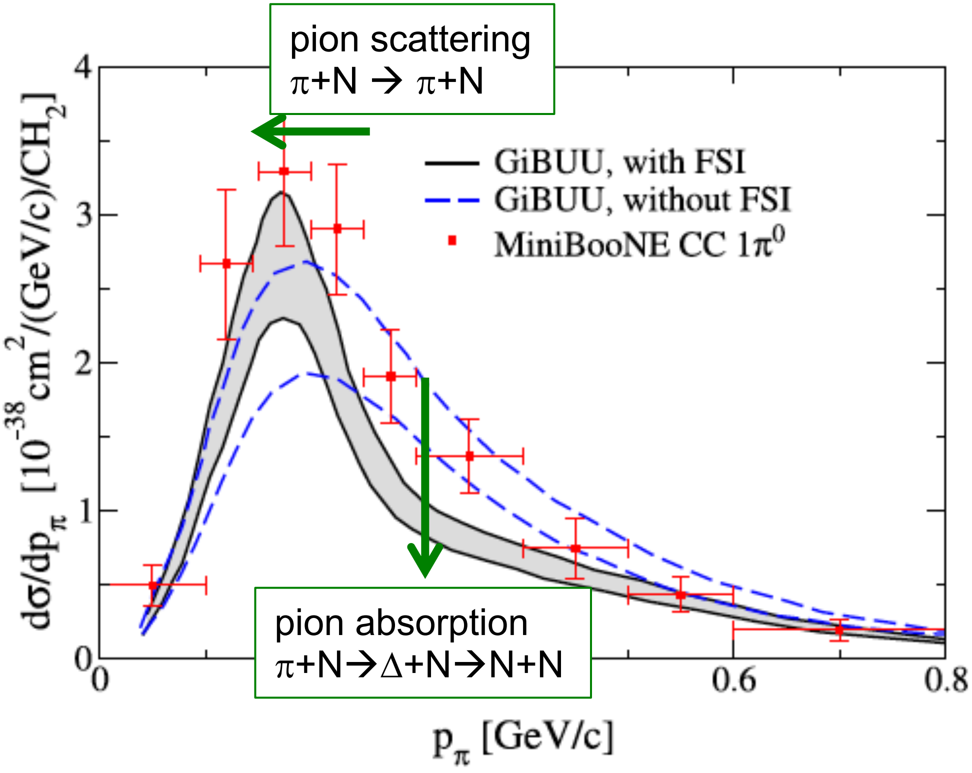


3. Neutrino Baryonic resonance data

Final state interaction

- Cascade model as a standard of the community
- Advanced models are not available for event-by-event simulation





pion scattering
$\pi+N \rightarrow \pi+N$

pion absorption
$\pi+N\rightarrow\Delta+N\rightarrow N+N$

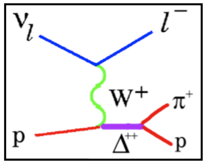
MiniBooNE πº momentum vs simulation

katori@fnal.gov

MiniBooNE,PRD83(2011)052009
Lalakulich et al,PRC87(2013)014602

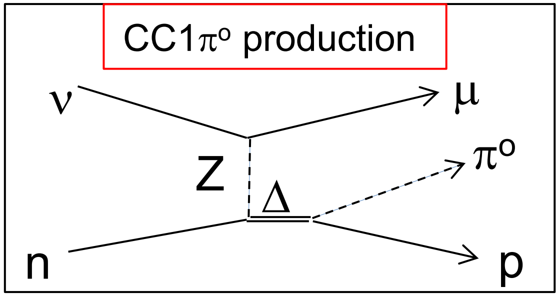
3. Neutrino Baryonic resonance data

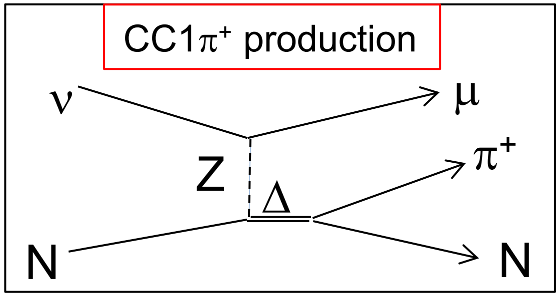
RESonance

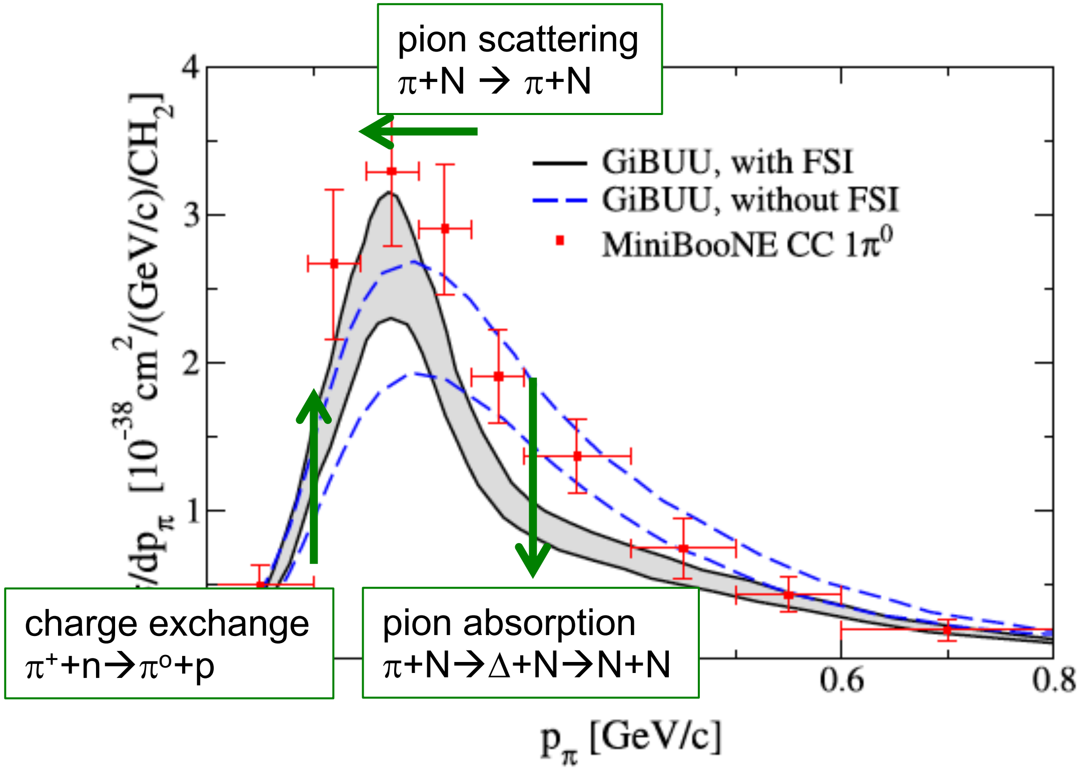


Final state interaction

- Cascade model as a standard of the community
- Advanced models are not available for event-by-event simulation


CC1π° production


CC1π⁺ production


pion scattering
$\pi+N \rightarrow \pi+N$

charge exchange
$\pi^++n \rightarrow \pi^o+p$

pion absorption
$\pi+N \rightarrow \Delta+N \rightarrow N+N$

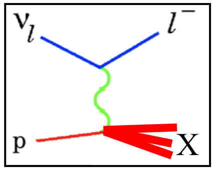
MiniBooNE π° momentum vs simulation

All neutrino baryonic resonance processes have ~30% errors

katori@fnal.gov

Nakamura et al,Rep.Prog.Phys.80(2017)056301

DIS

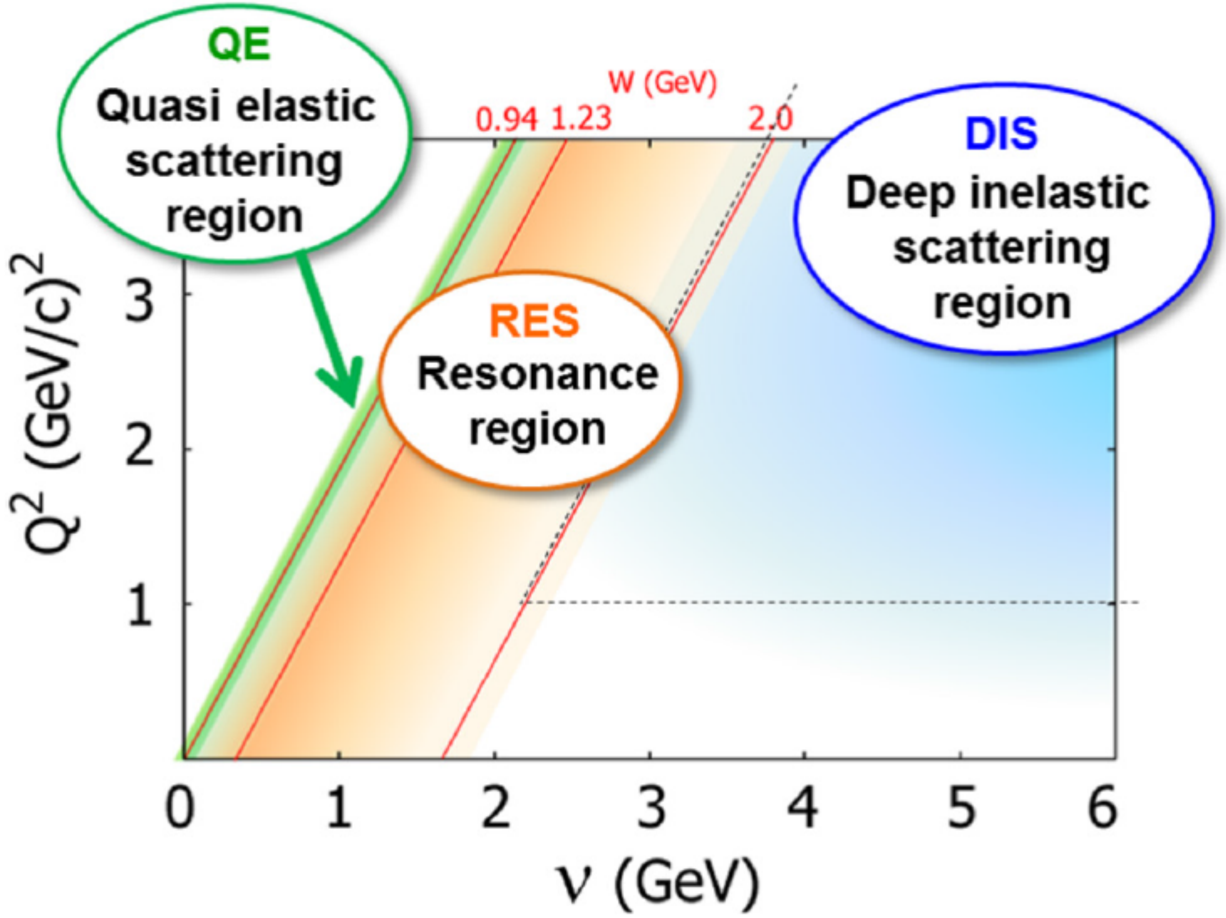


4. Shallow Inelastic Scattering (SIS)

Cross section

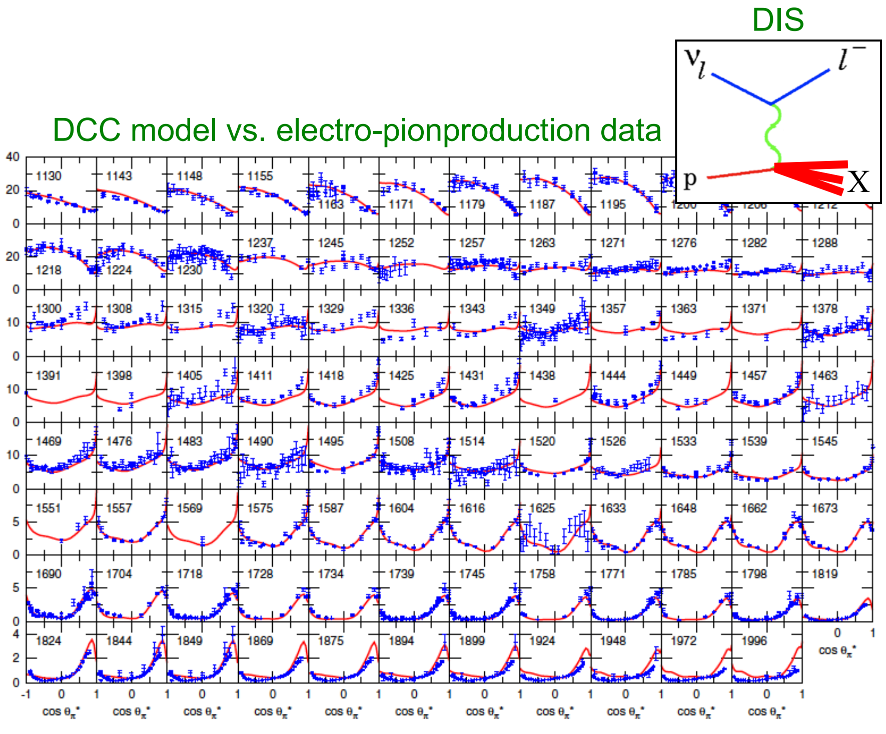
- Higher resonances and hadron dynamics
- low Q^2, low W DIS
- Nuclear dependent DIS

Neutrino experiment around 1-10 GeV is not quite DIS yet



Nakamura et al,Rep.Prog.Phys.80(2017)056301

4. Higher baryonic resonances



DCC model vs. electro-pionproduction data

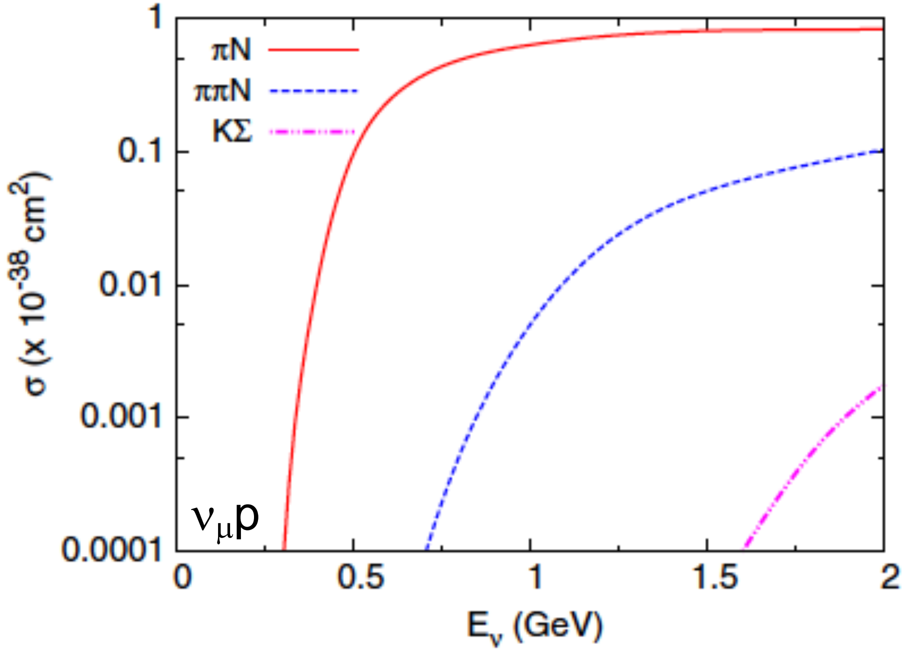
FIG. 8 (color online). Unpolarized differential cross sections, $d\sigma/d\Omega_\pi^*$ (μb/sr), for $\gamma n \to \pi^- p$. The data are from Refs. [55–78].

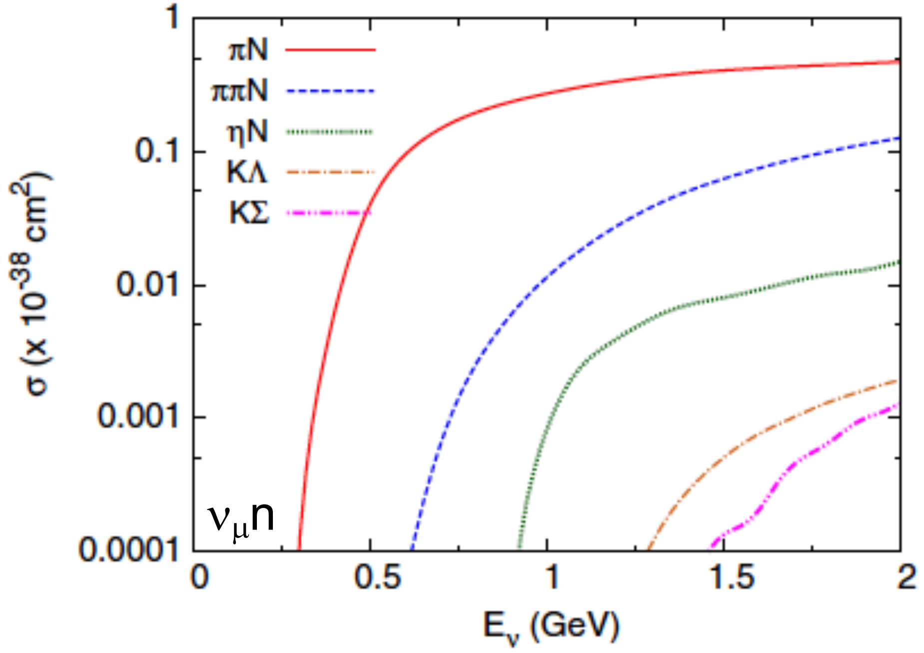
Cross section

- Higher resonances and hadron dynamics
- Quark-Hadron duality (low Q^2, low W DIS)
- Nuclear dependent DIS

DCC model

- Total amplitude is conserved
- Channels are coupled (πN, ππN, etc)
- 2 pion productions ~10% at 2 GeV



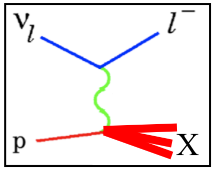


Bodek and Yang, AIP.Conf.Proc.670(2003)110,Nucl.Phys.B(Proc.Suppl.)139(2005)11

4. Quark-Hadron duality

Nachtmann variable $$\xi = \frac{2x}{\left(1+\sqrt{1+\frac{4x^2M^2}{Q^2}}\right)}$$

DIS



Cross section

- Higher resonances and hadron dynamics
- Quark-Hadron duality (low Q^2, low W DIS)
- Nuclear dependent DIS

GRV98 LO PDF + Bodek-Yang correction

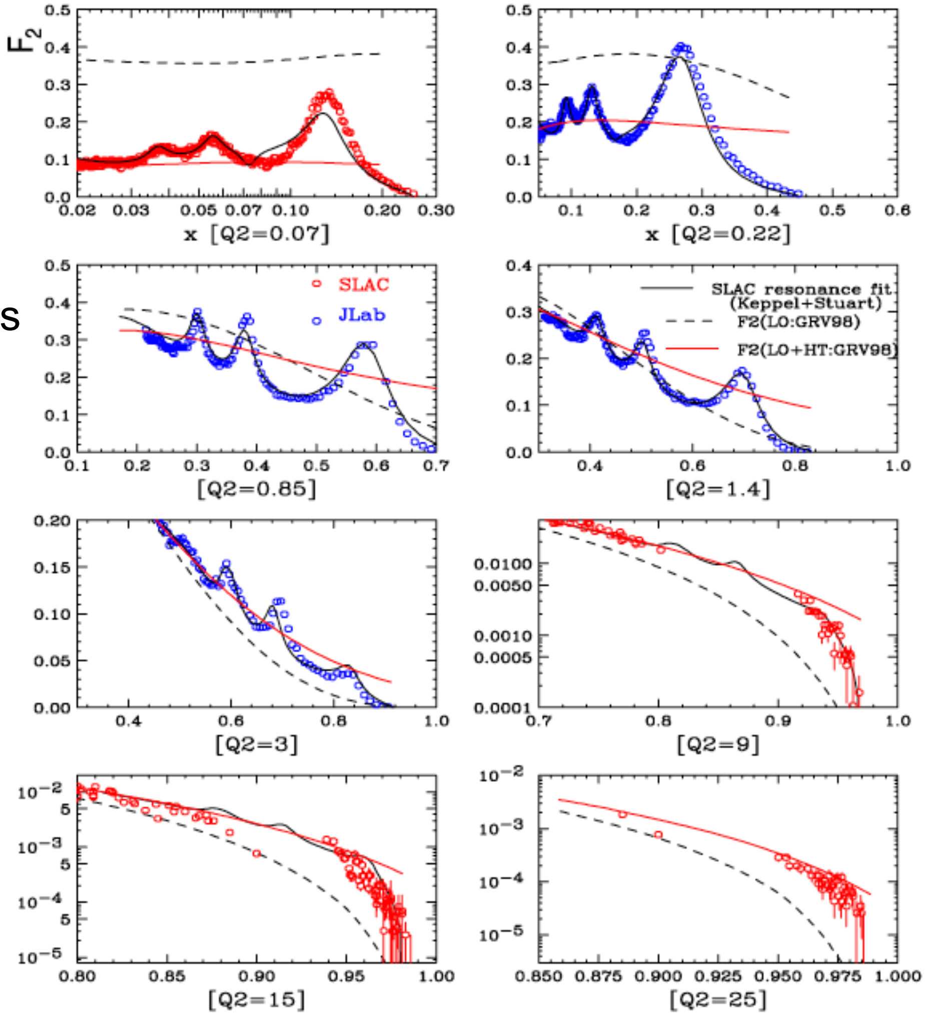
- GRV98 for low Q^2 DIS
- Bodek-Yang correction for QH-duality
- 20 years old, out-of-dated
- not sure how to implement systematic errors

$$\xi \rightarrow \xi_\omega = \frac{2x\left(1+\frac{M_f^2+B}{Q^2}\right)}{\left(1+\sqrt{1+\frac{4x^2M^2}{Q^2}}\right)+\frac{2Ax}{Q^2}}$$

$$K_{valence}(Q^2) = [1-G_D^2(Q^2)]\cdot\left(\frac{Q^2+C_{v2}}{Q^2+C_{v1}}\right)$$

$$K_{sea}(Q^2) = \frac{Q^2}{Q^2+C_{s1}}$$

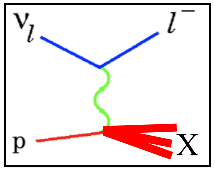
Proton F2 function GRV98-BY correction vs. data



HKN,PRC76(2007)065207, EPS,JHEP04(2009)065, FSSZ,PRD85(2012)074028
nCTEQ, PRD80(2009)094004

4. Nuclear dependent DIS

DIS



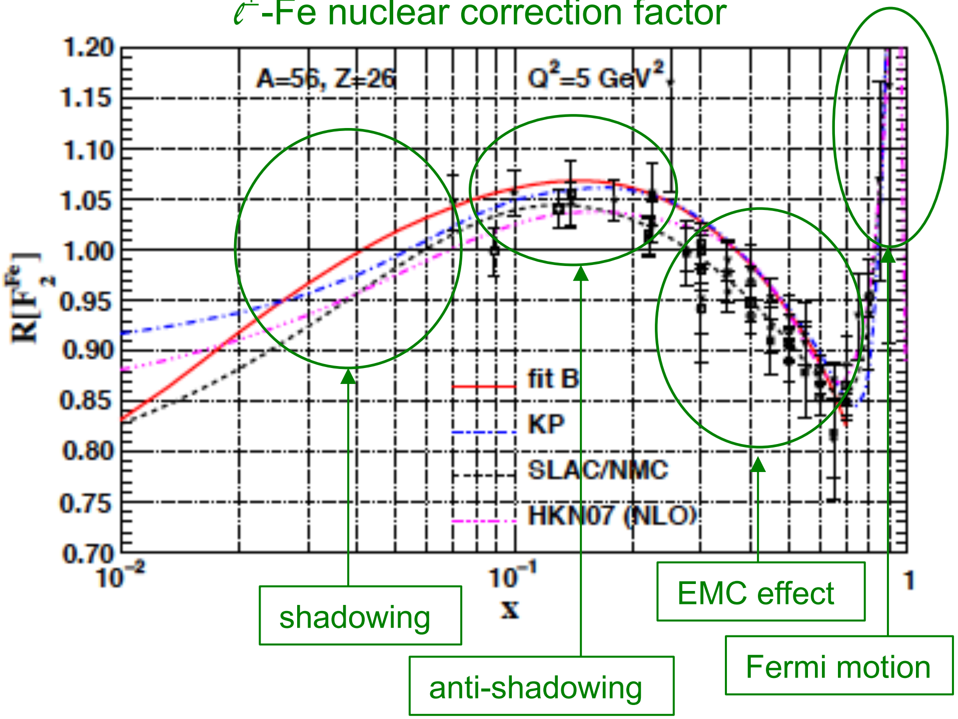
Cross section

- Higher resonances and hadron dynamics
- Quark-Hadron duality (low Q^2, low W DIS)
- Nuclear dependent DIS

Nuclear PDF

- Shadowing, EMC effect, Fermi motion
- Likely due to nucleon dynamics in nucleus
- Various models describe charged lepton data
- Neutrino data look very different

$\ell^\pm$-Fe nuclear correction factor



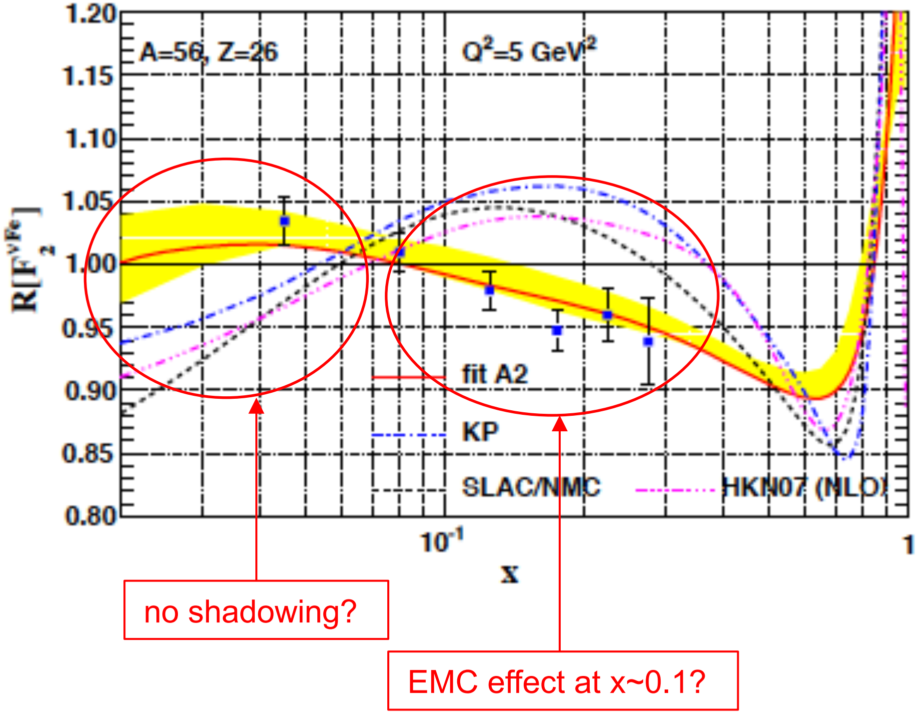
shadowing

anti-shadowing

EMC effect

Fermi motion

ν-Fe nuclear correction factor



no shadowing?

EMC effect at x~0.1?

Conclusion

ν-N scattering : important reactions for long baseline neutrino oscillation experiment
(T2K, NOvA, DUNE, Hyper-Kamiokande, etc)

CCQE: charged-current quasi-elastic, around 1 GeV
RES: baryonic resonance, around 2 GeV
DIS: deep inelastic scattering, 3 GeV to higher

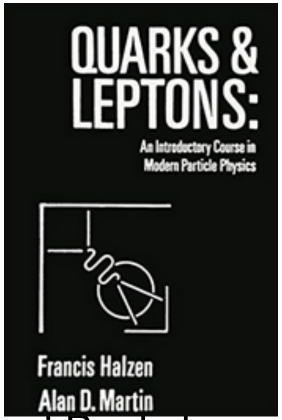
Nuclear physics sucks

- Fermi motion: nucleon motion smears kinematic reconstruction
- Pauli blocking: It limits low momentum transfer reaction
- Nuclear shell structure: separation energy (missing energy) for different nucleons
- Baryonic resonance: Often misidentified as CCQE
- Final state interaction: RES looks like CCQE, DIS looks like RES, etc
- Nucleon correlation: Physics between ν-N and ν-A interaction
- Quark-Hadron duality: Physics between ν-q and ν-N interaction
- Nuclear dependent PDF: Physics between ν-q and ν-A interaction

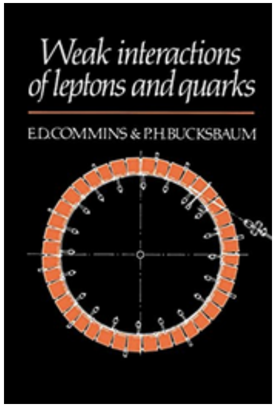
Currently, ~30% error is acceptable for many processes

References (books)

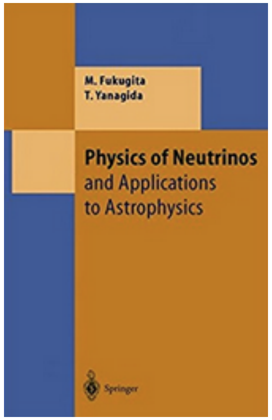
Quarks and Leptons (Halzen and Martin)
- show many calculations
- solutions for all exercises



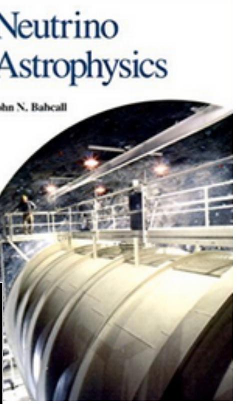
Weak interactions of Leptons and Quarks (Commins and Bucksbaum)
- show details of weak interaction calculations
- too many typos



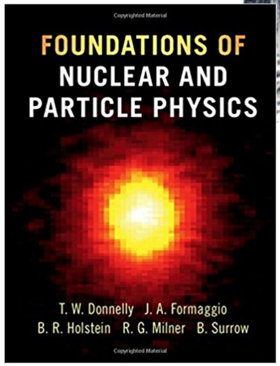
Physics of Neutrinos (Fukugita and Yanagida)
- very intense
- from solar neutrinos to SUSY



Neutrino astrophysics (Bahcall)
- good book to read



Foundation of Nuclear and Particle Physics (2017)
- Authors: Donnelly, Formaggio, Holstein, Milner, Surrow
- one and only one textbook on this subject
- buy if your PhD thesis topic is about neutrino cross section in T2K, NOvA, SBN, etc



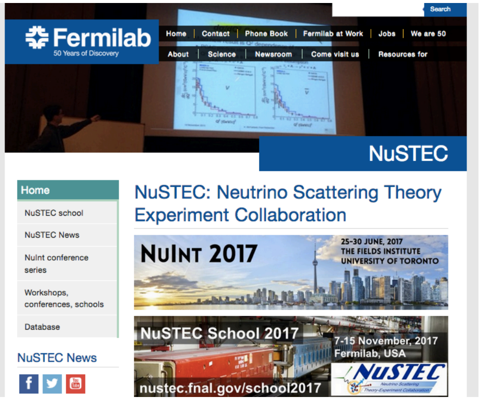
References (papers)

"From eV to EeV: Neutrino cross sections across energy scales"
- Authors: Formaggio and Zeller (MicroBooNE spokesperson)
- Rev.Mod.Phys.84(2012)1307, arXiv:1305.7513
- very good summary of neutrino cross sections

"Neutrino-Nucleus Cross Sections for Oscillation Experiments"
- Authors: Katori (me) and Martini (Martini model)
- J.Phys. G45 (2018) no.1, 013001
- A review both theoretical and experimental views

"NuSTEC White Paper: Status and challenges of neutrino–nucleus scattering"
- NuSTEC (Neutrino Scattering Theory-Experiment Collaboration)
- Prog.Part.Nucl.Phys. 100 (2018) 1-68
- Cover all open issues in the community

"NuSTEC News"
- http://nustec.fnal.gov/
- subscribe mailing list, "like" facebook page, use #nuxsec



NuSTEC: Neutrino Scattering Theory Experiment Collaboration

katori@fnal.gov